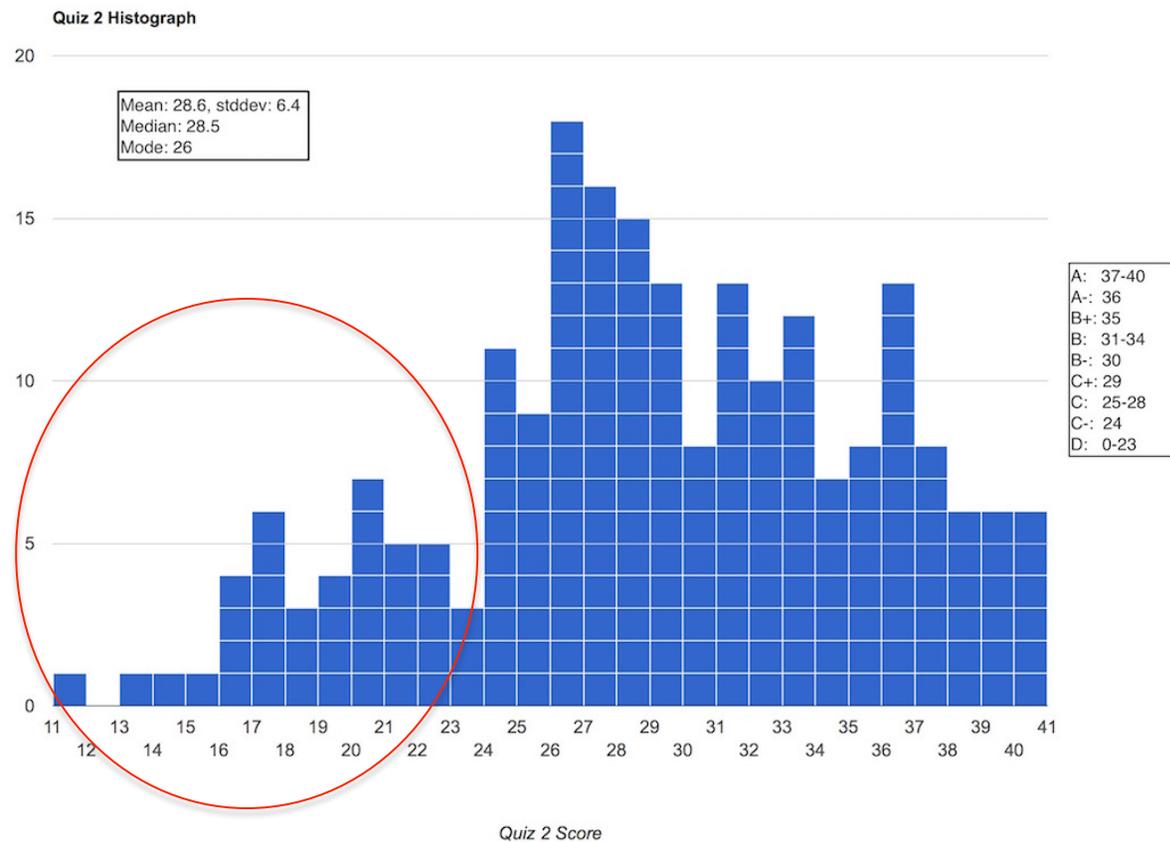


ANNOUNCEMENTS

- Results for Q2 are posted at class website



- Do better on quizzes!
 - Know the answers to homework questions
 - Attend sections (6 per week, you can fit one in) regularly
 - LSS tutor (yshtalen@ucsc.edu)
 - Send me an e-mail for office hours
 - Ask questions in class

SECTION 3

- Binary stars, nova, supernova I
- The origin of the elements
- Special relativity, General relativity, Black Holes
- The Milky Way Galaxy and other galaxies
- The distance scale of the Universe and the Hubble expansion



Evolution of Close Binary Systems

- There are many multiple star systems in the Galaxy, but for the vast majority, the separation of the stars is large enough that one star doesn't affect the evolution of the other(s).
- However, the minority of systems that are spatially close are important for several phenomena

The Algol Mystery

“Algol” is a double-lined eclipsing binary system with a period of about 3 days (very short). The two stars are:

Star A: $3.4M_{\text{Sun}}$ main-sequence star

Star B: $0.8M_{\text{Sun}}$ “subgiant” (evolved) star

What is wrong with this picture?

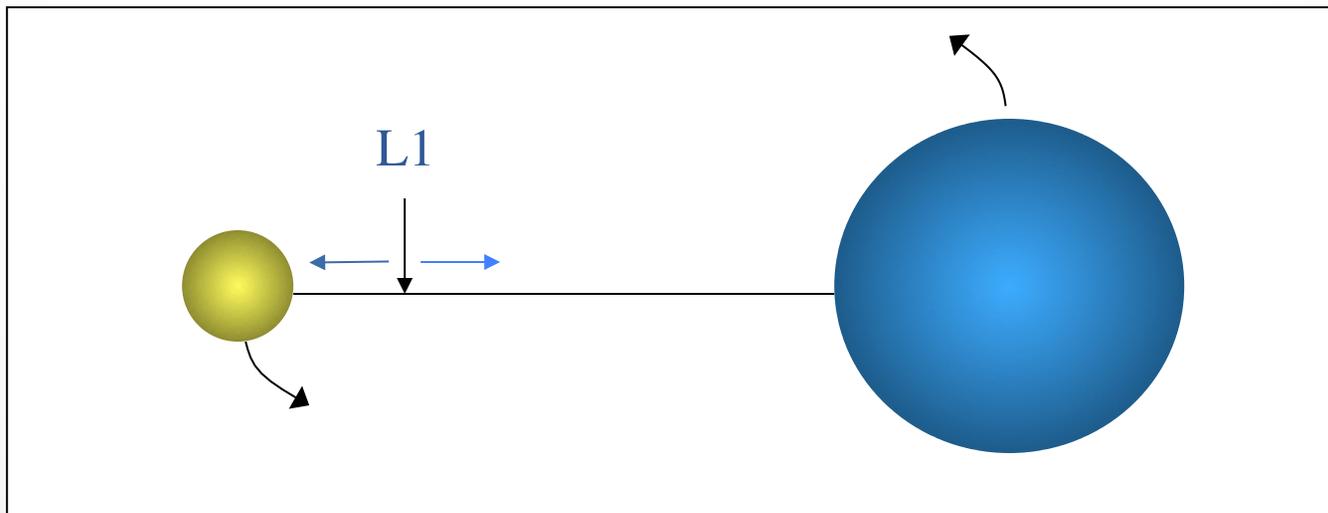
Algol

- The more massive star (A) should have left the main sequence and started up the RGB before the less massive star (B).

What is going on here?

