

What goes around comes around



What goes around comes around

Bulges of Galaxies: A Celebration of the 90th Birthday of Albert Whitford

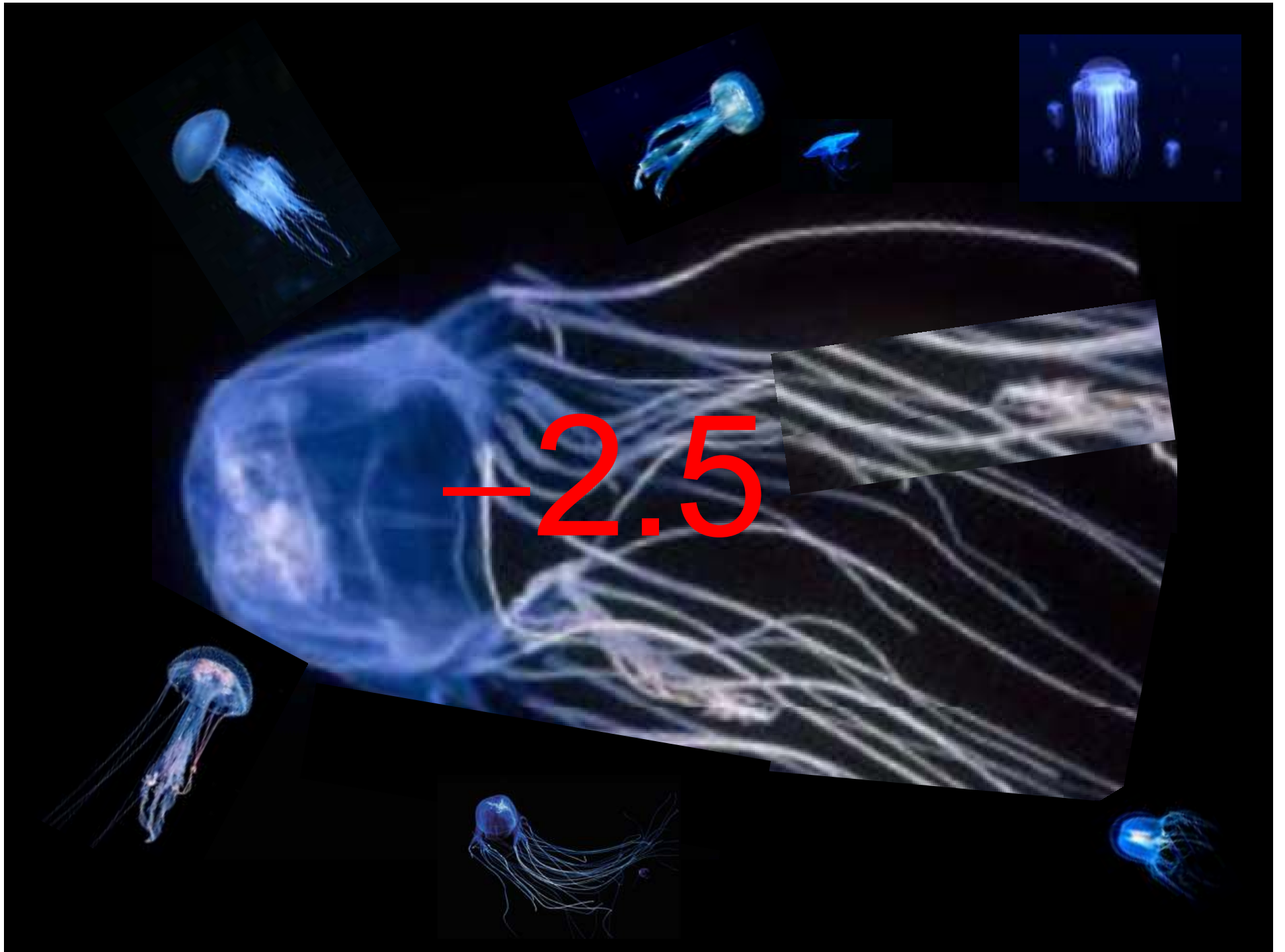
Rich & Terndrup 1997, PASP, 109, 571

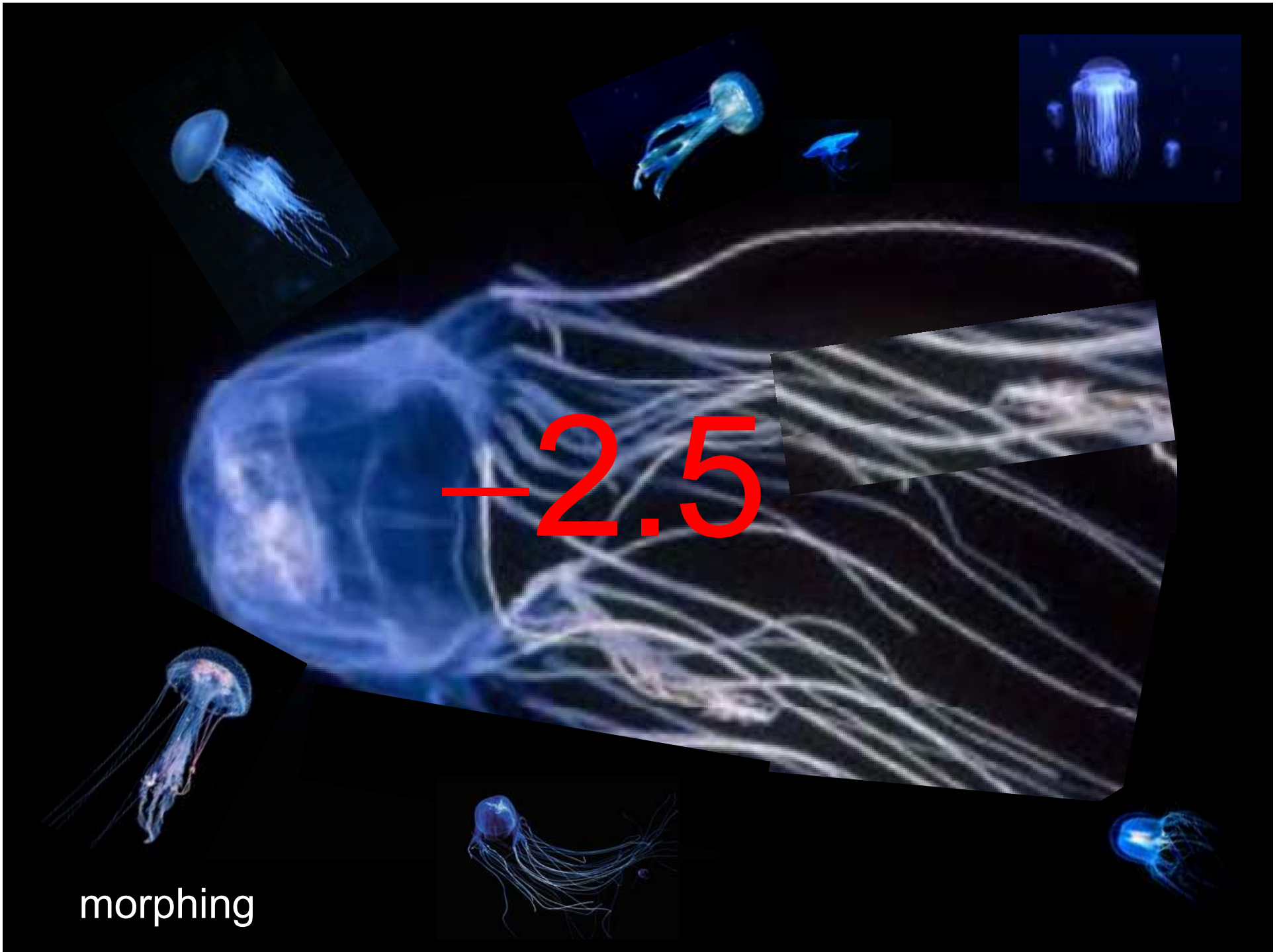
On the weekend of March 23–25, 1996, several dozen astronomers gathered under the redwoods at the University of California, Santa Cruz, to celebrate the 90th birthday of Albert Whitford, and to discuss the structure and stellar population of the bulge of our Galaxy and other related topics. Present were many of Whitford's lifelong friends and close associates, along with a number of researchers who more recently have been actively working on the central regions of the Galaxy.

The scientific organizing for the meeting consisted of Jay Frogel, Ruth Peterson, George Preston, Michael Rich, Nick Suntzeff, Donald Terndrup, and Bob Kraft (Chair). Local arrangements were supervised by Bob Kraft,

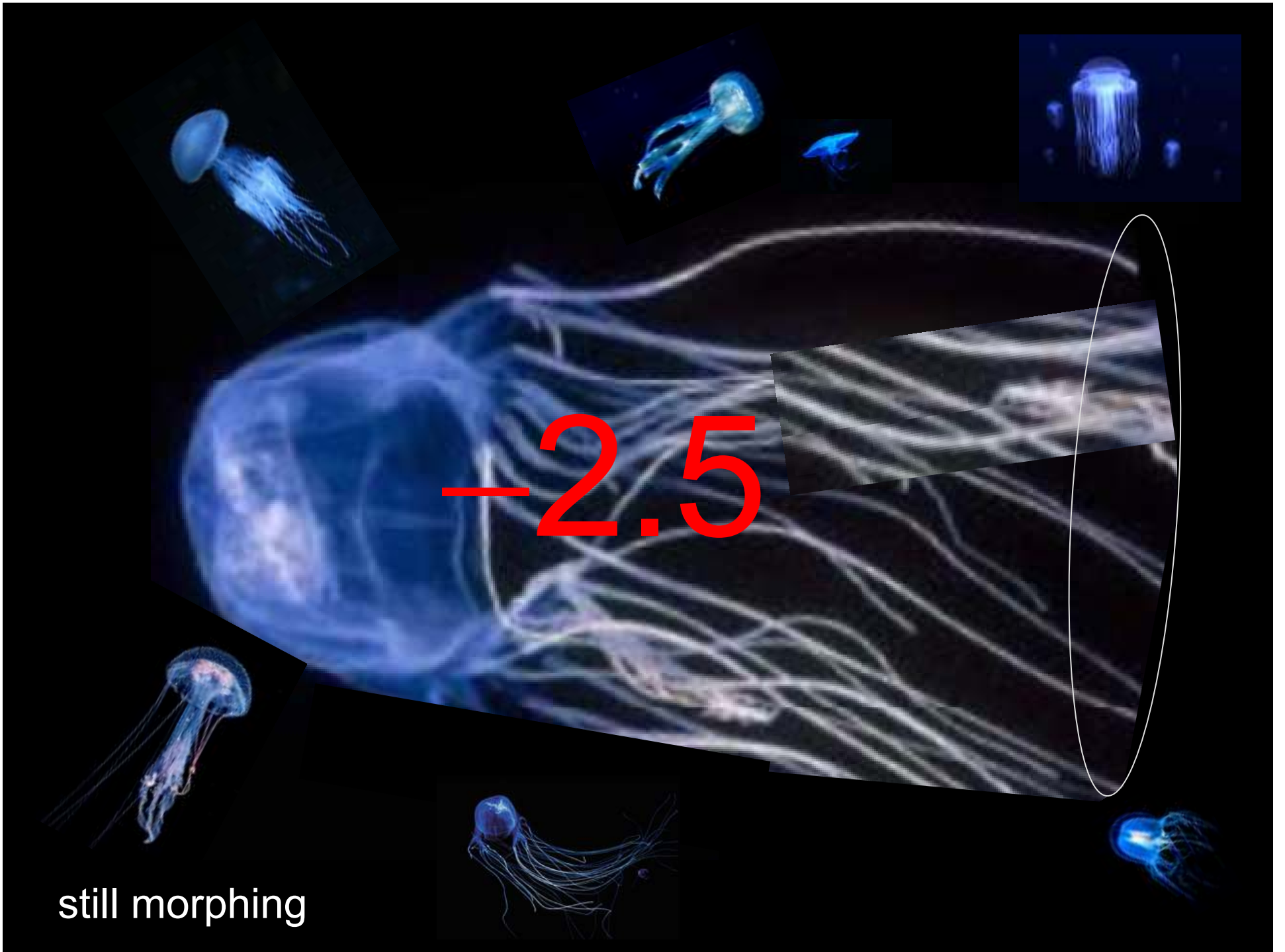
Chief cook & bottle washer

-2.5





morphing



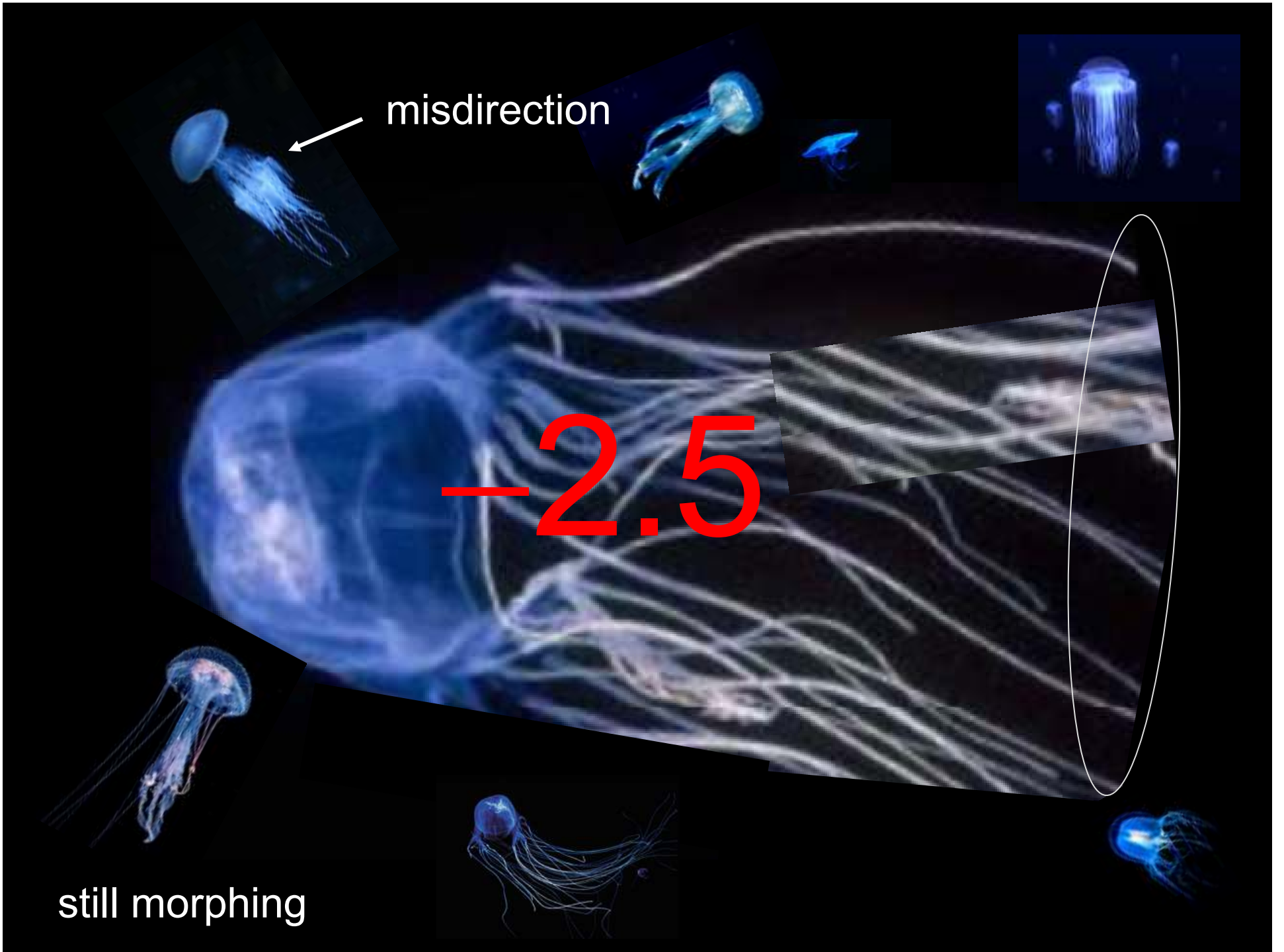
-2.5

still morphing

misdirection

-2.5

still morphing



You can't imagine where this is going

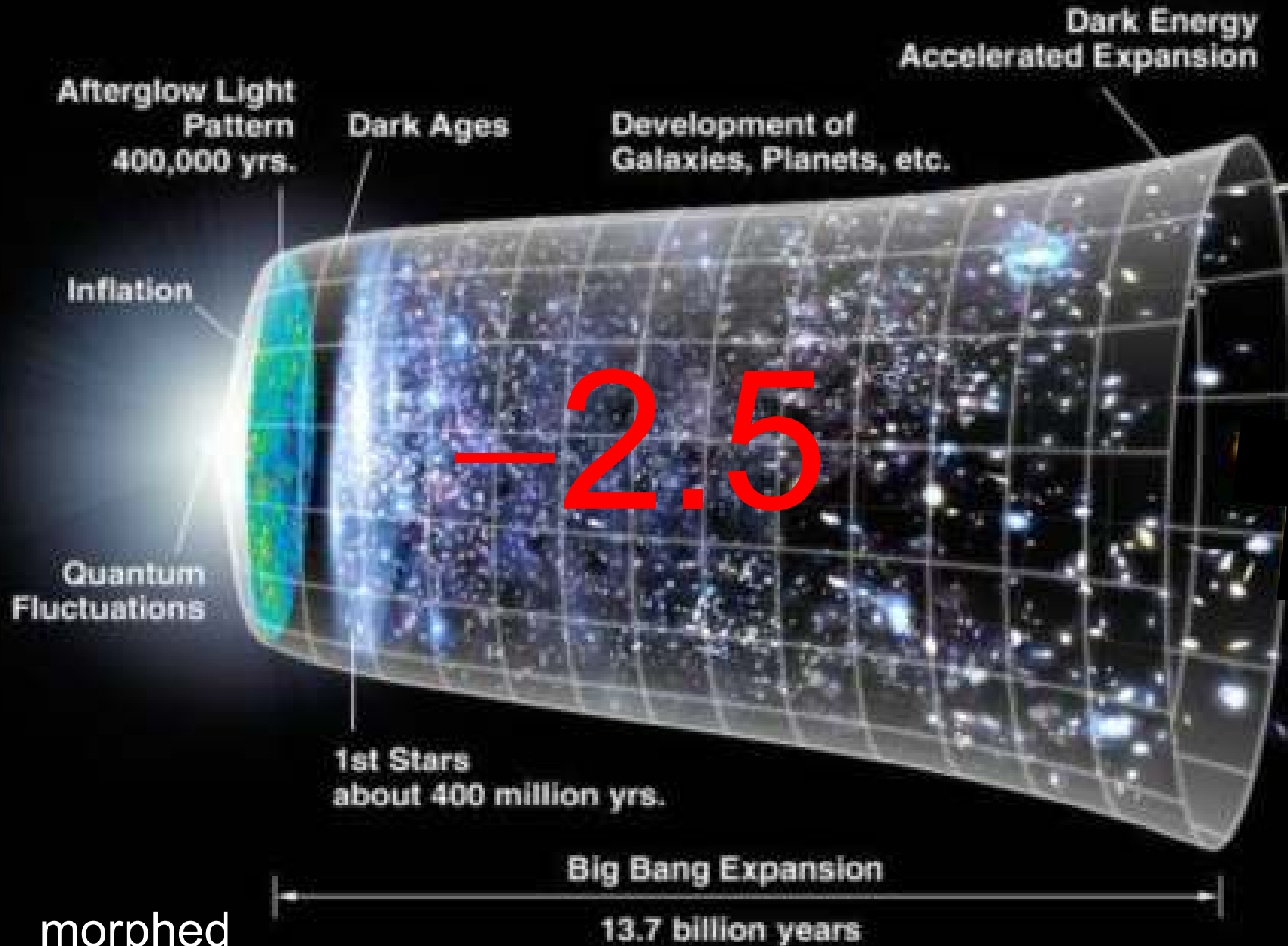


-2.5



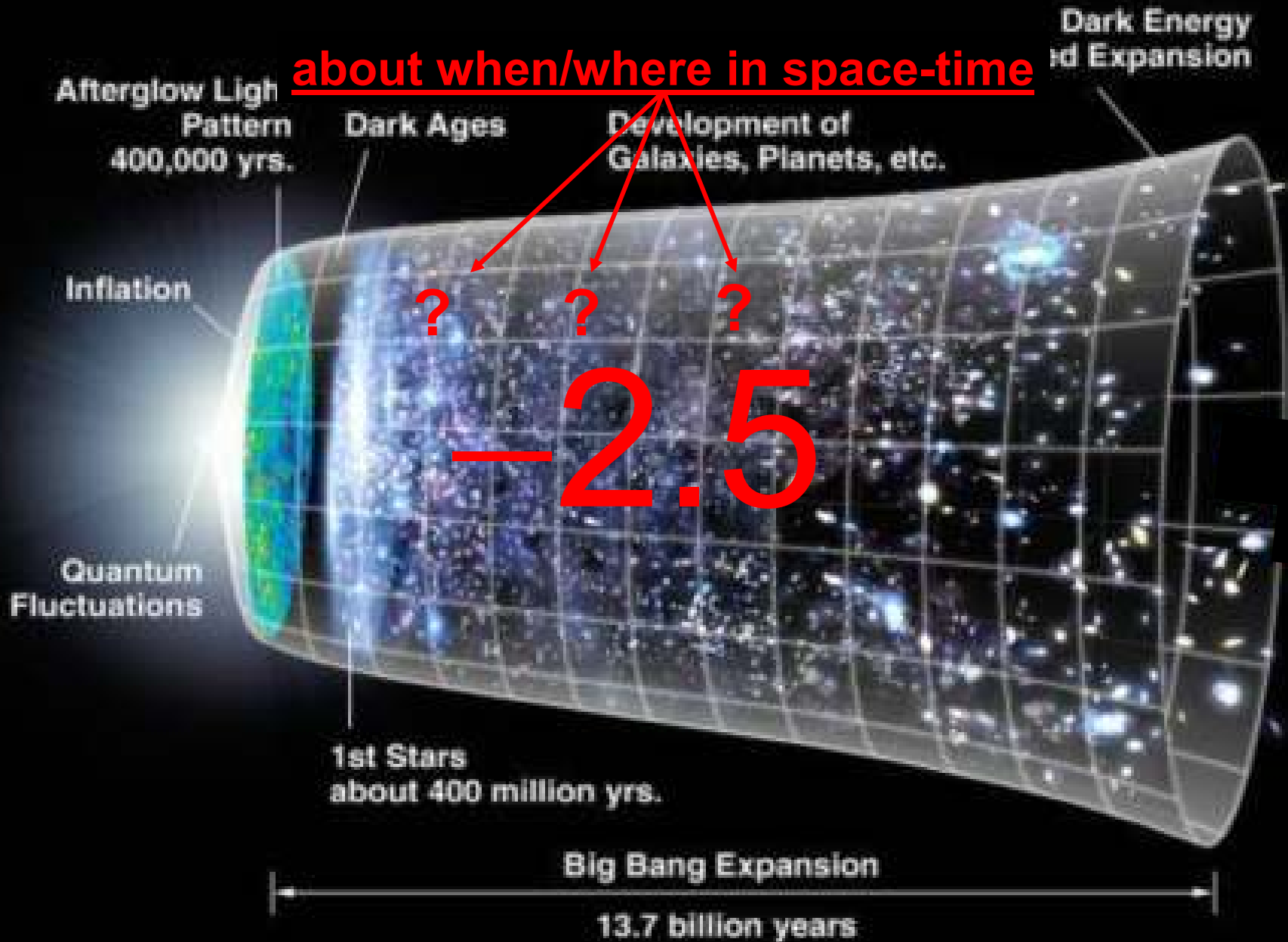
still morphing

-2.5 seems to be a special number in three related contexts

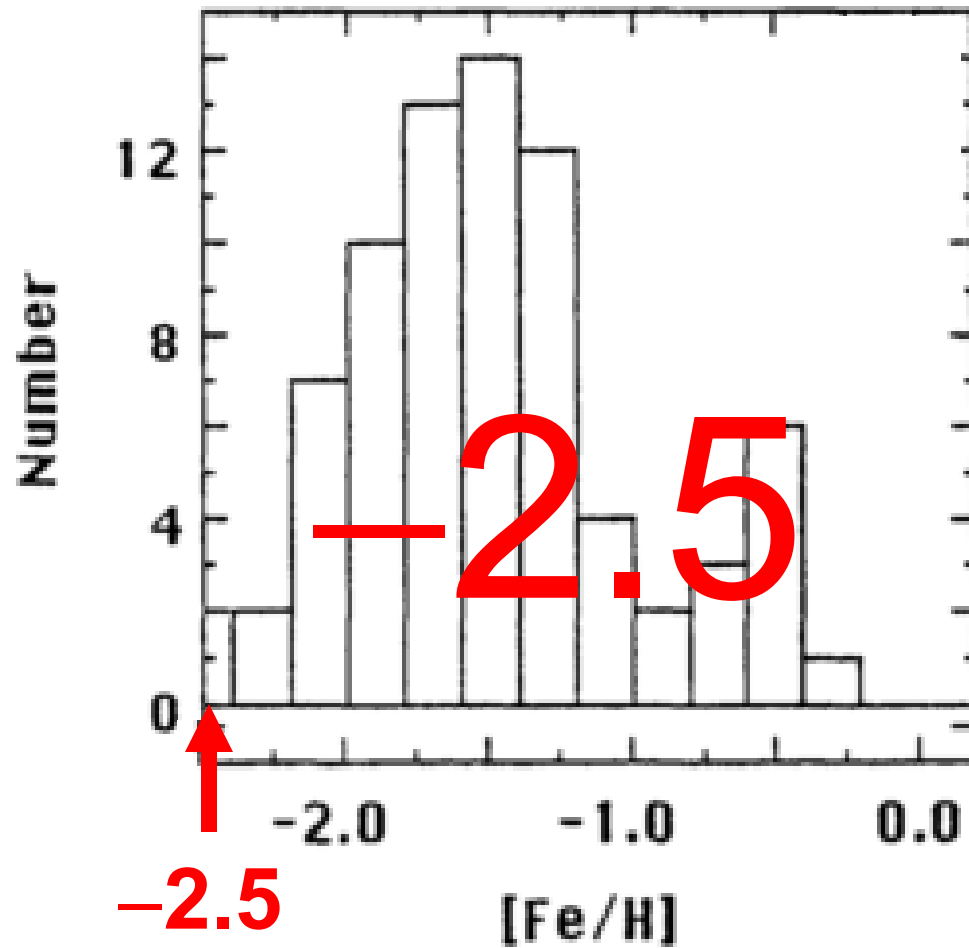


morphed

-2.5 seems to be a special number in three related contexts



(1) It is an apparent lower bound to the metallicity distribution of the Galactic Halo GCs (Zinn 1989 JRASC)



(1) It is an apparent lower bound to the metallicity distribution

A New Globular Cluster Metallicity Scale Based on the Abundance of Fe II

R. P. KRAFT¹ and I. I. IVANS²

*Carnegie Observatories Astrophysics Series, Vol. 4:
Origin and Evolution of the Elements, 2003*

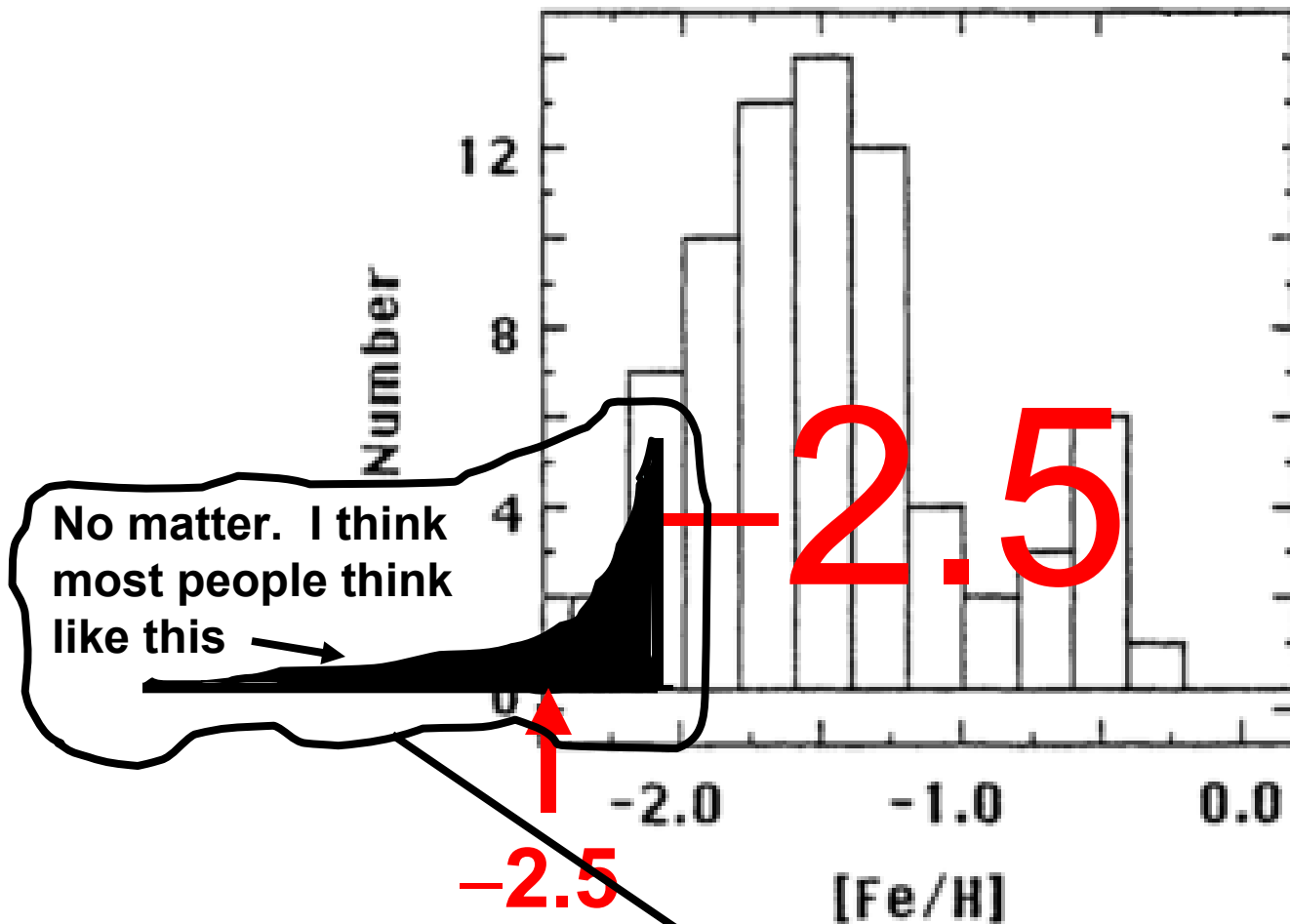
1.8 Conclusions

There exists no "definitive" set of cluster metallicities that are systematically reliable on the 0.05 dex level. Any discussion of cluster abundances (Galactic or extragalactic) must state clearly the underlying assumptions concerning the stellar atmosphere models used; the adopted T_{eff} -scale; what is meant by "metallicity" (eg. Fe I? Fe II? or a mean thereof?); what is the method used to derive $\log g$; and the origin of gf -values.