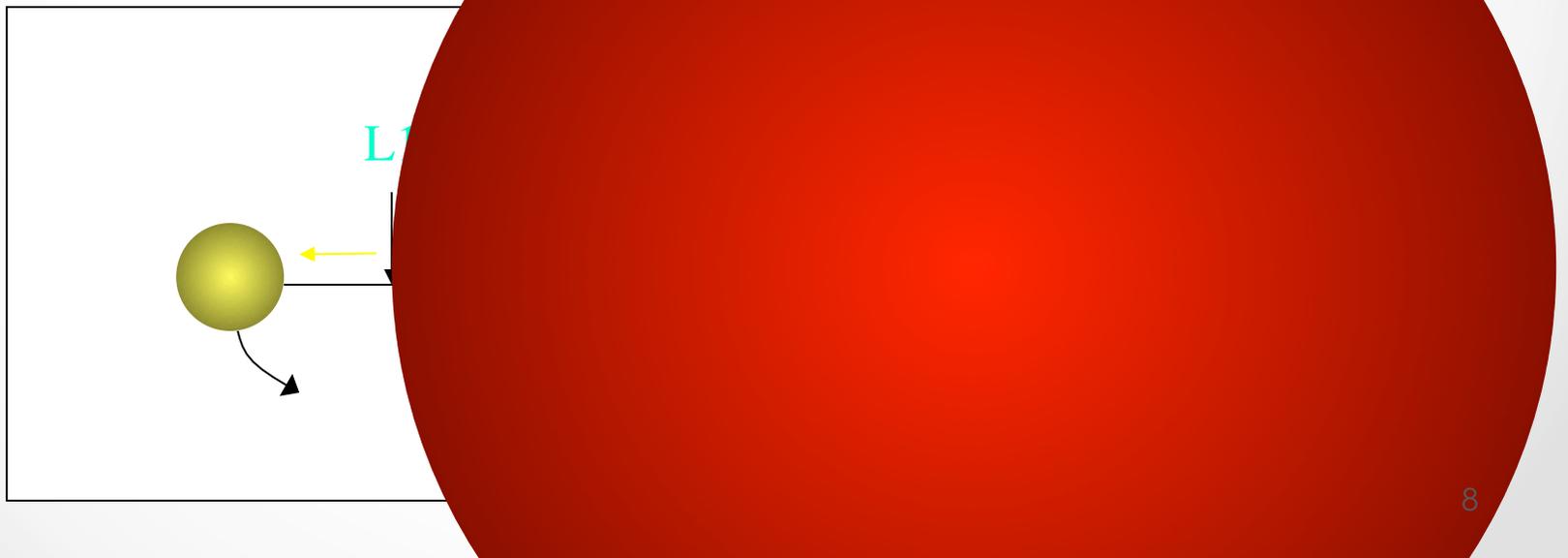
The Algol Story

- Originally the system contained Star A at $1.2M_{\text{Sun}}$ and Star B at $3.0M_{\text{Sun}}$.
- Between the two stars is a point where the gravitational forces of the two stars balance. This is called a “Lagrange point”.



The Algol Story

- Star B evolves up the giant branch and eventually its radius reaches L1.
- As it continues to expand, the Star B outer envelope has a stronger attraction to Star A and mass begins to transfer



Mass Transfer in Binaries

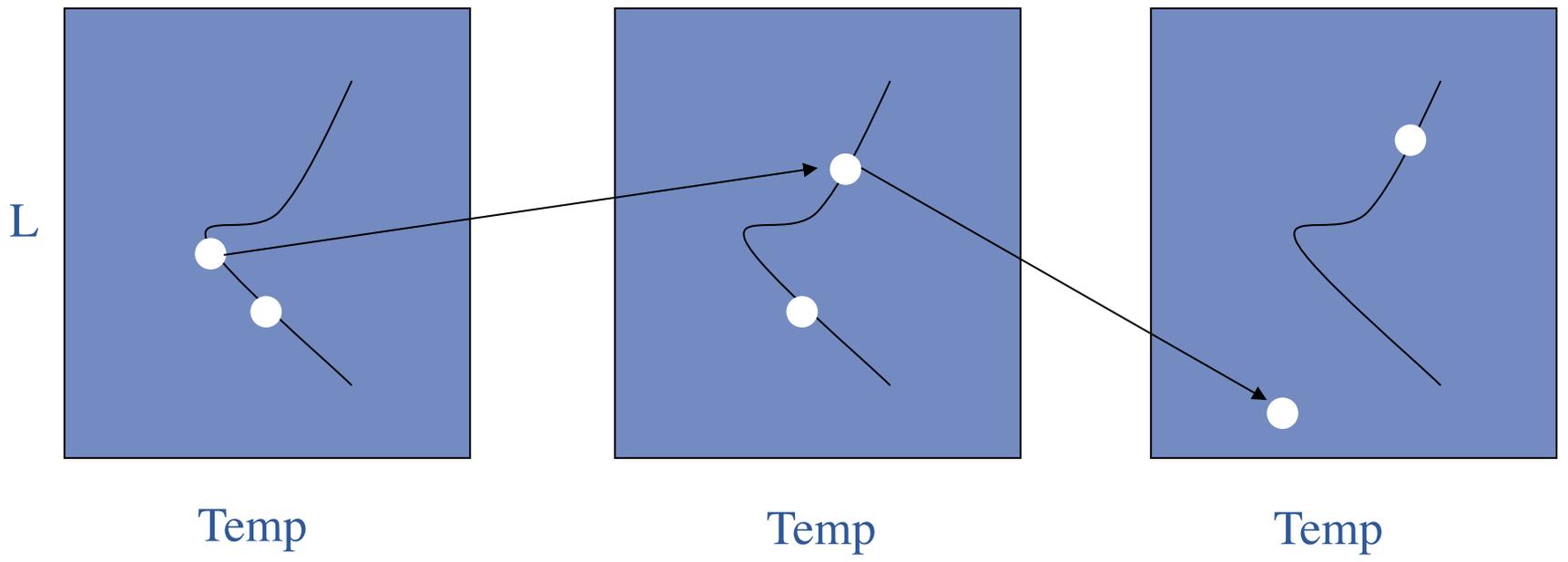
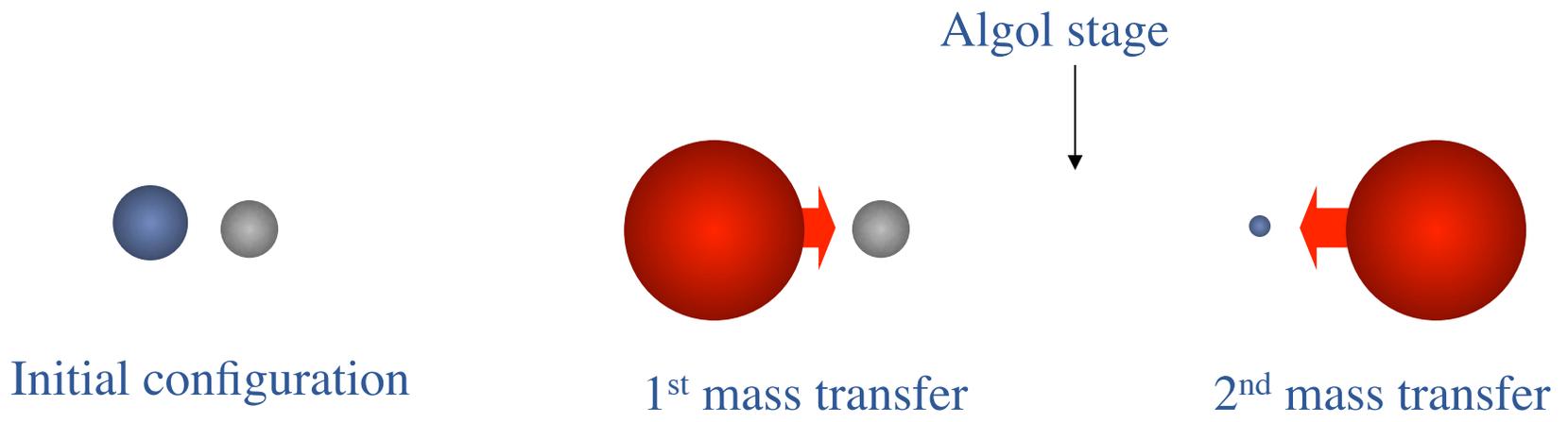
- In the case of Algol, Star B transferred $2.2M_{\text{Sun}}$ of material to Star A.