± 0.1 dex is the wiggle room,
according to Inese & Our Honoree

Is it a significant lower bound? .. Small number statistics,
yadda, yadda, yadda

(1) It is an apparent lower bound to the metallicity distribution of the Galactic Halo GCs (Zinn 1989 JRASC)



Is it a significant lower bound? .. Small number statistics, yadda, yadda, yadda

Time goes by

(2) It is also an apparent lower bound to abundances in damped Ly α systems at high redshift

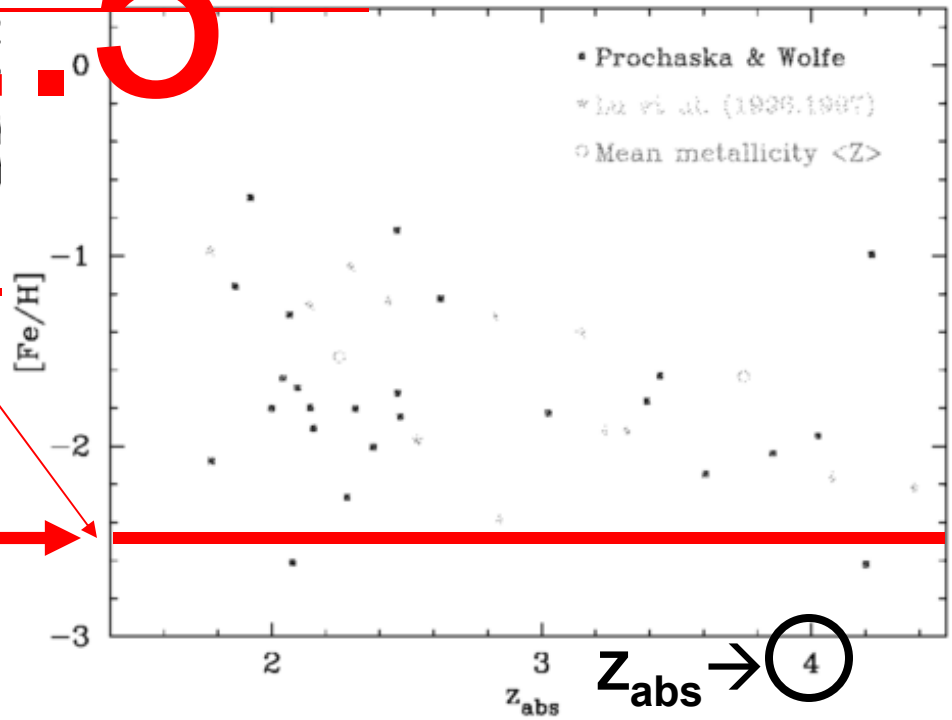
Received 1999 October 12; accepted 2000 February 28; published 2000 March 16

ABSTRACT

Observations of the damped Ly α systems provide direct measurements on the chemical enrichment history of neutral gas in the early universe. In this Letter, we present new measurements for four damped Ly α systems at high redshift. Combining these data with [Fe/H] values culled from the literature, we investigate the metallicity evolution of the universe from $z \approx 1.5$ to 4.5. Contrary to our expectations and the predictions of essentially every chemical evolution model, the $N(\text{H i})$ -weighted mean [Fe/H] metallicity exhibits minimal evolution over this epoch. For the individual systems, we report tentative evidence for an even lower mean and the scatter in [Fe/H], with the higher redshift systems showing lower values. We also note that no damped Ly α system has [Fe/H] < -2.7 dex. Finally, of small number statistics and dust on our conclusions and consider the implications for metallicity evolution in the early universe.

Significance of this lower bound? ..
Again, small number statistics,
more yadda, yadda, yadda

-2.5



METALLICITY EVOLUTION IN THE EARLY UNIVERSE

Jason X. Prochaska¹ * (and Art Wolfe)

* Jason still owes me one US dollar re: "24 steps"

(2) It is also an apparent lower bound to abundances in damped Ly α systems at high redshift

Received 1999 October 12; accepted 2000 February 28; published 2000 March 16

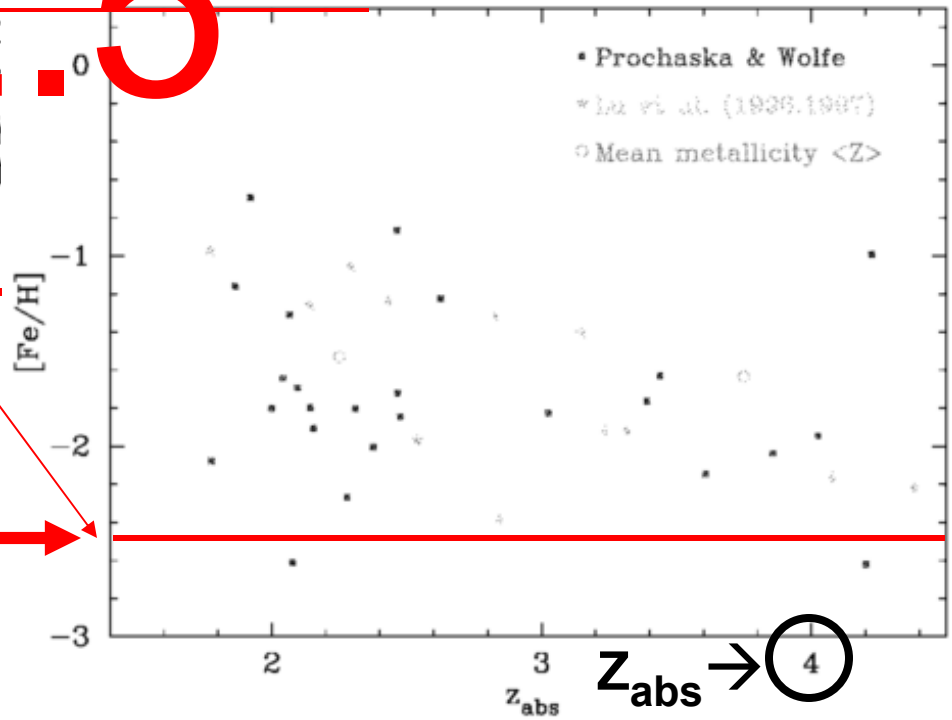
ABSTRACT

Observations of the damped Ly α systems provide direct measurements on the chemical enrichment history of neutral gas in the early universe. In this Letter, we present new measurements for four damped Ly α systems at high redshift. Combining these data with [Fe/H] values culled from the literature, we investigate the metallicity evolution of the universe from $z \approx 1.5$ to 4.5. Contrary to our expectations and the predictions of essentially every chemical evolution model, the $N(\text{H i})$ -weighted mean [Fe/H] metallicity exhibits minimal evolution over this epoch. For the individual systems, we report tentative evidence for a decrease in the mean and the scatter in [Fe/H], with the higher redshift systems showing lower values. We also note that no damped Ly α system has [Fe/H] < -2.7 dex. Finally, because of small number statistics and dust on our conclusions and consider the implications for metallicity evolution in the early universe.

2.5

Significance of this lower bound? .. Again, small number statistics, more yadda, yadda, yadda

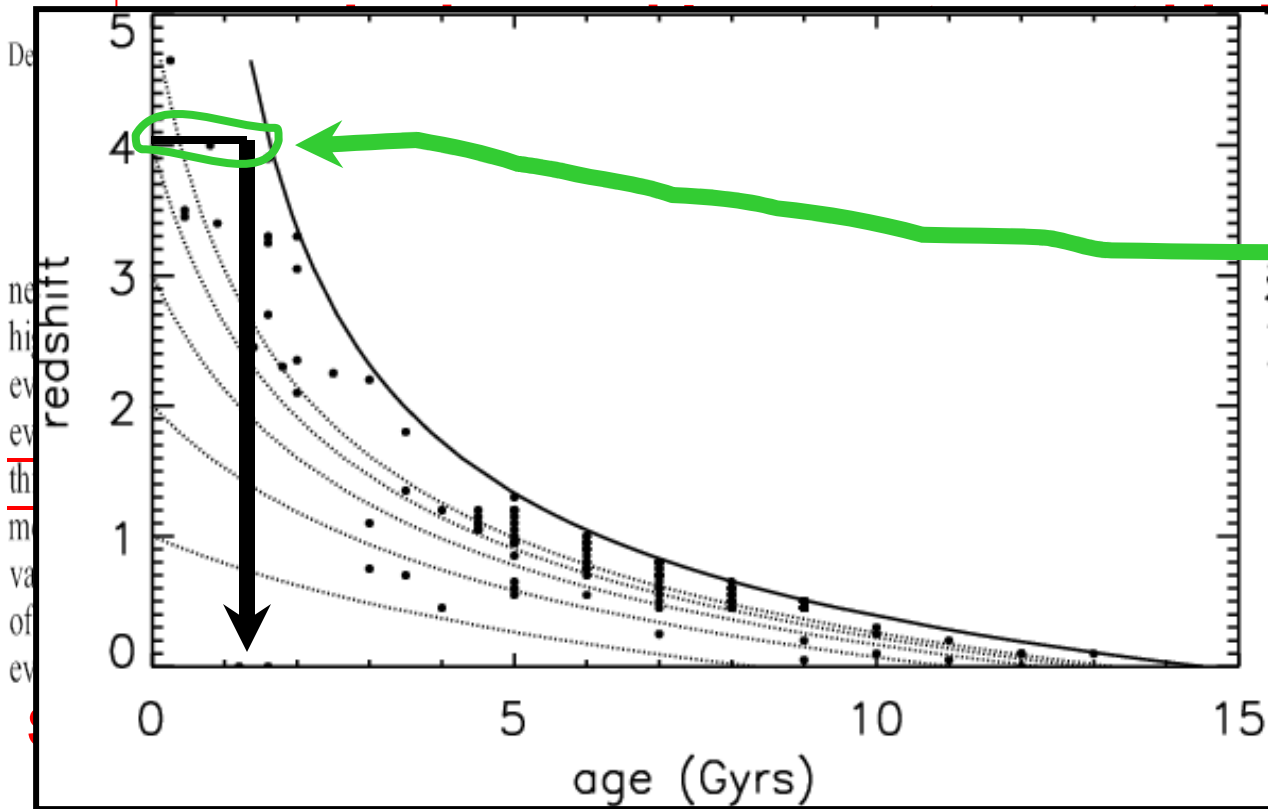
-2.5



METALLICITY EVOLUTION IN THE EARLY UNIVERSE

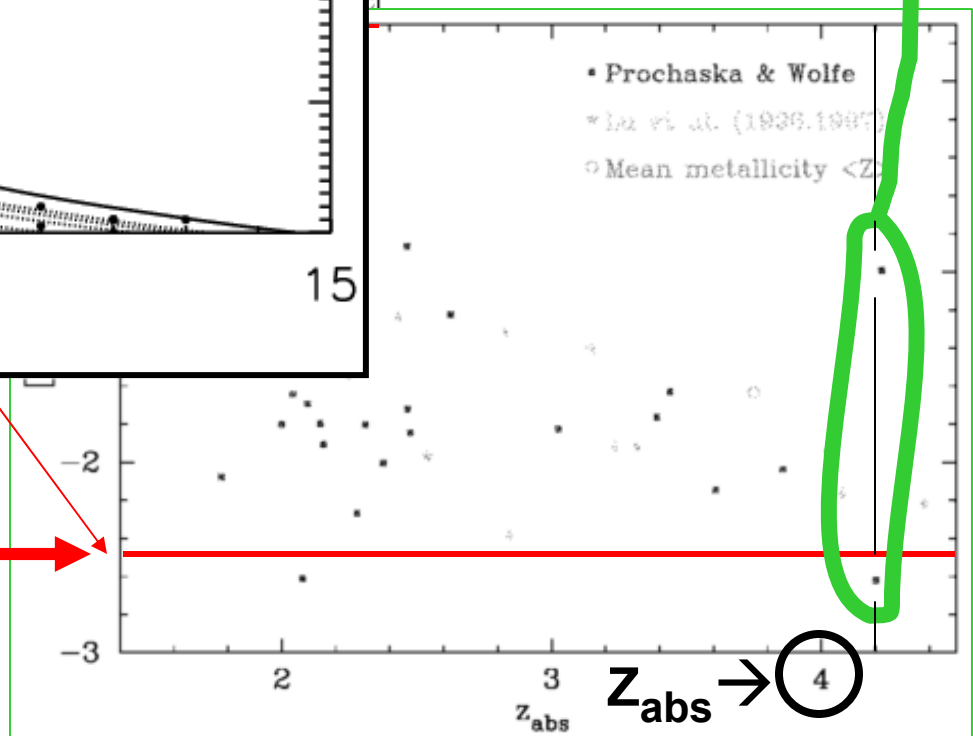
Jason X. Prochaska¹ (and Art Wolfe)

(2) It is also an apparent lower bound to abundances in redshift



more yadda, yadda, yadda

-2.5

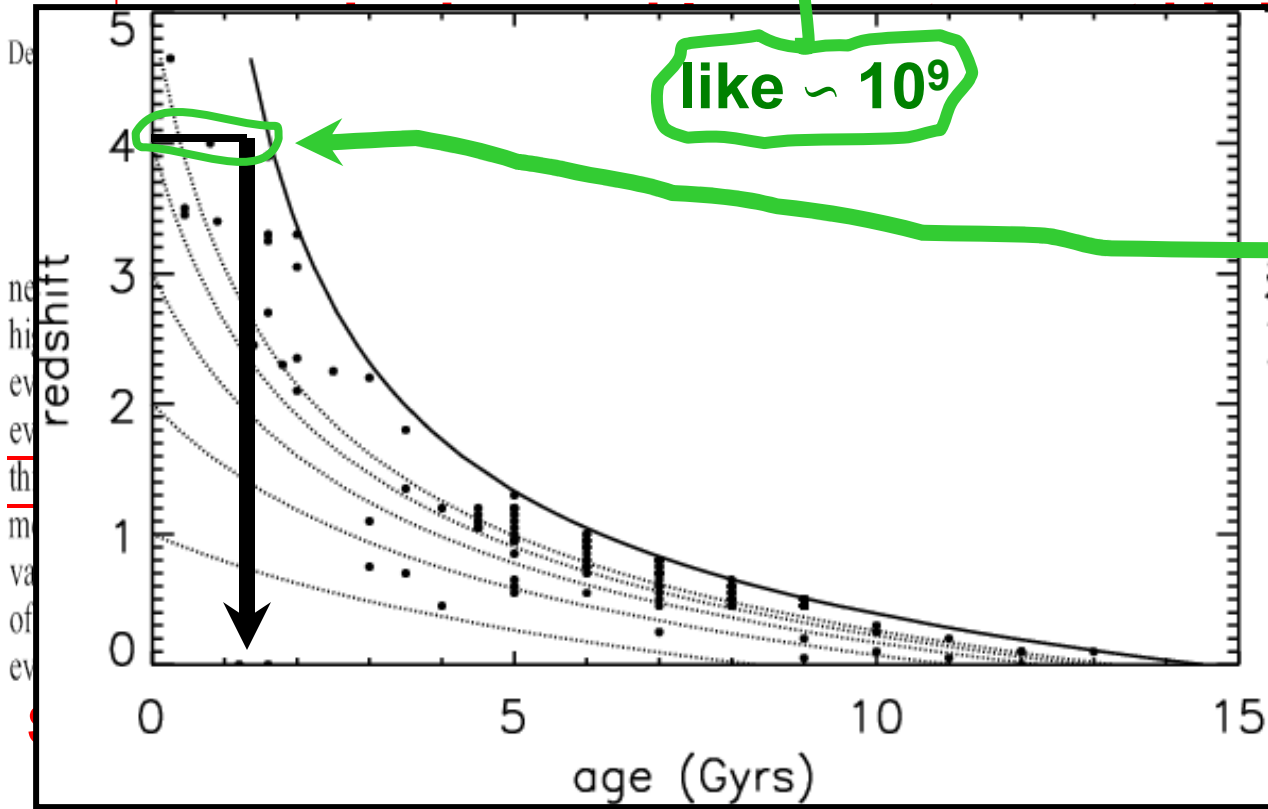


METALLICITY EVOLUTION IN THE EARLY UNIVERSE

Jason X. Prochaska¹ (and Art Wolfe)

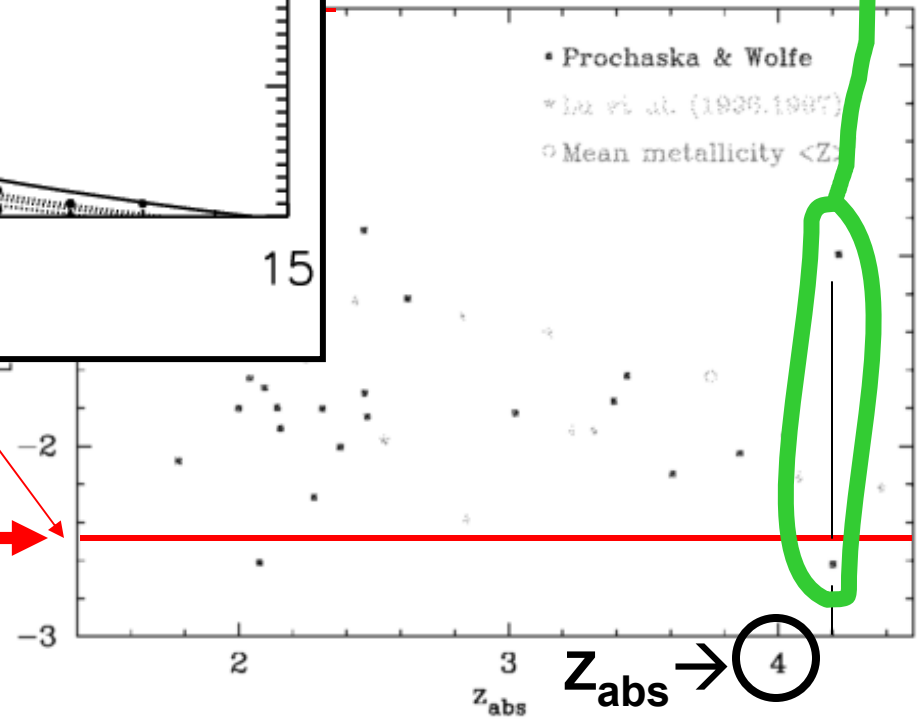
(2) It is also an apparent lower bound to abundances in redshift

like $\sim 10^9$



more yadda, yadda, yadda

-2.5

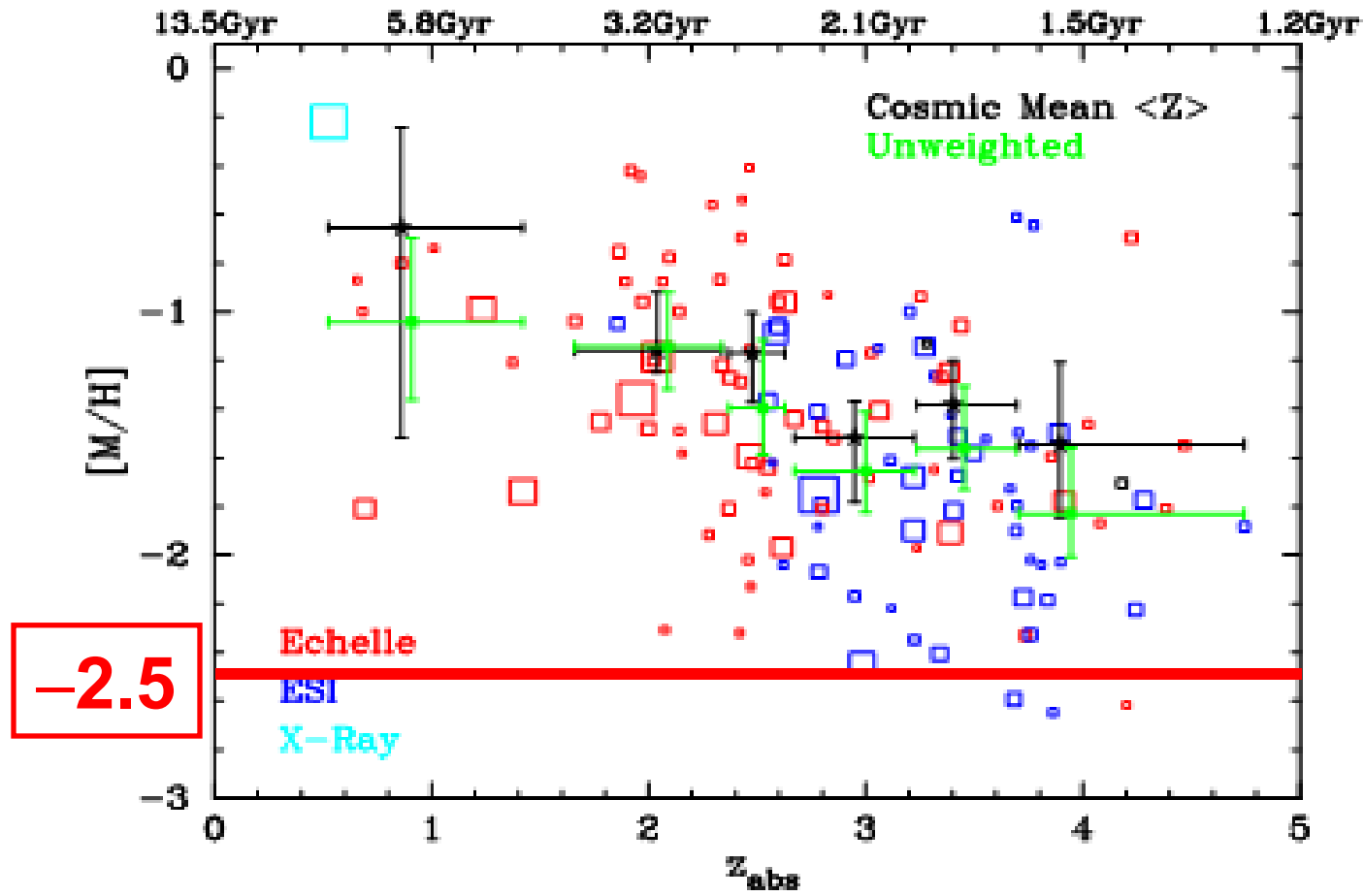


Here is a Prochaska 2004 update

Chemical Abundances in the Damped Ly α Systems

JASON X. PROCHASKA
UCO/Lick Observatories

*Carnegie Observatories Astrophysics Series, Vol. 4:
Origin and Evolution of the Elements*



http://pnm.itp.ucsb.edu/online/stars_c07/wolfe/pdf/Wolfe_StarFormationConf_KITP.pdf

KITP Conference: Star Formation, Then and Now: Friday, August 17, 2007

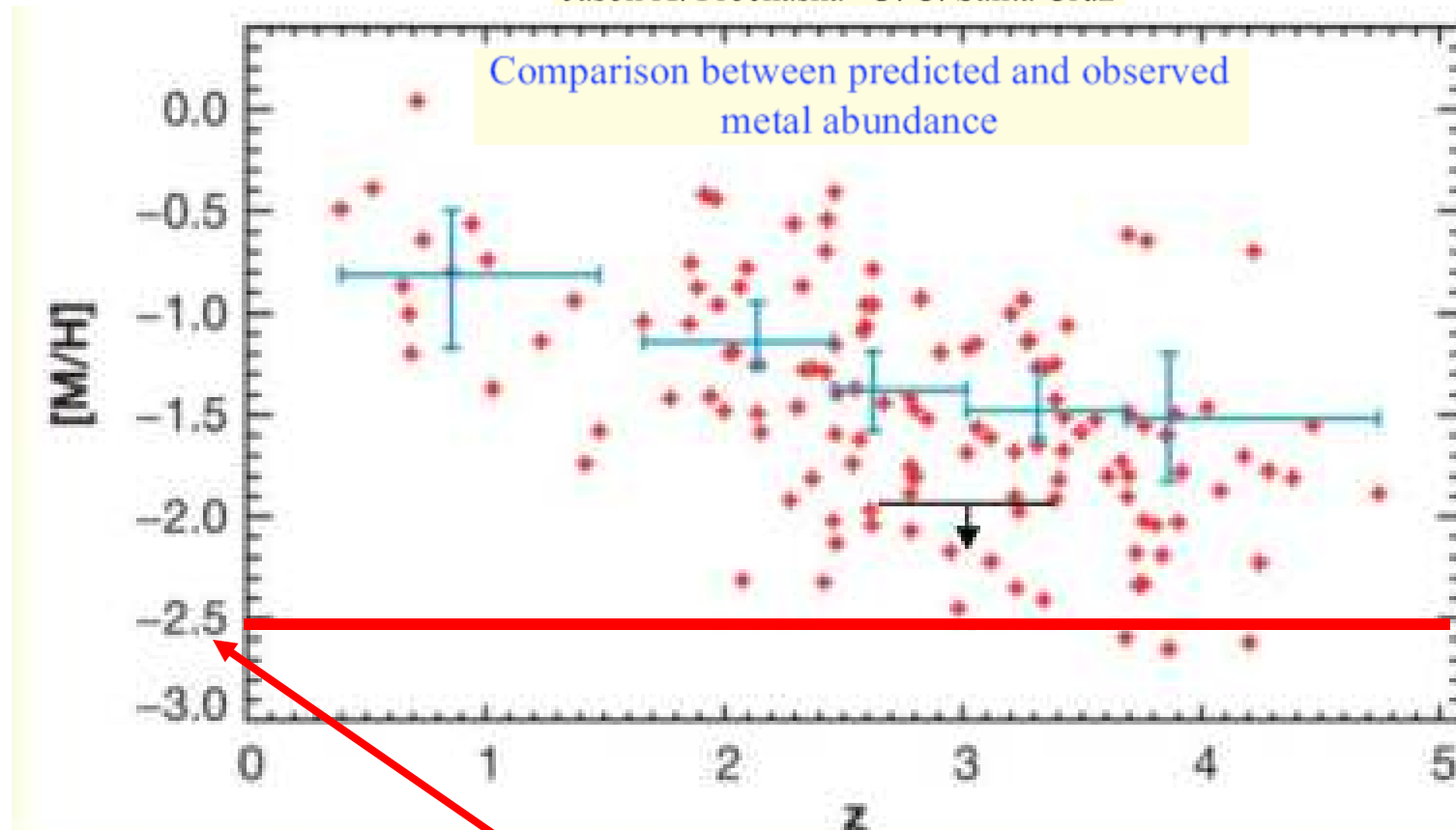
Bimodality in Damped Ly α Systems

Art Wolfe

Hsiao-Wen Chen--U. Chicago

Regina Jorgenson--U. C. San Diego

Jason X. Prochaska--U. C. Santa Cruz



And Wolfe et al presented similar data at KITP in 2007, so it looks like the J&P 2000 conclusion still stands.

(3a) Finally, It is the location of abrupt slope-changes in chemical evolution diagrams for stars of the Galactic halo.

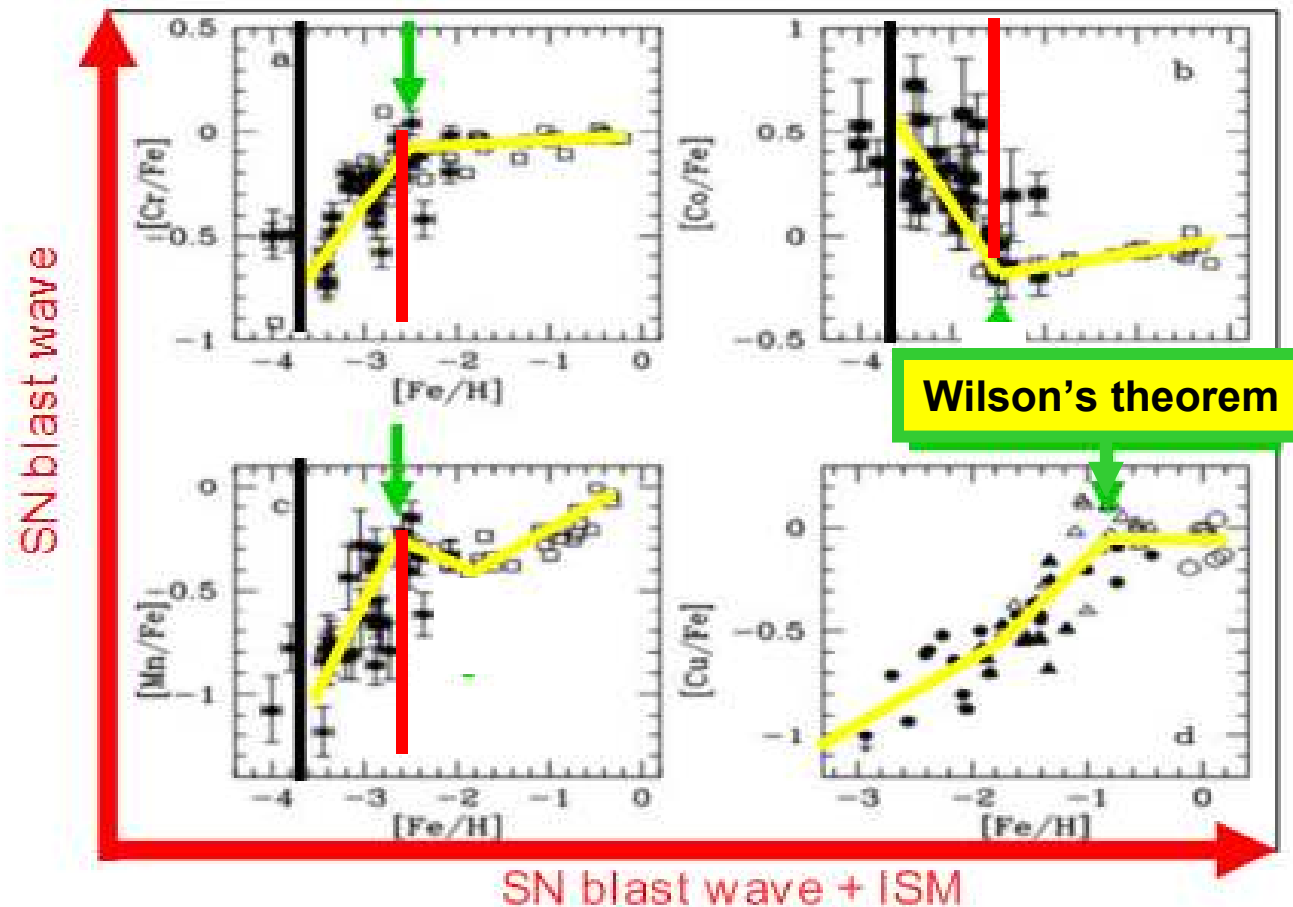
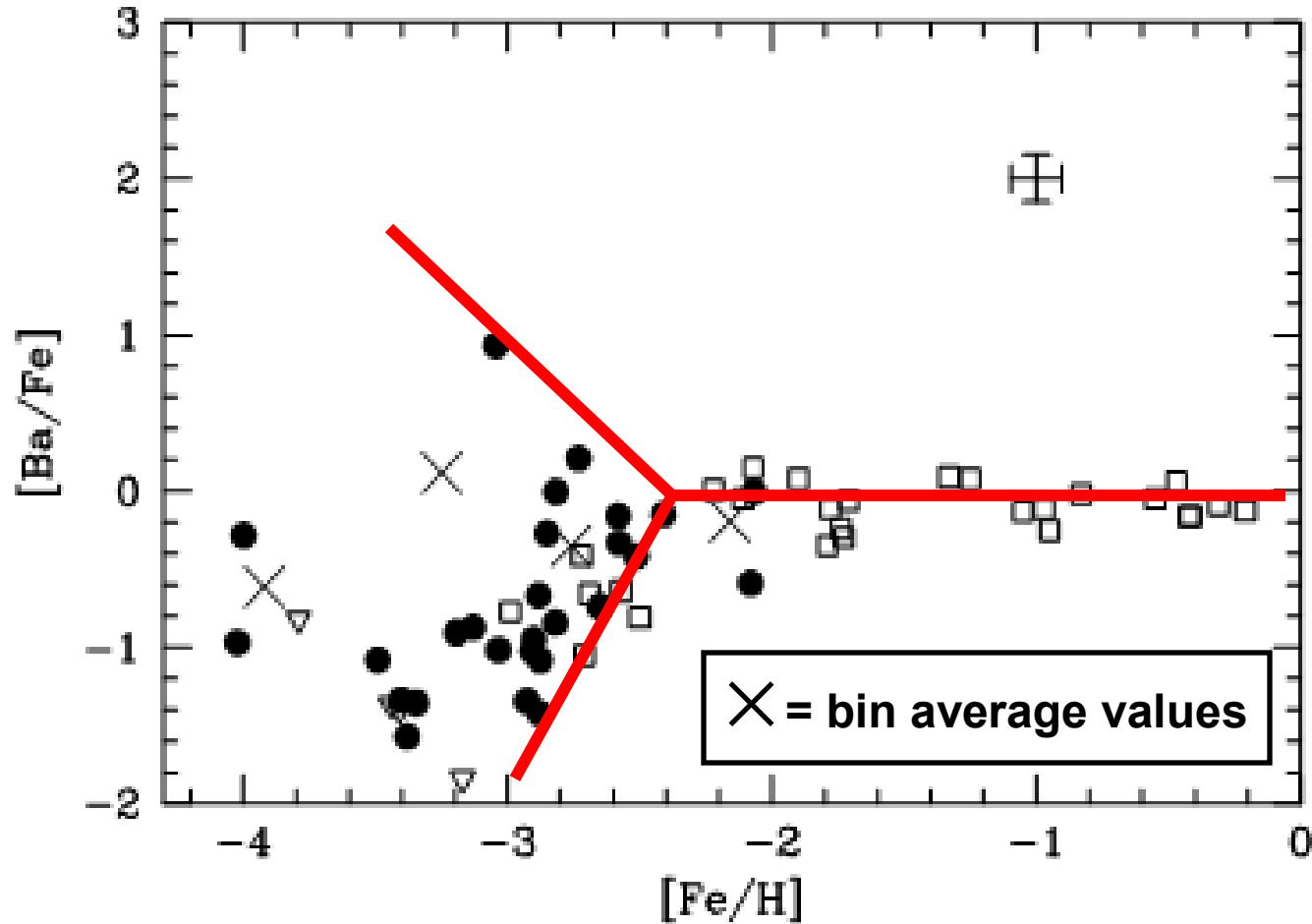