Star A: $1.2M_{\text{Sun}} \rightarrow 3.4M_{\text{Sun}}$

Star B: $3.0M_{\text{Sun}} \rightarrow 0.8M_{\text{Sun}}$

Mass Transfer Binaries

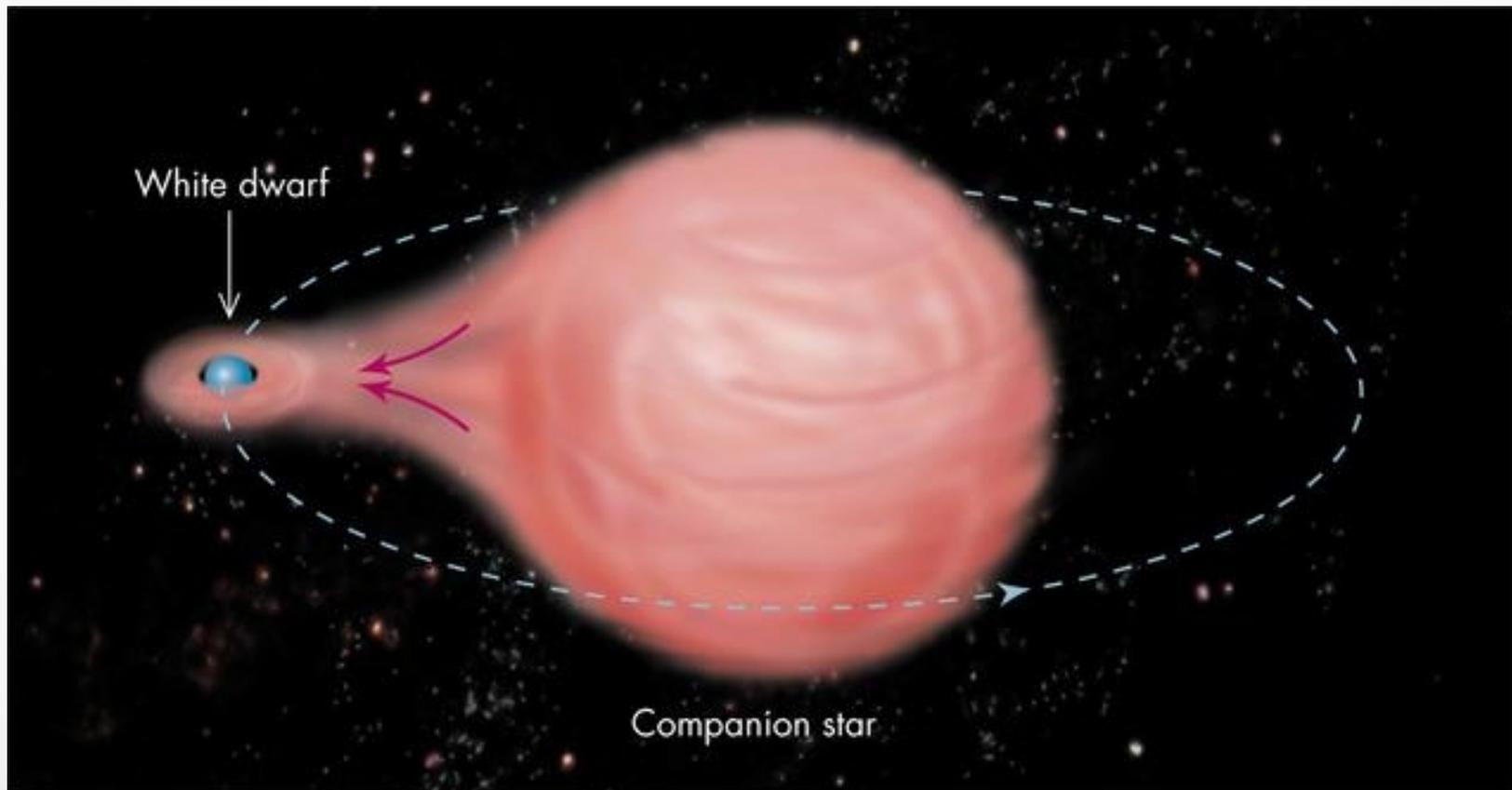
- Think about the continued evolution of Algol and you have the explanation for novae.
- If the original primary transfers most of its mass to the original secondary, you are left with a massive main-sequence star and a helium WD.
- When the original secondary starts to evolve up the RGB, it transfers some material back onto the helium WD.



Novae

- As the fresh hydrogen accumulates on the surface of the helium WD it is like an insulating blanket -- the temperature rises to 10^7 k and there is a *Hydrogen fusion explosion*.
- The star brightens by anywhere from a factor of 10 to a factor of 10,000.
- In some cases, this takes a star from too-faint to see to bright-enough to see so these objects were called Nova -- new star.

Novae/Supernovae I



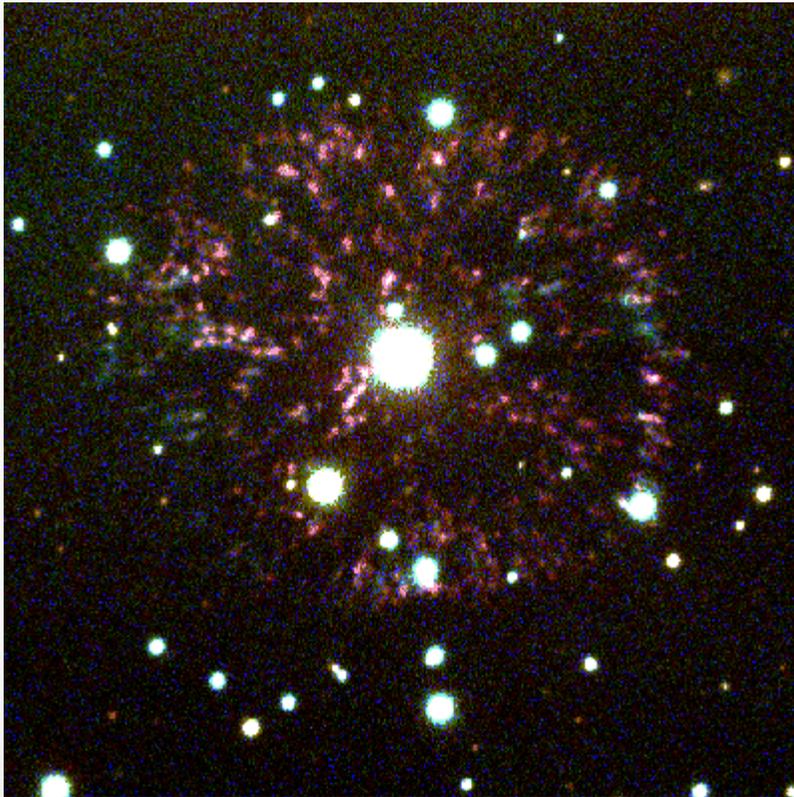
Note! Not to scale!