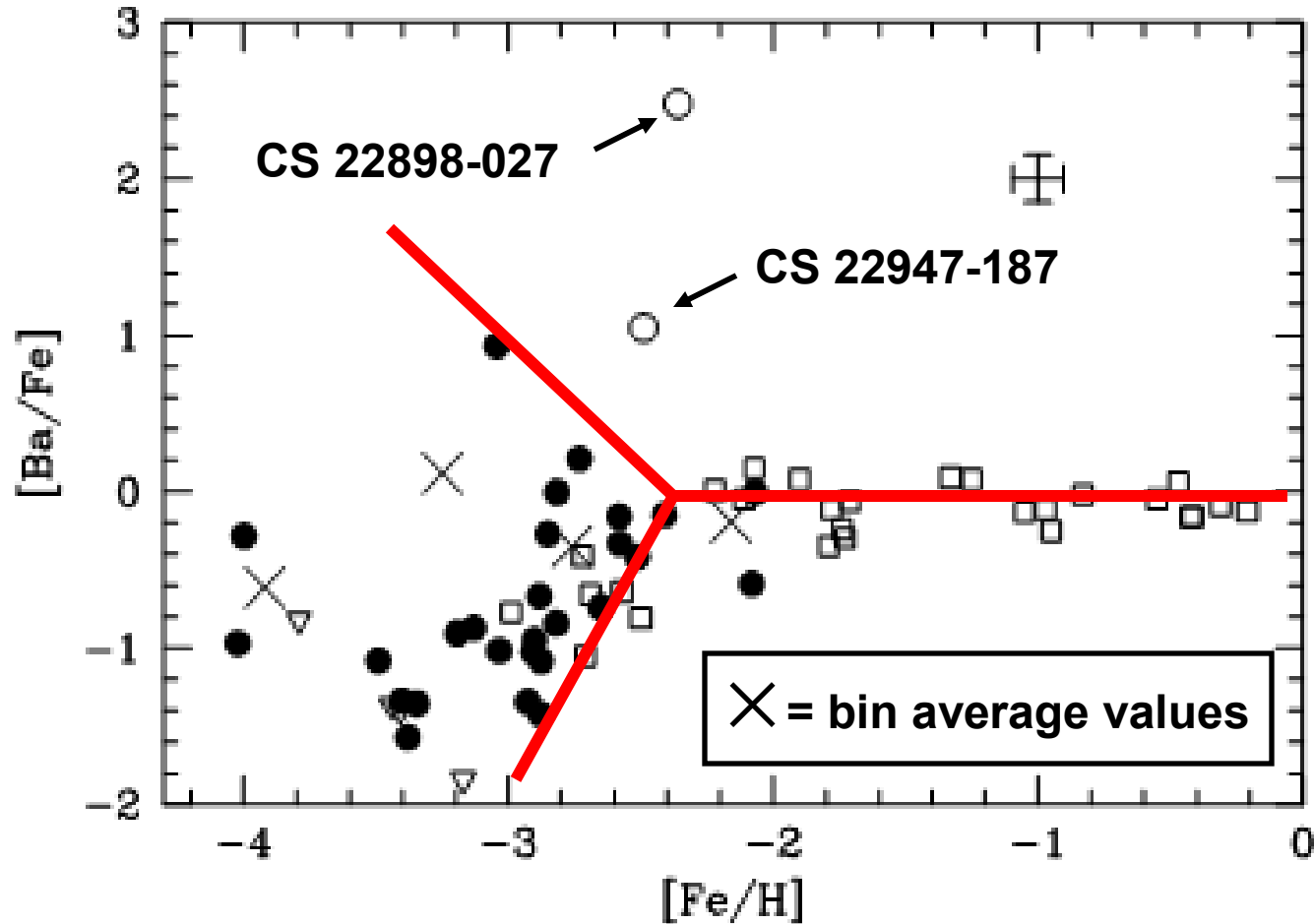
Fig. 1 Logarithmic element to Fe ratios n units of the solar values are plotted versus logarithmic Fe/H ratios, also in solar units for four Fe-peak elements (from McWilliam et al 1995, AJ 109, 2757). Yellow line-segments are eye-fitted empirical regressions. Green arrows mark changes in slope. Black vertical lines near $[Fe/H] = -3.6$ denote a speculative limit below which $[Fe/H]$ is no longer a surrogate clock (see Auldouze & Silk 1995, ApJ, 451 L49). The bold, red coordinate axes are included to dramatize the problem of explaining why element-to-Fe ratios are correlated with $[Fe/H]$.

(3b) And, approximately, it is where the Barium-to-Iron abundance ratio converges to its present-day solar value.



ANDREW MC WILLIAM THE ASTRONOMICAL JOURNAL, 115:1640-1647, 1998.

(3c) Pollution by AGB binary companions can produce occasional carbon stars with even larger Barium excesses



ANDREW MC WILLIAM THE ASTRONOMICAL JOURNAL, 115:1640-1647, 1998.

Accepting
Prompt Initial Enrichment (PIE)
by super-massive stars (SMS) during first 10^9 years

maybe

changes in nature of Chemical Evolution at $[\text{Fe}/\text{H}] = -2.5$
are due to the transition
SMS nucleosynthesis \rightarrow “ordinary” SN nucleosynthesis.

Cloud collapse \rightarrow SMS (or TH-IMF) \rightarrow PIE to $[\text{Fe}/\text{H}] \sim -3$

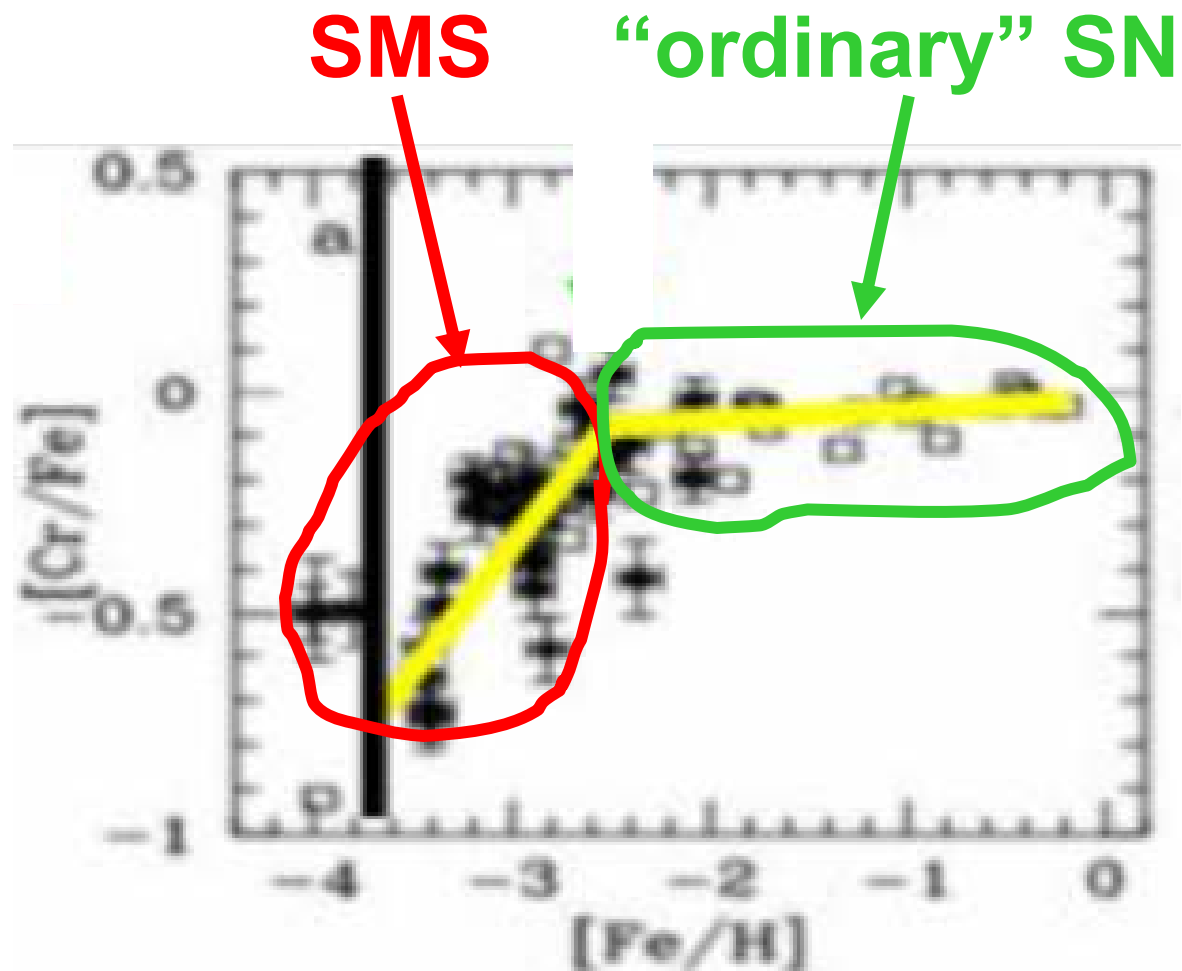


Metal enrichment alters mode of cloud collapse



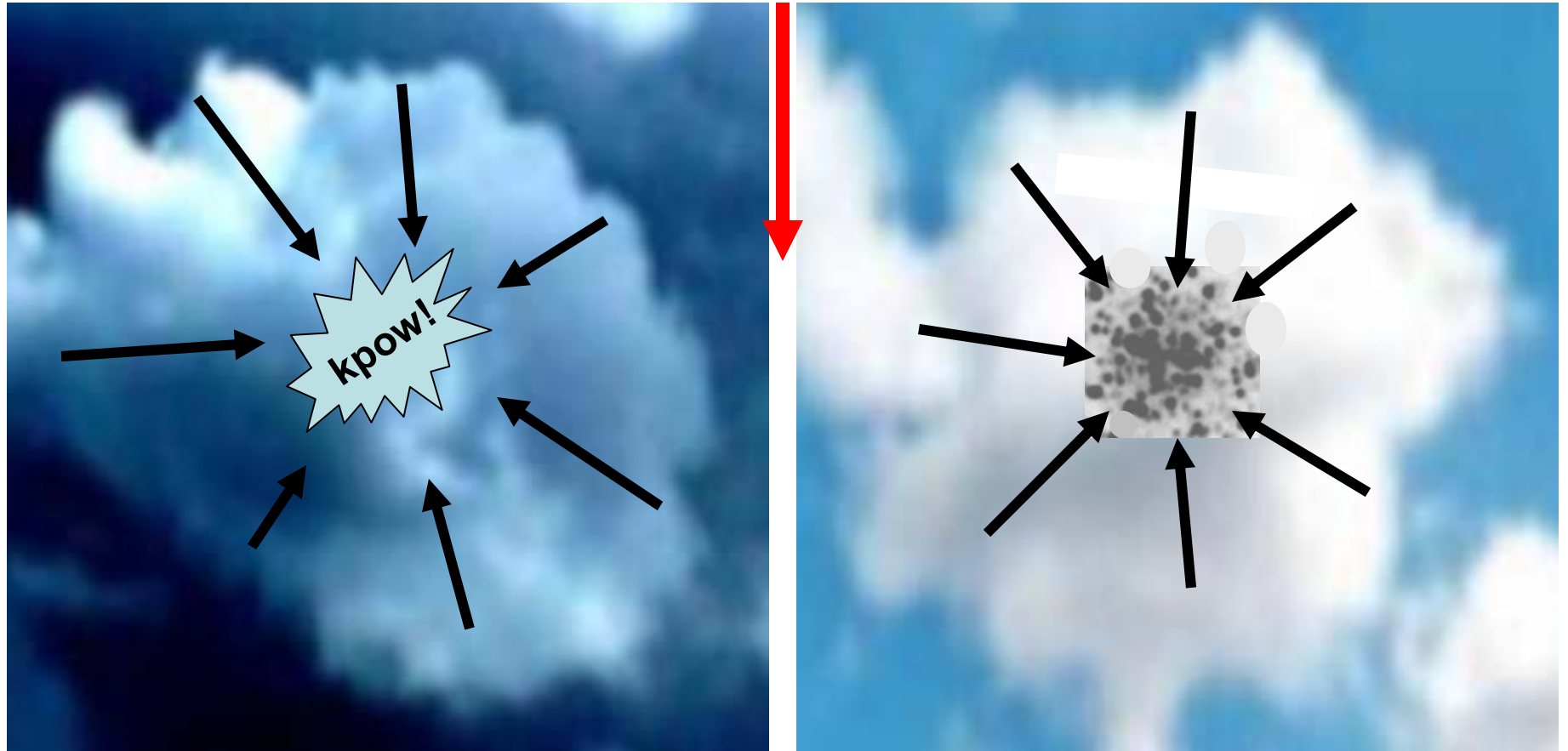
Cloud collapse \rightarrow fractionation \rightarrow GCs instead of SMS

So, { Different kind of progenitors
Different kind of SN
Different kind of Chem. Evol. } after $[Fe/H] \approx -2.5$



How physics of star formation at this epoch effects this transition is above my pay grade.

-2.7



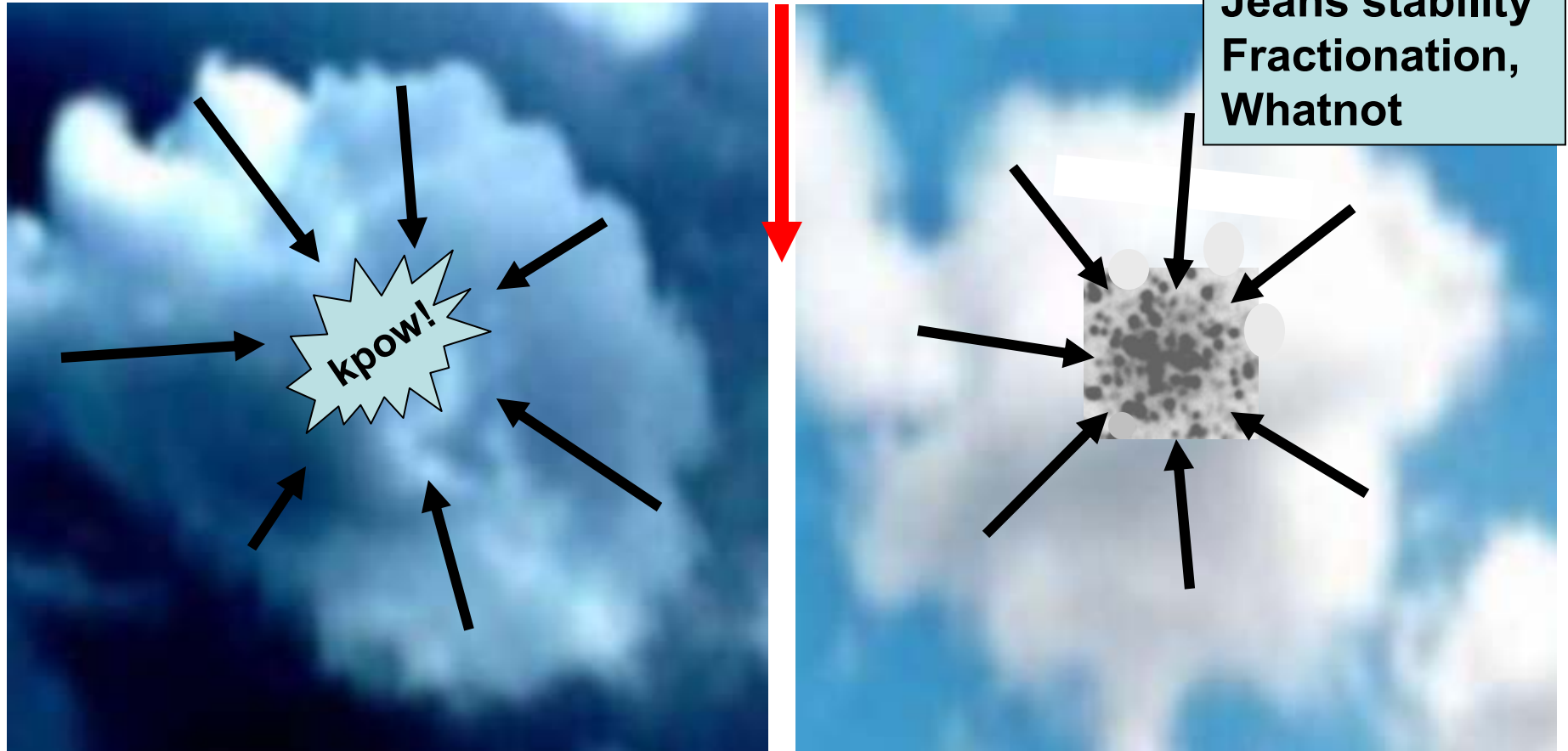
< -3.0

> -2.5

How physics of star formation at this epoch effects this transition is above my pay grade.

-2.7

Cooling agents
Jeans stability
Fractionation,
Whatnot



< -3.0

> -2.5

An aside: supermassive stars are old stuff, of course
– at least 34 years old

Linear Series of Stellar Models. VI.
Supermassive Stars

by

B. Paczyński

N. Copernicus Astronomical Center, Polish Academy of Sciences, Warsaw

and

M. Różyczka

Warsaw University Observatory

Received February 22, 1977

1977

These people think big!
(log M, not M, folks!)

ABSTRACT

Linear series of stellar models on hydrogen burning main sequence is constructed for Population I composition ($X = 0.7$, $Y = 0.27$, $Z = 0.03$) for masses in the range $1.8 \leq \log M/M_{\odot} \leq 12.0$. No terminating point is found in the supermassive star region. For $\log M/M_{\odot} > 9$ the main sequence models develop a giant type structure and various assumptions underlying our computations become invalid.

Models with $\log M/M_{\odot} > 5.7$ are known to be thermally unstable even though there is no turning point of the linear series. This phenomenon is caused by the onset of dynamical instability at the same mass.

An aside: supermassive stars are old stuff, of course
– at least 34 years old

Linear Series of Stellar Models. VI.
Supermassive Stars

by

B. Paczyński

**A bit of ancient Lick Observatory history :
Paczynski, a graduate student of Warsaw University,
worked as a research assistant at LO in 1962-63.**

Warsaw University Observatory

Received February 22, 1977

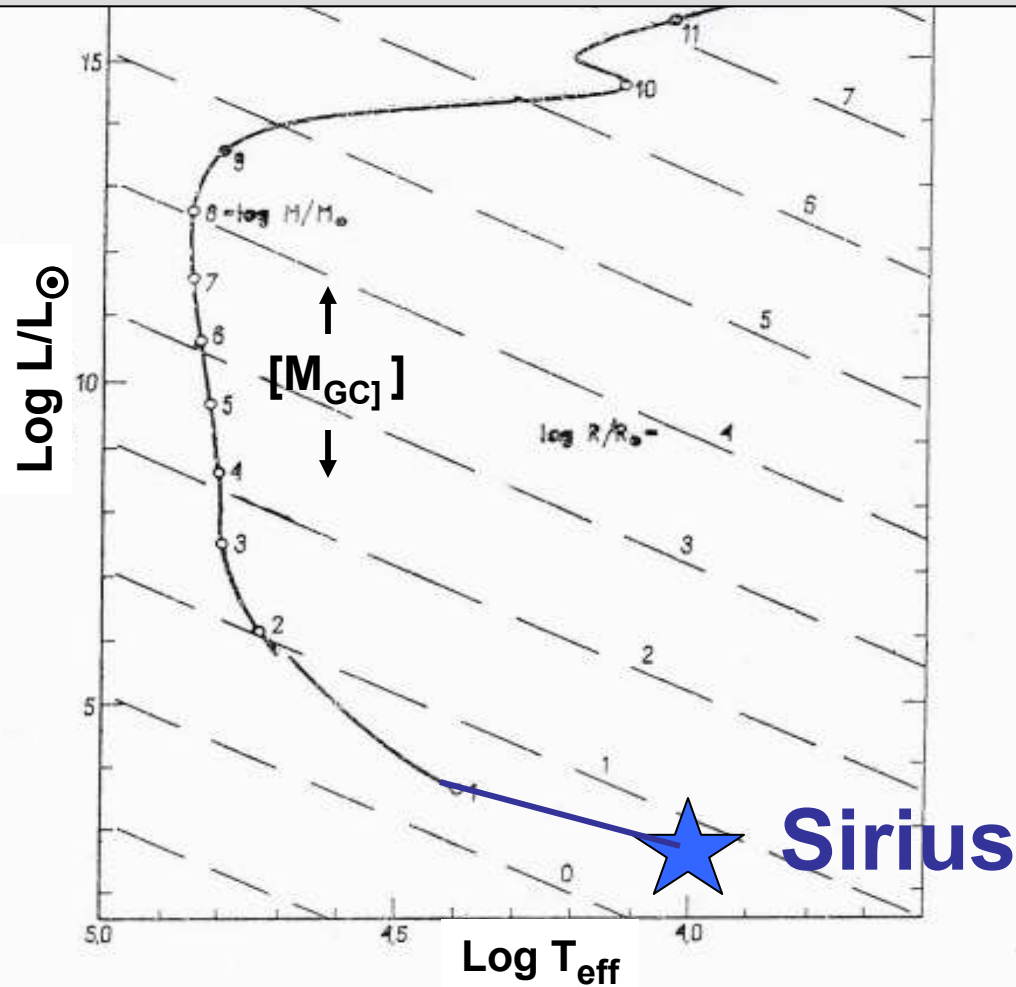
**These people think big!
(log M, not M, folks!)**

ABSTRACT

Linear series of stellar models on hydrogen burning main sequence is constructed for Population I composition ($X = 0.7$, $Y = 0.27$, $Z = 0.03$) for masses in the range $1.8 \leq \log M/M_{\odot} \leq 12.0$. No terminating point is found in the supermassive star region. For $\log M/M_{\odot} > 9$ the main sequence models develop a giant type structure and various assumptions underlying our computations become invalid.

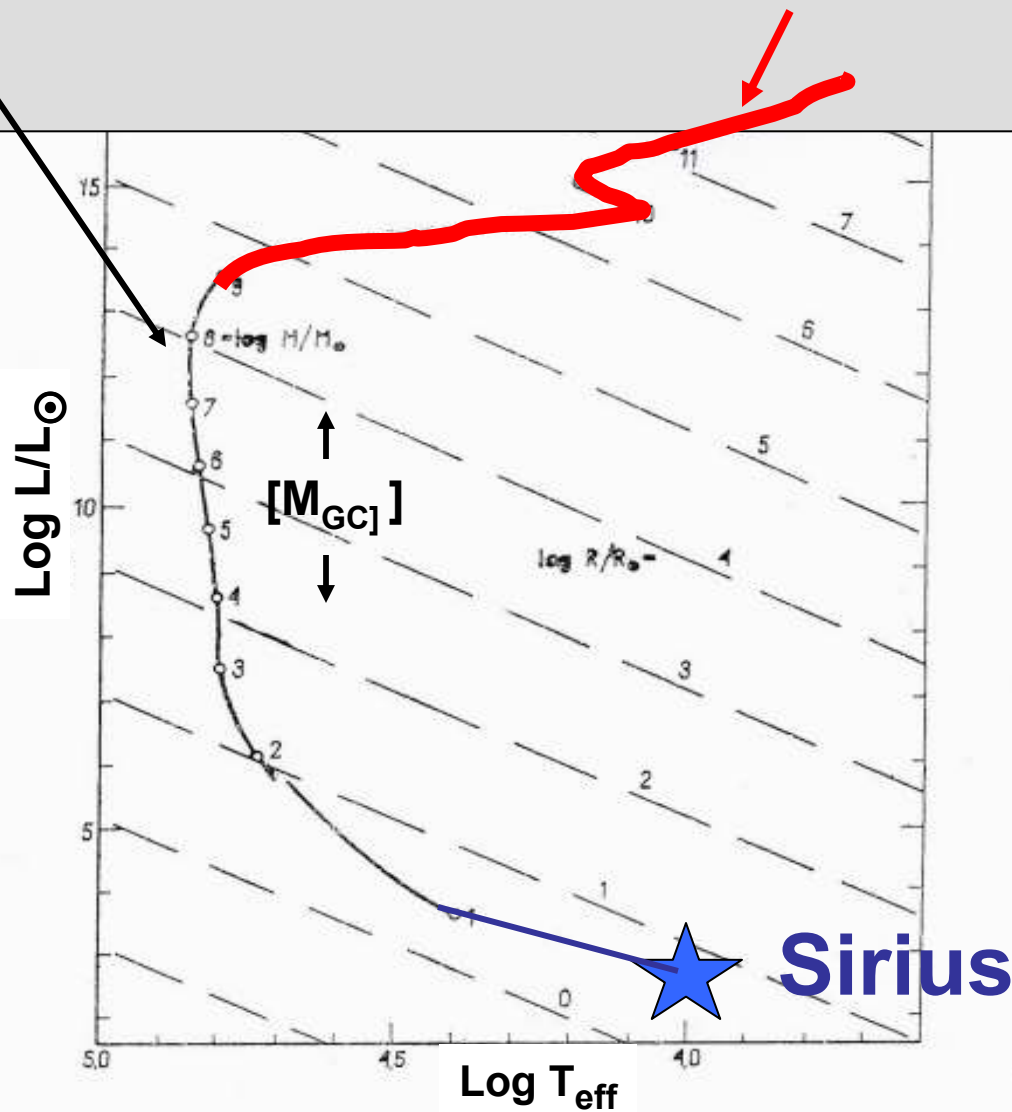
Models with $\log M/M_{\odot} > 5.7$ are known to be thermally unstable even though there is no turning point of the linear series. This phenomenon is caused by the onset of dynamical instability at the same mass.

With regard to their SMS main sequence I quote the authors:
“For masses greater than $10^9 M_{\odot}$,
main sequence models develop a giant type structure.”



1. Position of the main sequence models ($X = 0.7, Z = 0.03$) on the H-R diagram. Open circles are labelled with a logarithm of stellar mass.

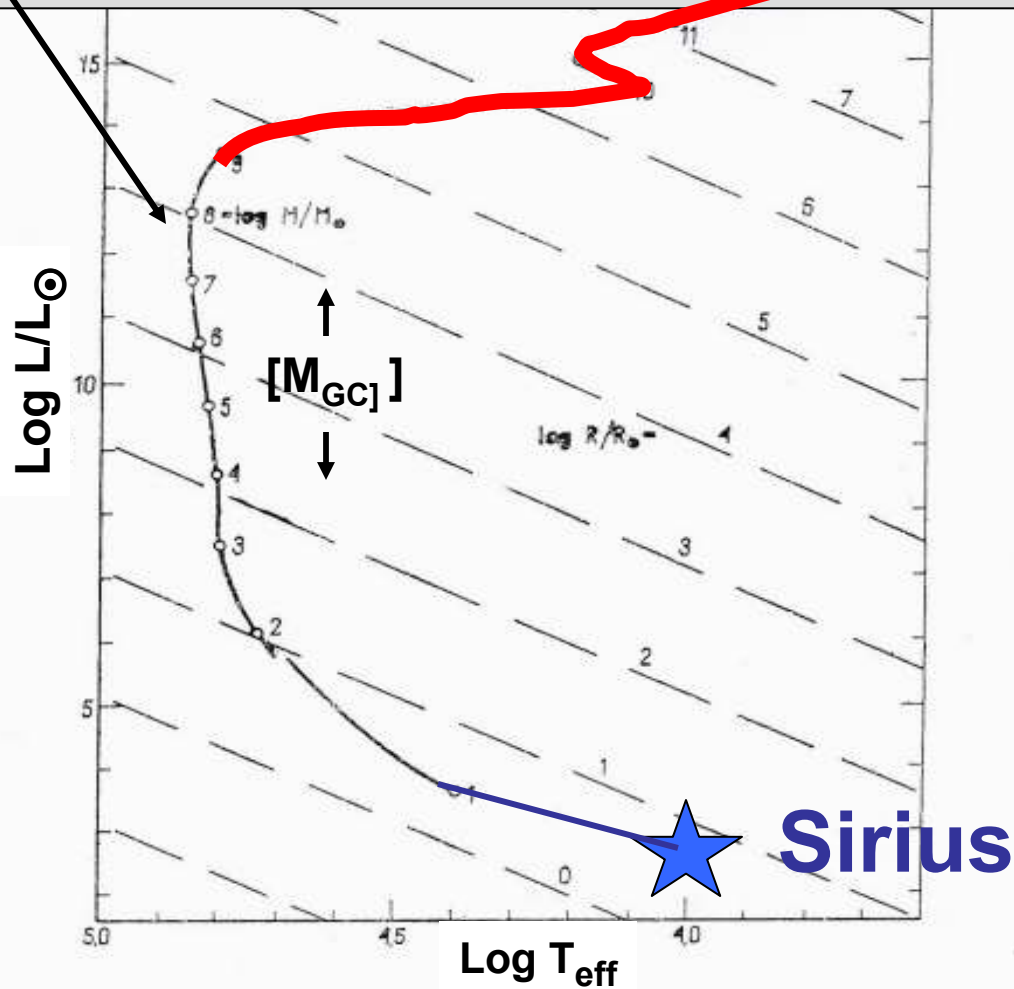
With regard to their SMS main sequence I quote the authors:
“**For masses greater than $10^9 M_{\odot}$,
main sequence models develop a giant type structure.**”



1. Position of the main sequence models ($X = 0.7, Z = 0.03$) on the H-R diagram. Open circles are labelled with a logarithm of stellar mass.