Novae



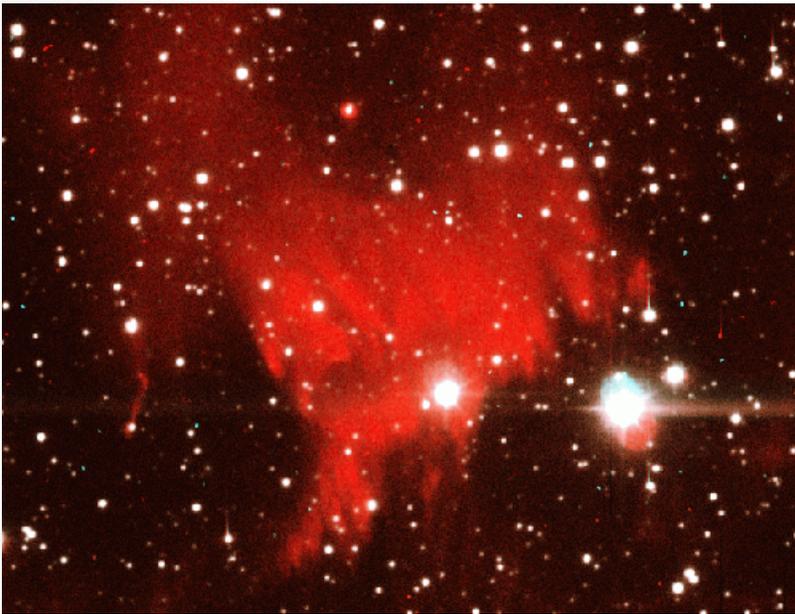
- Nova Vel 1998 (3rd magnitude)

Novae



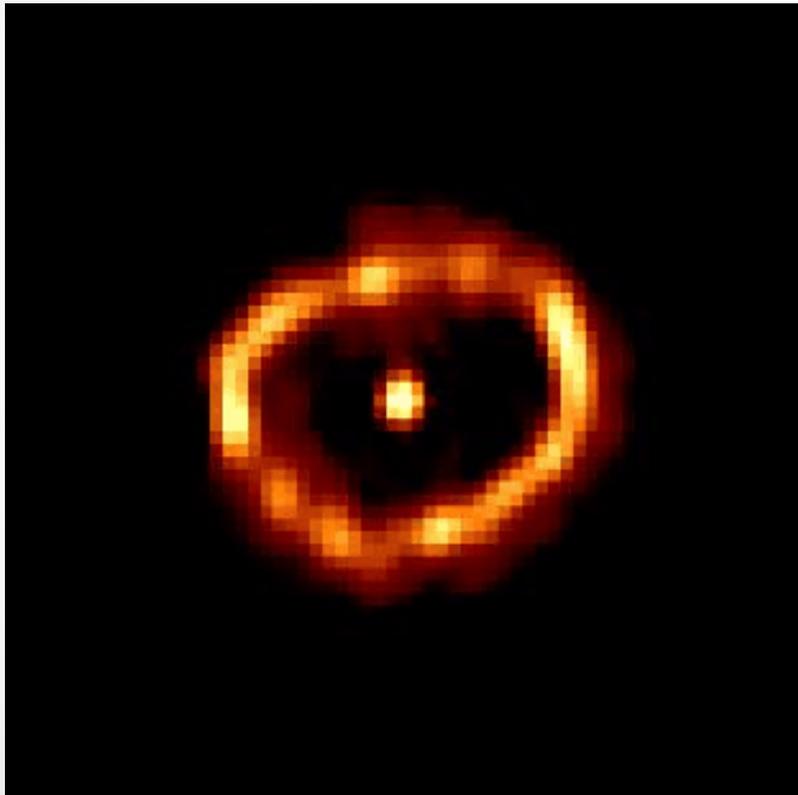
- Nova Persei became one of the brightest stars in the sky in 1901. Look there now and see the expanding shell from the explosion. The velocity of the material is $\sim 2000\text{km/sec}$

Novae



- Nova Cyg (1992) illuminated a cloud of nearby Hydrogen gas.
- The expanding shell of the nova could be seen a few years later with HST.

Novae



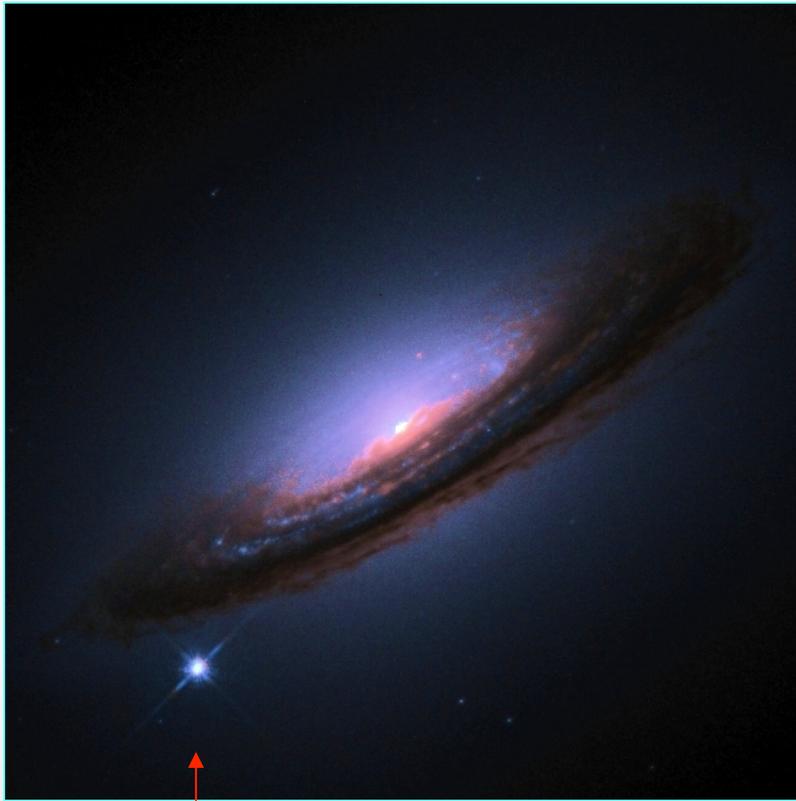
- Nova Cyg in 1994.
- Most nova are “recurrent”.
- Every year there are 20 - 30 novae observed in the Galaxy. Naked eye nova occur more like one per decade.

Mass Transfer in Binaries

- The scenario that leads to nova explosions can produce an even wilder phenomenon.
- In the early 1900s `novae' were sometimes observed in other galaxies and were used to help set the distances to galaxies.
- But, when it became clear that even the nearest galaxies were much further away than anyone had thought this suggested that the extragalactic “nova” were much brighter than Galactic nova -- the term supernova was coined.

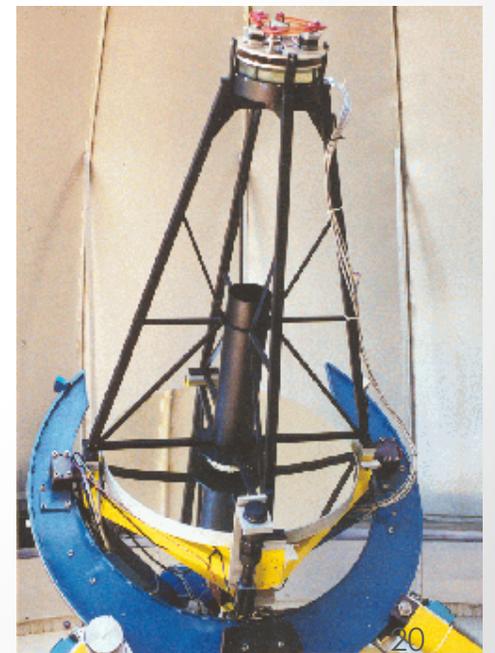
Supernova

- Early on it was realized there were two distinct types of SN.
- SN I have no hydrogen in their spectra and are seen in all types of galaxies
- SN II have hydrogen and are only seen in spiral galaxies and near star-forming regions (discussed in the last section)



Supernova

- There is a robotic telescope up at Mt. Hamilton that does an automatic search for SN every clear night.
- Take images of lots of galaxies, digitally subtract them, look for any residual.