With regard to their SMS main sequence I quote the authors:
“**For masses greater than $10^9 M_{\odot}$,
main sequence models develop a giant type structure.**”

Oh.... OK.



1. Position of the main sequence models ($X = 0.7, Z = 0.03$) on the H-R diagram. Open circles are labelled with a logarithm of stellar mass.

Now,
we turn to the globular clusters & our honoree
Mrr. Krraft,
as his mentorr Otto Strruve called him.

As we have learned
during this meeting

**Globular cluster stars
rank among
the great mysteries
of the world**

Globular cluster stars rank among the great mysteries of the world



creators of such mysteries

Leonardo da Vinci

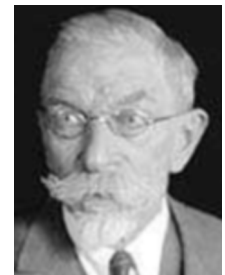
Pierre de Fermat

Sergei N. Blazhko

Pieter Th. Oosterhoff

Sidney van den Bergh

Robert P. You-Know-Who



Is she or isn't she?

1519

Mona Lisa's smile a mystery
by [Ewen Callaway](#)

If you have been puzzled by the Mona Lisa's smile, you are not alone. It is a moment of ambiguity and serious thought that happens because our eyes are not sure about her smile. Different categories of information encode data about an object in the visual field.



The Mystery of Mona Lisa's Smile
Linked to Flickering Eyes

e's radi
are over. It
the brain

The Italians have a word to explain Mona Lisa's smile: **sfumato**. It means blurry, ambiguous and up to the imagination.

Who is *Mona Lisa*?

Many questions arose over the identity of the lady in the portrait, which means "the light and shadow." The name *La Joconde*, carries a

and theories about the *Mona Lisa*'s smile. One popular [theory](#) suggests that the lady is the Duchess of Milan, Isabella of Aragon.

And who was she?

Neuroscientists at the Institute of Neuroscience in Alicante, Spain claim to have solved the mystery as to why some people see Mona Lisa's smile and some don't.

Mona Lisa's smile mystery solved



[Art Glossary: Sfumato - Painting Techniques Sfumato](#)

[painting.about.com/od/artglossary/a/01defsfumato.htm - Cached/Similar](#)
An explanation of what the painting technique *sfumato* involves.

[Sfumato and Chiaroscuro -- Painting in the Style of Old Masters ...](#)

[painting.about.com/od/oldmstert/techniques/c/sfumato_chiaroscuro.htm - Cached/Similar](#)
Don't be kept in the dark by these two important terms, *sfumato* and ...
[Show more results from about.com](#)

[Sfumato - Wikipedia, the free encyclopedia](#)

[en.wikipedia.org/wiki/Sfumato - Cached/Similar](#)
Detail of the face of Mona Lisa showing the use of *sfumato*, particularly in the shading ...
Renaissance Art: A Topical Dictionary, Greenwood Press, pp. 263. ...

[Glossary of Art Terms and Definitions](#)

[www.funartwork.com/Glossary_of_Art_Terms_and_Definitions.html - Cached/Similar](#)
Sep 29, 2006 Sfumato - The term *sfumato* was coined by Italian Renaissance artist, Leonardo da Vinci, and refers to a fine art painting technique of ...

[sfumato \(painting technique\) -- Britannica Online Encyclopedia](#)

[www.britannica.com/EBchecked/topic/31737/sfumato - Cached/Similar](#)
Facts about *sfumato*: development during Renaissance, as discussed in Western painting (art); Leonardo da Vinci; No Additional Content ...

[sfumato: Definition from Answers.com](#)

[www.answers.com/Library/Literature & Language - Cached/Similar](#)
sfumato n. The blurring or softening of sharp outlines in painting by subtle and ... A soft form in art created by blurring the light and dark edges. ...

Is she or isn't she?

1519

Mona Lisa's smile a m

by

If y

mo

hap

abo

cate

encode data about an o

in the visual field.

The M

I just had to include da Vinci
after I learned about "sfumato".
What a wonderful word!

sa's Smile

s

m

The Italians have a word to express... up to the imagination.

sfumato. It means blurry, ambiguous...
(just like lots of contemporary astrophysics - GWP)

Neuroscientists at the Institute of Neuroscience in Alicante, Spain claim to have solved the mystery as to why some people see Mona Lisa's smile and some don't.

Mona Lisa's smile mystery solved

And who was she?

Who was she?

Many questions arose o

of the lady in the portrait

which means "the light

La Joconde, carries a

and theories about the *Mona Lisa's* smile. One popular theory

suggests that the lady is the Duchess of Milan, Isabella of Aragon.

Fermat's last theorem is a theorem first proposed by Fermat in the form of a note scribbled in the margin of his copy of the ancient Greek text *Arithmetica* by Diophantus. The scribbled note was discovered posthumously

Did he or didn't he?

1637

Arithmeticonum Lib. II.

85

tertio quadratum, & Casus idem hic etiam locus habebunt, ut manifeste
 Pencil

QVÆSTIO VIII

PROPOSITVM quadratum
 dividere in duos quadratos.
 Imperatum sit ut 16. dividatur
 in duos quadratos. Ponatur
 primus 1. Q. Oportet igitur 16
 - 1. Q. equali esse quadrato.
 Fingo quadratum à numeris
 quotquot libuerit, cum defe-
 ctu tot unitatum quot conti-

ΤΟΝ τετραγώνον πρῶτον
 διαιρῆν εἰς δύο τετραγώνους. ἰ-
 πεταχθέν δὲ τὸ 16 εἰς δύο τε-
 τραγώνους. καὶ πρῶτον ἰσοπέπ-
 δυνάμους μᾶς. Ἄρα ἄρα μὴδ-
 δας 15 λείπει δύναμους μᾶς ἴσας
 τῷ πρῶτῳ. πλάσσει τὸ τετραγώ-
 νον



$x^n + y^n = z^n$
 $n > 2$
 $x, y, z \neq 0$

Fermat's last theorem is a theorem first proposed by Fermat in the form of a note scribbled in the margin of his copy of the ancient Greek text *Arithmetica* by Diophantus. **Did he or didn't he?** **1637** The scribbled note was discovered posthumously

Arithmeticonum Lib. II 8c



In the words of mathematical historian [Howard Eves](#),
"Fermat's Last Theorem has the peculiar distinction
of being the mathematical problem for which the greatest
number of incorrect proofs have been published."
(including Fermat's own?)

in duas quadratos. Ponatur
primus 1 . Q . Oportet igitur 16
 $- 1$ Q . equalis esse quadrato.
Fingo quadratum à numeris
quotquot libuerit, cum defe-
cta tot unitatum quot conti-

πλάζου δὲ ἢ ἰς δύο π-
ραγμάτων. καὶ πλάζου ἰσοῦται
δυνάμει μίας. Ἄρα ἄρα μιά-
δος ἰσὺ λαμβάνει δύναμις μίας ἰσας
ἢ πλάζου. πλάζου ἢ πλάζου

$x^n + y^n = z^n$
 $n > 2$
 $\neq 2!!$

Fermat's last theorem is a theorem first proposed by Fermat in the form of a note scribbled in the margin of his copy of the ancient Greek text *Arithmetica* by Diophantus. The scribbled note was discovered posthumously

Did he or didn't he?

1637

Arithmeticonum Lib. II &c



In the words of mathematical historian [Howard Eves](#),
"Fermat's Last Theorem has the peculiar distinction
of being the mathematical problem for which the greatest
number of incorrect proofs have been published."
(including Fermat's own?)

in duas quadratos. Ponatur
primus 1 . Oportet igitur
 -1 Q . equalis esse quadrato.
Fingo quadratum à numero
quotquot libere, cum defe-
cta tot unitatum quot conti-

LIKE
[A Short-Form Proof of Fermat's Last Theorem](#)
www.fermatproof.com/ - [Cached Similar](#)
Tom Ballard solves *Fermat's Last Theorem*,
Puts Wiles To Shame.

БЛАЖКО ЕФФЕКТИ 1907

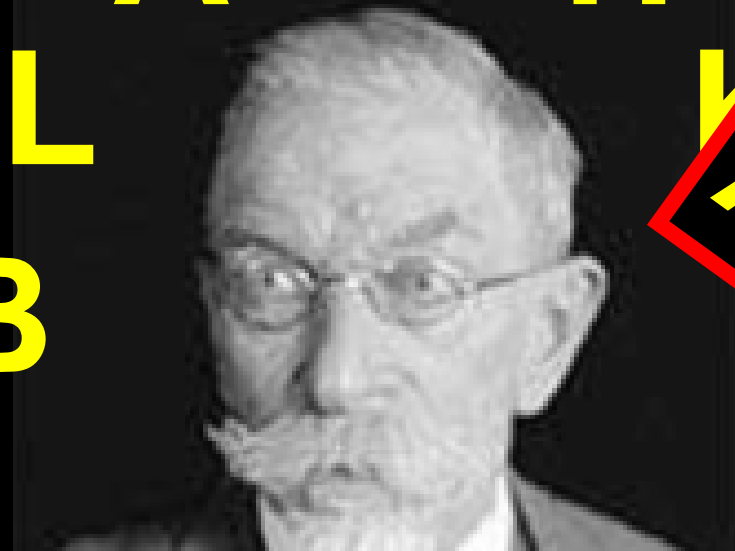


Sergei Nikolaevich Blazhko

АБВГГДЪГЕЁЄЖЗСИЙІЙКЛЉМНЊОПРСТЪКУЎФХЦЦШЩЪЫЬЭЮ

БЛАЖКО ЕФФЕКТ! 1907

A plague for students
of Solon Bailey's "cluster-type" variable stars



L
B

Aaargh!

A mystery only pundits can love.
I spare you the details.

Sergei Nikolaevich Blazhko

АБВГГДЪГЕЁЄЖЗСИЙІЙКЛЉМНЊОПРСТЪКУЎФХЦЦШЩЪЫЬЭЮ

Do they or don't they?

1939

The Oosterhoff effect

van den Bergh, S. 1993, MNRAS, 262, 588.

“It is found (**with greater than 99 percent confidence**) that globular clusters of Oosterhoff class I lie on retrograde orbits more frequently than do clusters of Oosterhoff class II. A possible explanation for the observed correlation is **blah, blah, blah,**”

A plethora of other distinguishing oddities are known about Oosterhoff clusters.

Sidney cannot be blamed for all of them.

Do they or don't they?

1939

The Oosterhoff effect

[van den Bergh, S.](#) 1993, MNRAS, 262, 588.

This conclusion was reaffirmed by

LEE & CARNEY, 1999, AJ, 118, 1373
RR LYRAE LUMINOSITY DIFFERENCES BETWEEN OOSTERHOFF
GROUP I AND II CLUSTER SYSTEMS AND THE ORIGIN OF
THE OOSTERHOFF DICHOTOMY

but it has not attracted much attention.

Sidney cannot be blamed for all of them.

“It is (with more than 99 percent confidence) that
globally, RR Lyrae stars lie on retrograde orbits
m... Oosterhoff class II. A
pos... blah,
A plethora of other dist... about
Oosterhoff...

**Finally,
we come to
Bob
Himself**

Bob has been near the center of globular cluster research ...



... for a long time

1973

ON THE METAL ABUNDANCE OF RR LYRAE STARS
IN THE GLOBULAR CLUSTER M22*

DENNIS BUTLER, ROBERT P. KRAFT, J. S. MILLER, AND LLOYD B. ROBINSON

Lick Observatory, Board of Studies in Astronomy and Astrophysics,
University of California, Santa Cruz

Received 1972 November 8

ABSTRACT

The Preston system for measuring metal abundances of RR Lyrae stars has been reestablished from equivalent-width measurements of $H\gamma$, $H\delta$, and $Ca\ II\ K$ in a slit image-tube scanner. The RR Lyraes of M22 are found to be metal-rich, in agreement with the expectation of metal richness as judged from the small slope of the giant branch.

ApJ 1973



... for a long time

1973

ON THE METAL ABUNDANCE OF RR LYRAE STARS
IN THE GLOBULAR CLUSTER M22*

DENNIS BUTLER, ROBERT P. KRAFT, J. S. MILLER, AND LLOYD B. ROBINSON

Lick Observatory, Board of Studies in Astronomy and Astrophysics,
University of California, Santa Cruz

Received 1972 November 8

ABSTRACT

The Preston system for measuring metal abundances of RR Lyrae stars has been reestablished from equivalent-width measurements of $H\gamma$, $H\delta$, and $Ca II K$ in a spectrograph image-tube scanner. The RR Lyraes of M22 are found to be metal-poor, in agreement with the expectation of metal richness as judged from the small slope of the giant branch.

ApJ 1973



GCs are his **current** targets after fifty distinguished years of:

Cepheids and RR Lyrae stars: atmospheres, ρu^3 , etc

CV/Novae/Binary phenomena

gravitational radiation

mass transfer

Main Sequence Rotation

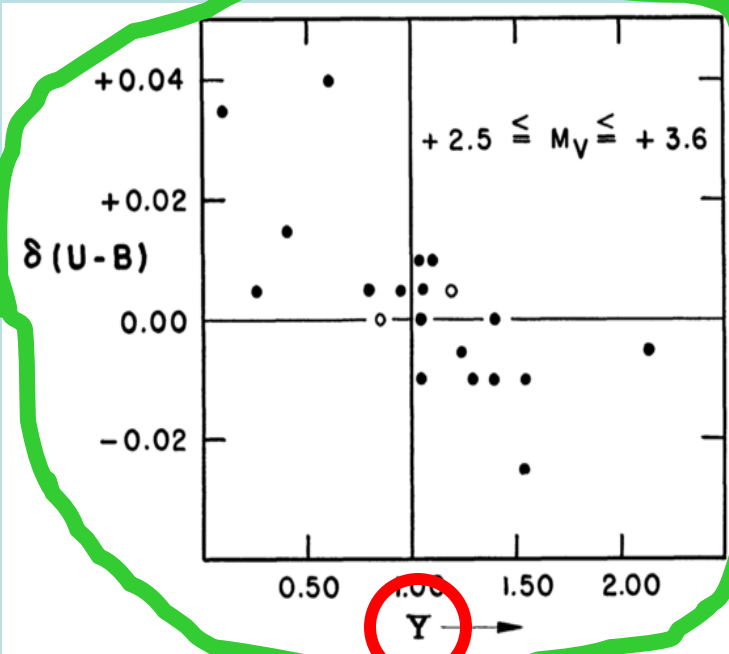
And on and on and on ...

More later

Another aside: Bob was talking about effects of stellar rotation at an IAU symposium somewhere in Europe.



**speaking
at the podium**



**sitting
In the audience**

Y is some godawful stellar rotation parameter

Main Sequence Rotation
And on and on.