Supernova I

- What is going on here? It took a long time to sort this out.
- Remember WD mass transfer binaries and the Chandrasekar limit.
- What would happen if mass transfer nudged the mass of a WD above the $1.4M_{\text{Sun}}$ limit for degenerate electron gas pressure?

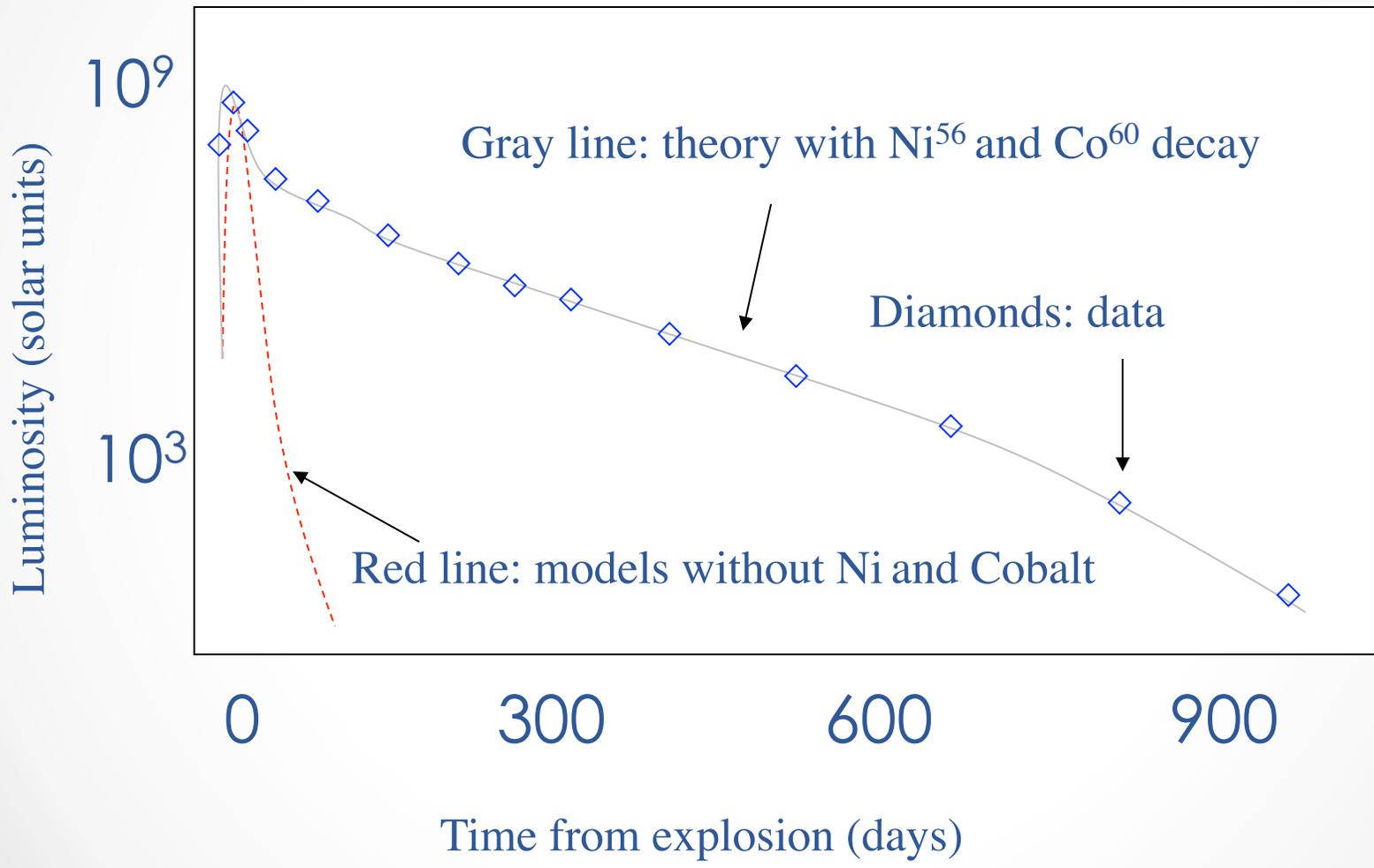
Supernova I

- When a WD exceeds the Chandrasekar limit the electron degeneracy pressure can not support it against gravity and the WD collapses.
- The temperature skyrockets and within a second a fusion chain reaction fuses elements all the up to radioactive nickel.
- *This is a runaway thermonuclear explosion!*
- Different that the core collapse in a high-mass star Fe core (also triggered at $\sim 1.4M_{\text{Sun}}$) because the WD is primarily made of He.

Supernova I



- What is RIGHT about this theory?
 - (1) Will see these objects in `old' populations.
 - (2) Models for the detonation of a $1.4M_{\odot}$ WD give the right total energy
 - (3) The predicted amount of radioactive Ni^{56} in the explosion fit the light curve perfectly



Supernova I

- Because the explosion mechanism for SN I is the same from one to the next (WD with mass transfer that pushes it over the Chandrasekar mass), these explosions are pretty good “standard candles” and will play a roll in establishing the expansion history of the Universe.

Historical Galactic SN

- We miss many in the Galaxy because of dust obscuration.
- From radio surveys for SN remnants, we have discovered 49 remnants for an inferred rate of 3.4 SN/century.
- There are several 'historical supernovae' -- bright new stars that appeared in the sky and were recorded by various people.

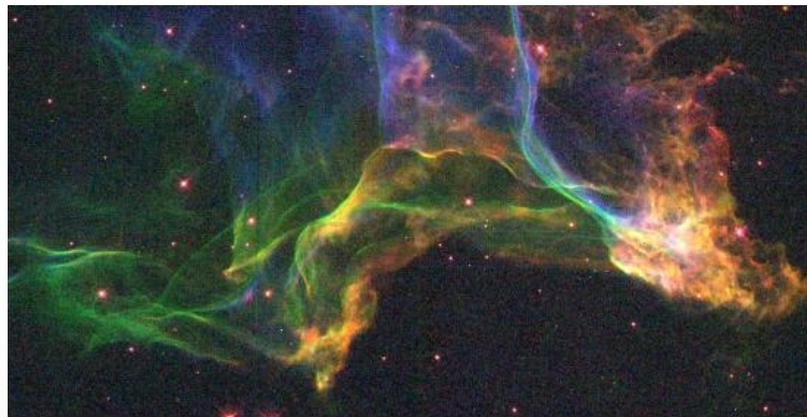
Historical SN

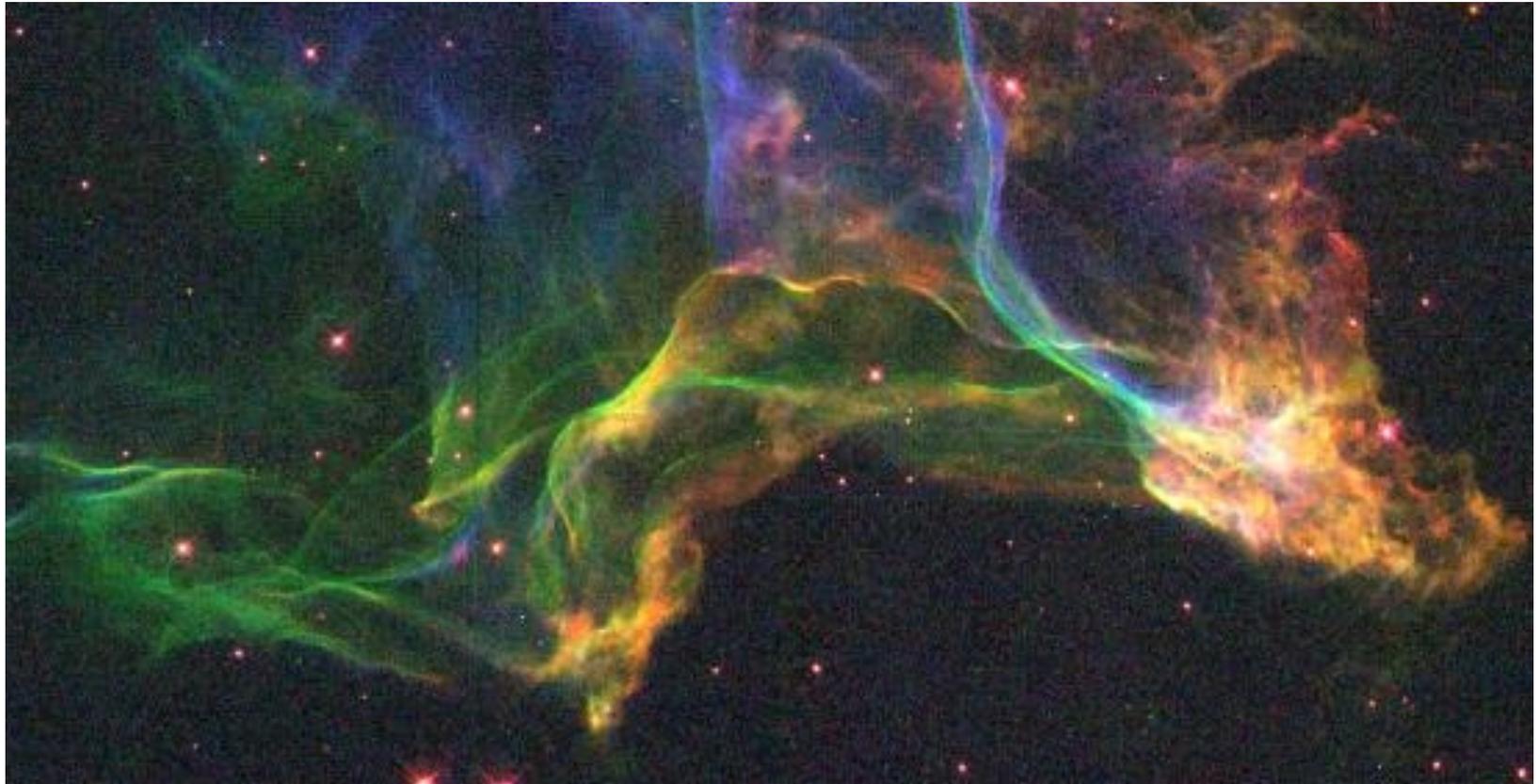
- 1006, 1054, 1181, 1572, 1604 and 1658 were years when bright `guest stars` were widely reported

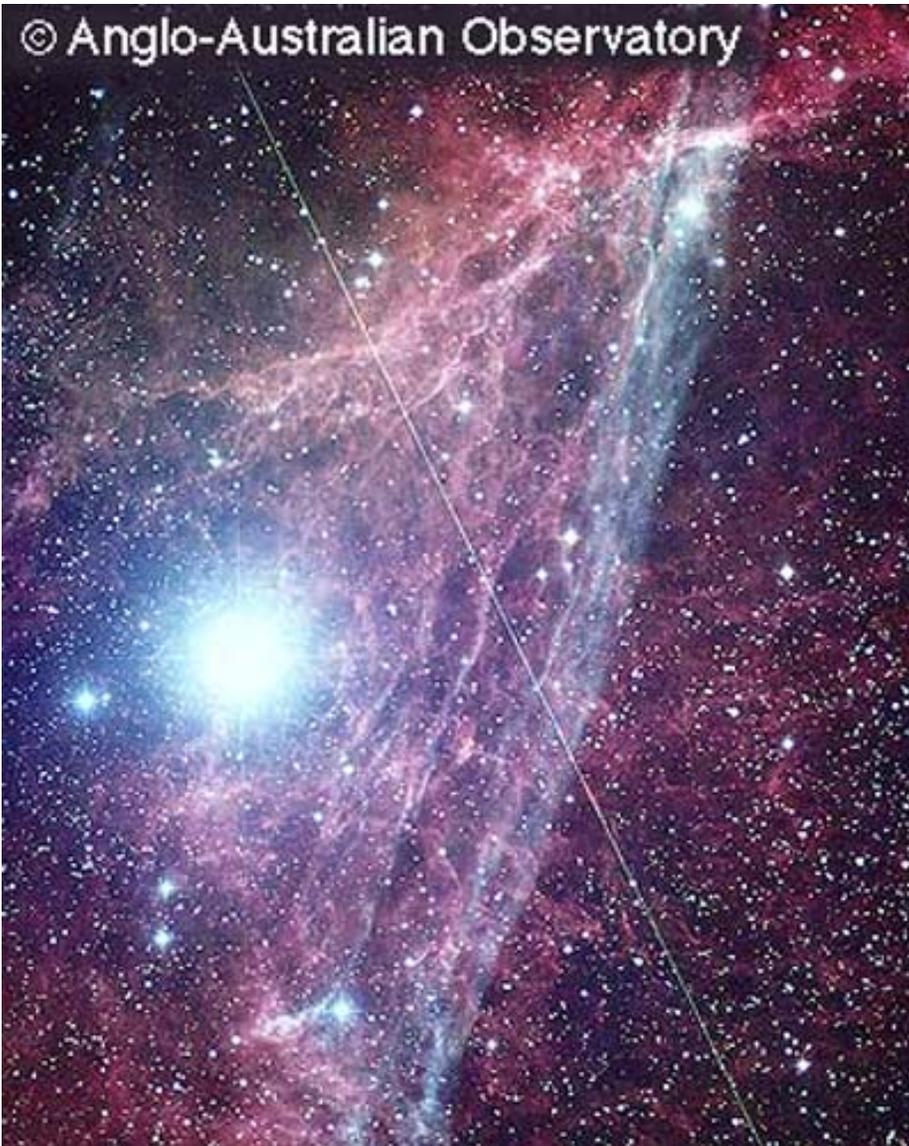


Historical SN

- For all the guest stars, point a modern telescope at the position and see a rapidly-expanding shell of material.
- In two cases, the remnant was discovered first