1965

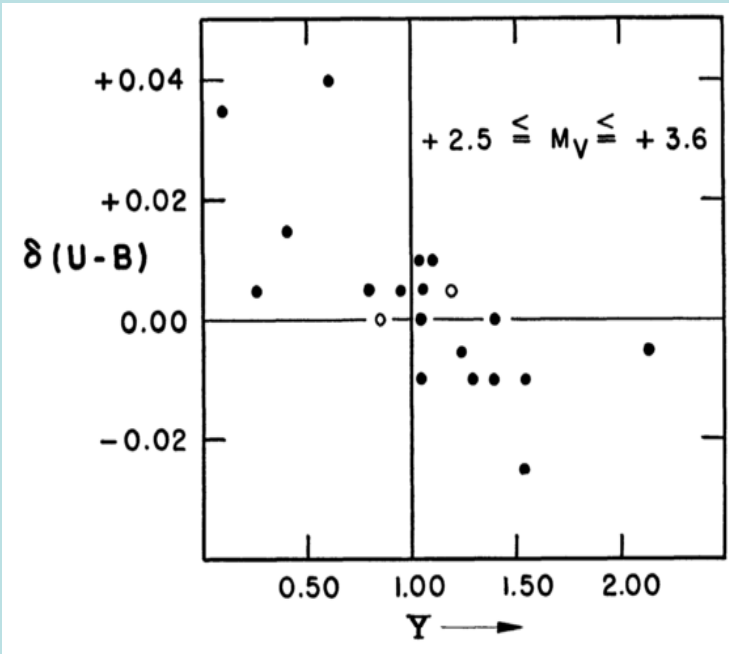
DENN
The Pr
from equ
image-tub
expectatio

GCs
Cep
CV/
g
mass transfer

Another aside: Bob was talking about effects of stellar rotation at an IAU symposium somewhere in Europe.



**speaking
at the podium**



**sitting
In the audience**

Overheard afterwards in the audience:

“All I wanted to know is what is the mean angular momentum density in the Galaxy!”

Main Sequence Rotation
And on and on.

1965

... for a long time

1975

ON THE METAL ABUNDANCE OF RR LYRAE STARS
IN THE GLOBULAR CLUSTER M22*

DENNIS BUTLER, ROBERT P. KRAFT, J. S. MILLER, AND LLOYD B. ROBINSON

Lick Observatory, Board of Studies in Astronomy and Astrophysics,
University of California, Santa Cruz

Received 1972 November 8

ABSTRACT

The Preston system for measuring metal abundances of RR Lyrae stars has been reestablished from equivalent-width measurements of $H\gamma$, $H\delta$, and $Ca II K$ in a
image-tube scanner. The RR Lyraes of M22 are found to be metal
expectation of metal richness as judged from the small slope of the giant

ApJ 1973

05.05.04 Carbon and Nitrogen Abundances in the
Atmospheres of Subgiant and Asymptotic Giant Branch
Stars in M92. DENNIS BUTLER*, DUANE CARBON and
ROBERT P. KRAFT, Lick Obs., Bd. of Studies in Ast. and
Ap., U. of Calif., Santa Cruz. - Using the Wampler-
Robinson Image Tube Scanner operated at the Cassegrain
focus of the 120-inch telescope, we have observed a
sample of subgiant (SG) and asymptotic giant branch
(AGB) stars in M92, in the wavelength interval
 $\lambda\lambda 3100 \text{ \AA} - 5100 \text{ \AA}$, with resolution \sim
of 28 stars is confined largely to the

BAAS 1975



... for a long time

1975

ON THE METAL ABUNDANCE OF RR LYRAE STARS
IN THE GLOBULAR CLUSTER M22*

DENNIS BUTLER, **ROBERT P. KRAFT**, J. S. MILLER, AND LLOYD B. ROBINSON

Lick Observatory, Board of Studies in Astronomy and Astrophysics,
University of California, Santa Cruz

Received 1972 November 8

ABSTRACT

The Preston system for measuring metal abundances of RR Lyrae stars has been reestablished from equivalent-width measurements of $H\gamma$, $H\delta$, and $Ca II K$ in a slit image-tube scanner. The RR Lyraes of M22 are found to be metal-poor, in agreement with the expectation of metal richness as judged from the small slope of the giant branch in the $H-R$ diagram.

ApJ 1973

05.05.04 Carbon and Nitrogen Abundances in the Atmospheres of Subgiant and Asymptotic Giant Branch Stars in M92. DENNIS BUTLER*, DUANE CARBON and **ROBERT P. KRAFT**, Lick Obs., Bd. of Studies in Ast. and Ap., U. of Calif., Santa Cruz. - Using the Wampler-Robinson Image Tube Scanner operated at the Cassegrain focus of the 120-inch telescope, we have observed a sample of subgiant (SG) and asymptotic giant branch (AGB) stars in M92, in the wavelength interval $\lambda\lambda 3100 \text{ \AA} - 5100 \text{ \AA}$, with resolution $\sim 10 \text{ \AA}$. The abundance of 28 stars is confined largely to the

BAAS 1975



characteristically playing second fiddle, so that youngsters could get their names up in lights

... for a long time

1975

ON THE METAL ABUNDANCE OF RR LYRAE STARS
IN THE GLOBULAR CLUSTER M22*

DENNIS BUTLER, ROBERT P. KRAFT, J. S. MILLER, AND LLOYD B. ROBINSON

Lick Observatory, Board of Studies in Astronomy and Astrophysics,
University of California, Santa Cruz

Received 1972 November 8

ABSTRACT

The Preston system for measuring metal abundances of RR Lyrae stars has been reestablished from equivalent-width measurements of $H\gamma$, $H\delta$, and $Ca II K$ in a spectrograph image-tube scanner. The RR Lyraes of M22 are found to be metal-rich, in agreement with the expectation of metal richness as judged from the small slope of the giant branch in the $H-R$ diagram.

ApJ 1973

05.05.04 Carbon and Nitrogen Abundances in the Atmospheres of Subgiant and Asymptotic Giant Branch Stars in M92. DENNIS BUTLER*, DUANE CARBON and ROBERT P. KRAFT, Lick Obs., Bd. of Studies in Ast. and Ap., U. of Calif., Santa Cruz. - Using the Wampler-Robinson Image Tube Scanner operated at the Cassegrain focus of the 120-inch telescope, we have observed a sample of subgiant (SG) and asymptotic giant branch (AGB) stars in M92, in the wavelength interval $\lambda\lambda 3100 \text{ \AA} - 5100 \text{ \AA}$, with resolution $\sim 10 \text{ \AA}$. The abundance distribution of 28 stars is confined largely to the

BAAS 1975



while he assembled his overview

ON THE NONHOMOGENEITY
OF METAL ABUNDANCES IN
STARS OF GLOBULAR CLUSTERS
AND SATELLITE SUBSYSTEMS
OF THE GALAXY

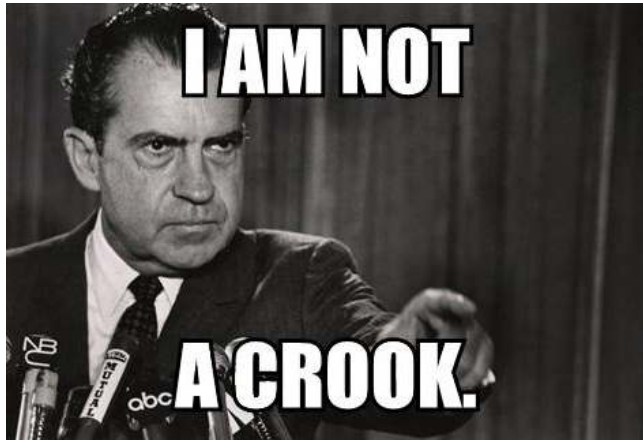
Robert P. Kraft

Lick Observatory, Board of Studies in Astronomy and Astrophysics,
University of California, Santa Cruz, California 95067

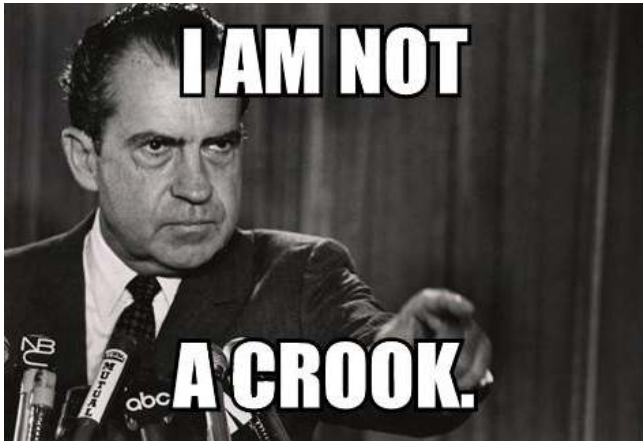
ARA&A 1979

**While we recall all
of his admirable traits,
we must also acknowledge
that**

Bob has a severe Watergate problem.

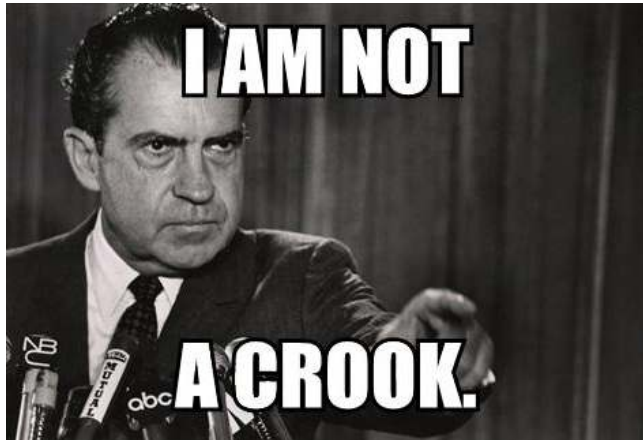


Bob has a severe Watergate problem.



It is an issue that has bedevilled
our honoree for decades,

Bob has a severe Watergate problem



It is an issue that has bedevilled
our honoree for decades,
giving rise to sporadic,
sometimes uncontrollable outbursts

like this



RPK's bête noir

1972



Growls accompanied
by shaking of jowls

RPK, impersonator *nonpareil*



I was not present, but Steve Shectman tells me that Bob, while President of the AAS, lost control of himself in this manner for some time during an after-dinner speech. Questioned later, Bob could not recall the episode for 18 minutes.

However, his passion is a redeeming virtue



**Whenever I think of Bob,
this pose reminds me of his passions,
particularly in all matters of social justice,**

However, his passion is a redeeming virtue



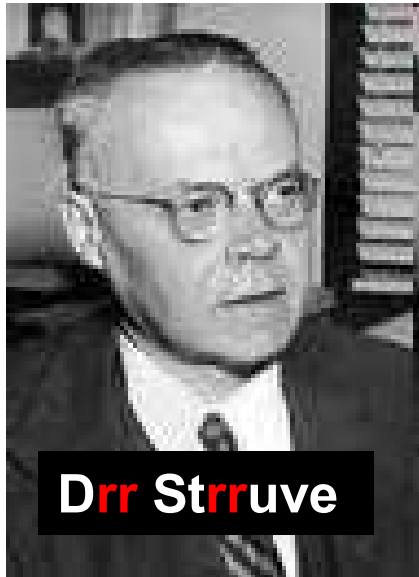
Whenever I think of Bob,
this pose reminds me of his passions,
particularly in all matters of social justice,
and particularly after a bottle of fine wine!

A few shared experiences from the distant past

Berkeley Astronomy 1955



The Tiger



Drr Strruve

Mr. Preston.. feel in your bones? Mr Kraft please stand up ..



LO 1963

H α

Microphotometer
room

Wasiutinski and all that ...

What really impressed Bob was

ρU^3

He would use it every time the topic of cool stellar atmospheres came up.

UCSC 1967 \pm

NSF Proposal

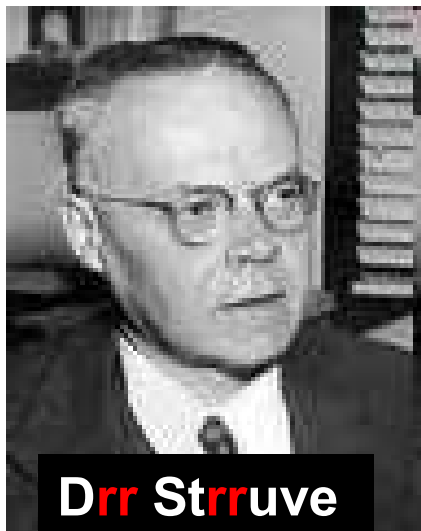
(The recently transplanted LO staff wasn't quite up to running a distinguished academic program on the Santa Cruz campus)

A few shared experiences from the distant past

Berkeley Astronomy 1955



The Tiger



Drr Strruve

What a friendly department!



1963
H α
photometer
room

Wasiutinski and all that

What really impressed [redacted] was

How to use it every time the topic of cool stellar atmospheres came up.

UCSC 1967±

Proposal

(The recently [redacted] LO staff wasn't [redacted] up running a distinguished academic program on the Santa Cruz campus)

A few shared experiences from the distant past

Berkeley Astro



The Tiger

During construction of his stellar evolution code Henyey relied heavily on **Wasiutinski's treatment of convection**, and Henyey's esteem for JW rubbed off on all of his students ...

Drr Struve

What a friendly department!

Mr. Preston.. feel in your bones? Mr Kraft please stand up ..

Wasiutinski and all that ...
What really impressed Bob was

$$\rho u^3$$

He would use it every time the topic of cool stellar atmospheres came up.

Jeremi Wasiutynski



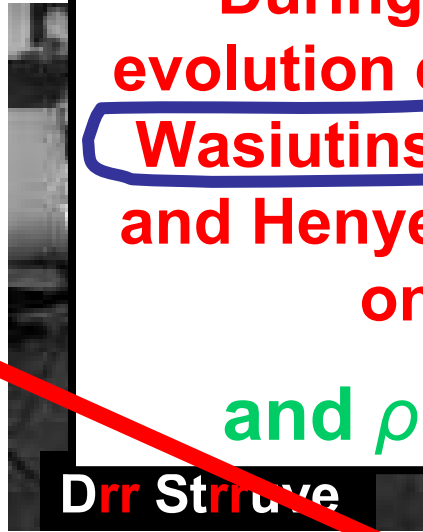
Norske Videnskapsakademie 1946

A few shared experiences from the distant past

Berkeley Astro



The Tiger



Drr Struve

During construction of his stellar evolution code Henyey relied heavily on **Wasiutinski's treatment of convection**, and Henyey's esteem for JW rubbed off on all of his students ...

and ρu^3 stuck in Bob's head

What a friendly department!

Mr. Preston.. feel in your bones? Mr Kraft please stand up ..

Wasiutinski and all that ...
What really impressed Bob was

ρu^3

He would use it every time the topic of cool stellar atmospheres came up.

Jeremi Wasiutynski



Norske Videnskapsakademie 1946

A few shared experiences from the distant past

A sad day at Lick Observatory in 1963

$H\alpha$

In the Microphotometer Room
Our work came to an abrupt end.



THE WIDTH OF $H\alpha$ AS A DISCRIMINANT OF LUMINOSITY
IN THE SPECTRA OF LATE-TYPE STARS*

ROBERT P. KRAFT

Mount Wilson and Palomar Observatories
Carnegie Institution of Washington, California Institute of Technology

GEORGE W. PRESTON

Lick Observatory, University of California, Mount Hamilton, California

AND

SIDNEY CARNE WOLFF

Berkeley Astronomical Department, University of California, Berkeley, California

Received February 17, 1964

An Astronomical Life Salted by Pure Chance

Kraft, Robert P. [2009ARA&A..47, 1](#)

Both of us having graduated from the Berkeley Ph.D. program, Preston and I were great admirers of Otto Struve and of the "Yerkes Observatory style," in which an observatory-based graduate program included a cadre of theorists to complement the observers. Although Lick had appointed Bodenheimer, a theoretician in the field of stellar structure and evolution, to the mountain-top staff shortly before the move to UCSC, we saw the need for theoretical expertise in a number of other areas. At the time, the NSF was offering a group of "Science Departmental Improvement" grants, and with Whitford's approval **Preston and I wrote a proposal** for such a grant to Lick/UCSC in the amount of \$600 K. Following a site visit involving the NSF astronomy program director and an evaluation committee, the funds were secured for a five-year program

UCSC 1967±

NSF Proposal

(The recently transplanted LO staff wasn't quite up to running a complete academic program on the Santa Cruz campus)

In conclusion: Berkeley 1955



**Sunday afternoons
in Albany student village:
4-finger Mozart**

In conclusion: Berkeley 1955



**Sunday afternoons
in Albany student village:
4-finger Mozart**

**The Prestons watched tearfully
as the Krafts
departed to Pasadena.**