NGC 1850 in LMC



Historical SN

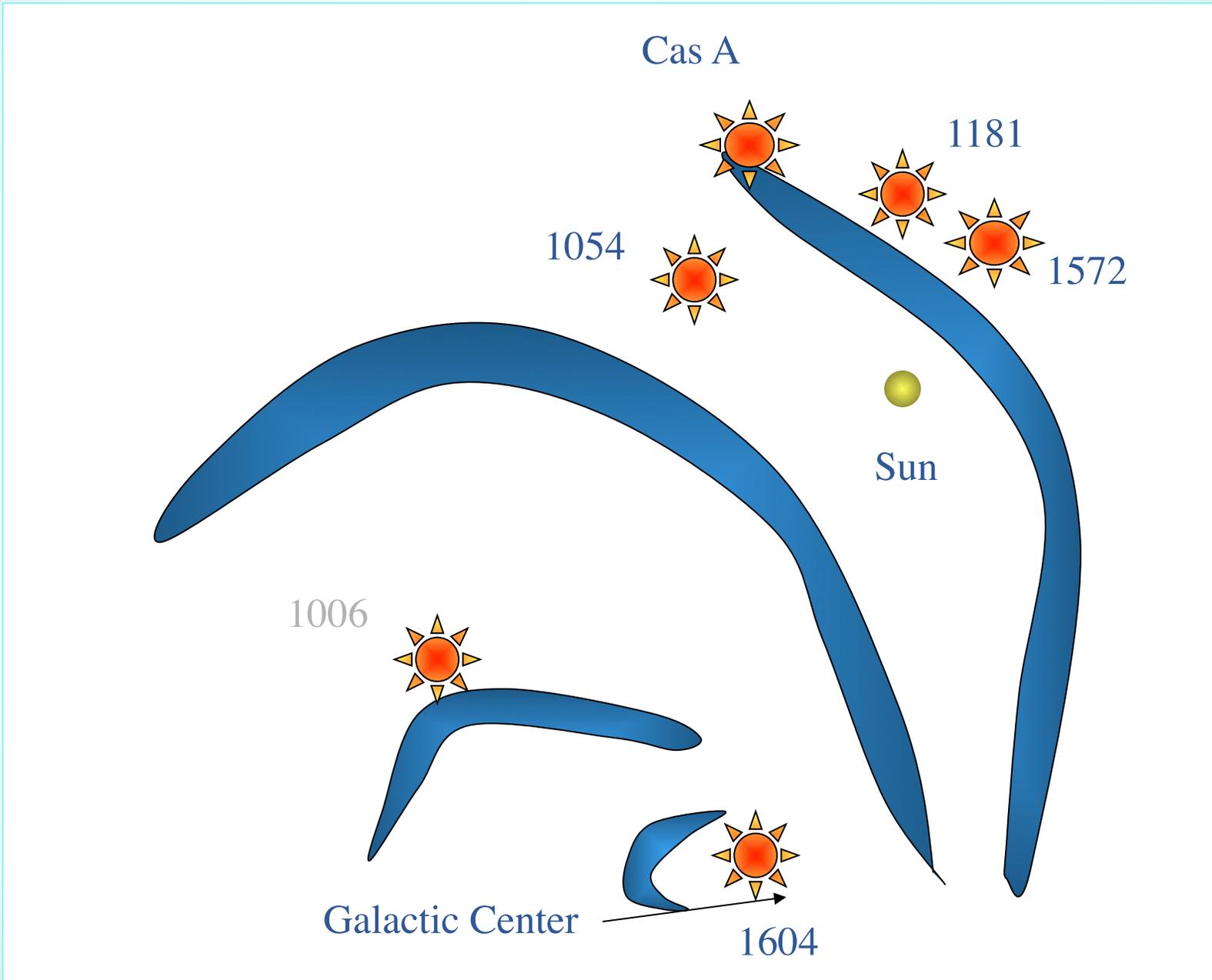
- The 1054AD event was so bright it cast shadows during the day -- this is seen today at the position of the Crab Nebula





Historical SN

- The nearest SN remnant is the 'Gum' nebula from around 9000BC. Four times closer than the Crab, it would have been as bright as the full moon.
- A mystery is 'Cas A' -- this was a SN at about 1600AD, should have been very bright, but no records of it exist.

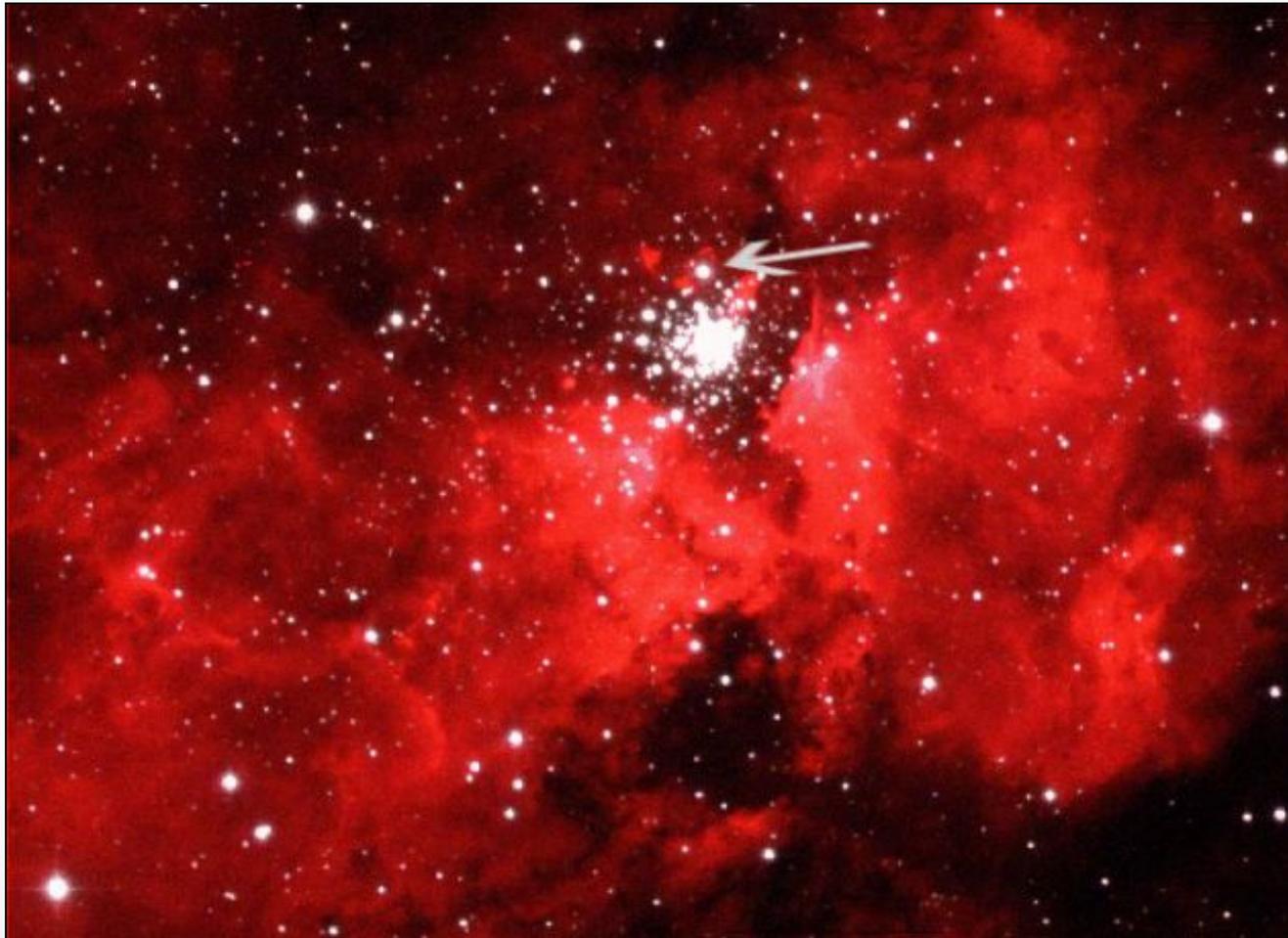


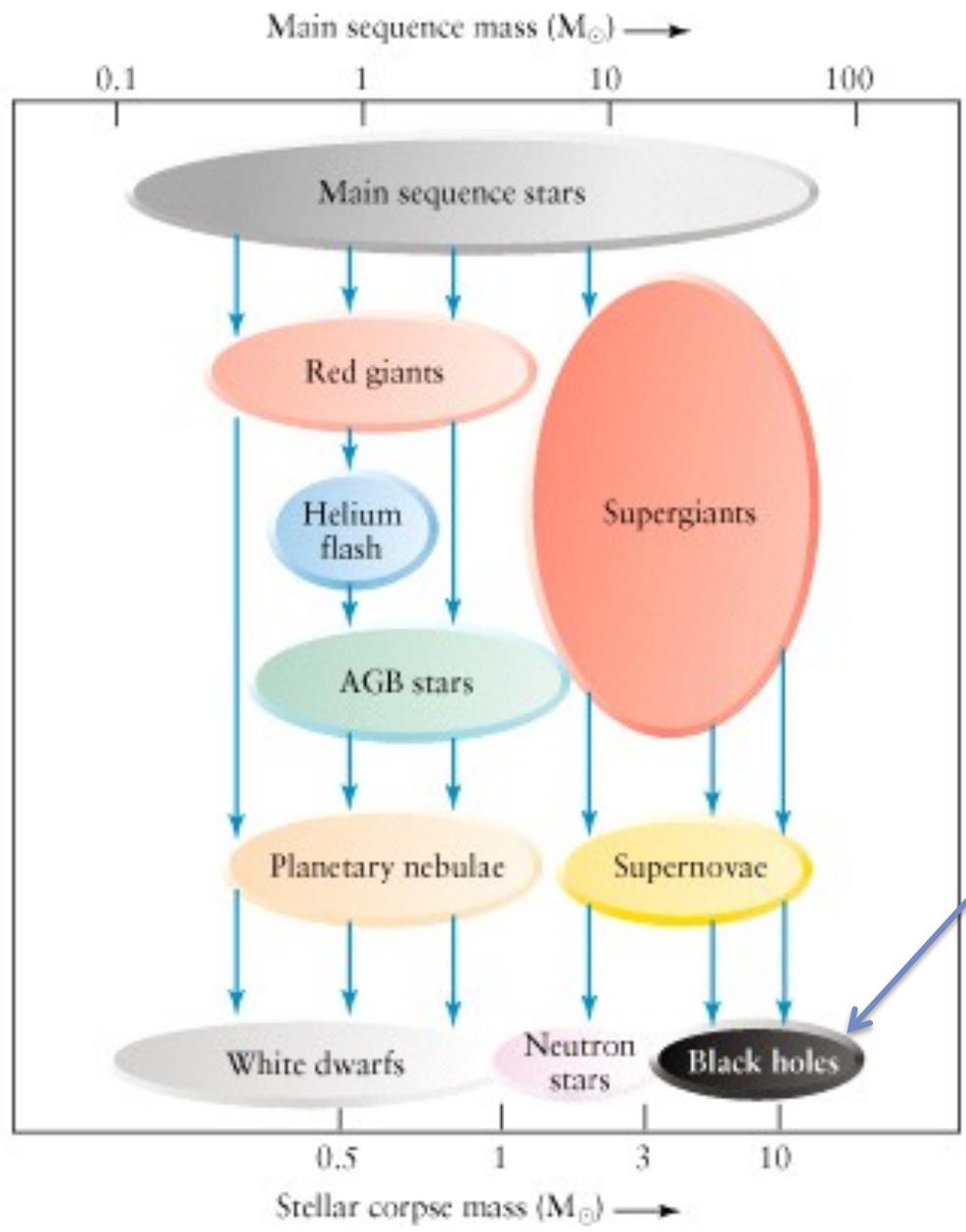
Supernovae in the Galaxy

- We are long overdue for a bright Galactic Supernova.
- For a while, a nearby SN was a valid candidate for the source of the demise of the dinosaurs.
- There are the products of short-lived radioactive isotopes locked up in primitive meteorites which suggest a SN in the vicinity of the Solar System about 100,000 years before the Sun formed. A SN may have triggered the collapse of the proto-Sun.



Next Galactic SN?





Next section

Chemical Evolution Of the Universe

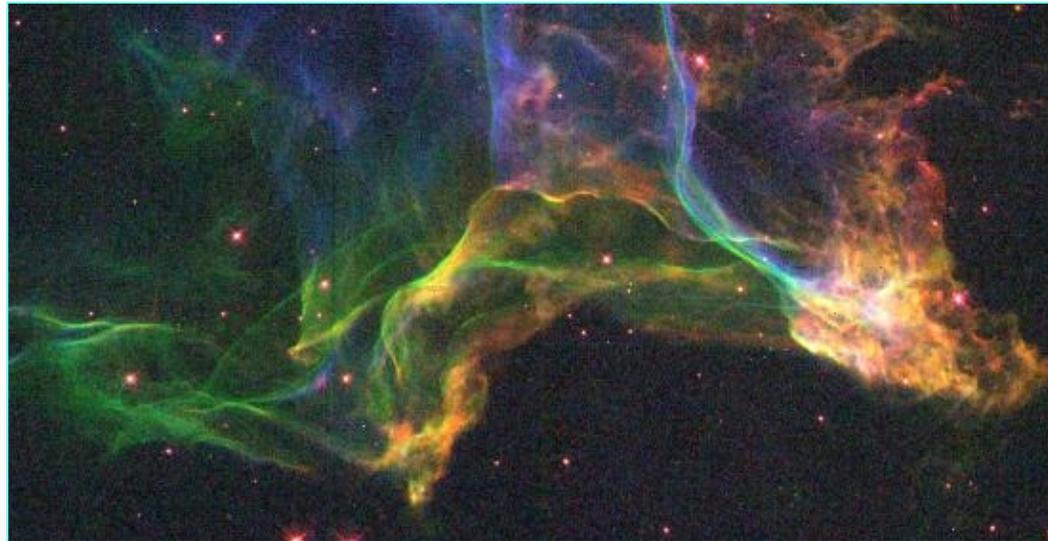
- Low-mass stars synthesize “new” He, C, O, Ne, Ar, Mg during the main-sequence and giant-branch phases.
- These freshly-minted elements are brought to the surface via convection and re-distributed via stellar winds and planetary nebulae into the interstellar medium to be incorporated into later generations of stars.



© Royal Observatory, Edinburgh / Anglo-Australian Observatory

Chemical Evolution II

- For more massive stars, 'equilibrium' fusion reactions produce elements all the way up to Fe.
- Freshly made elements are delivered via stellar winds or, sometimes more spectacularly via supernova explosions



PERIODIC TABLE OF THE ELEMENTS

1 IA																	18 VIIIA	
1 H Hydrogen 1.0079	2 IIA		<div style="border: 1px solid black; padding: 5px; width: fit-content; margin: auto;"> <p style="text-align: center;">14 ← Group IUPAC IVA ← Group CAS</p> <p>Atomic Number → 6 -4 Symbol → C +2 Name → Carbon +4 Electron Configuration → 2-4 Selected Oxidation States Atomic Mass → 12.011 Atomic Mass</p> </div>										13 IIIA	14 IVA	15 VA	16 VIA	17 VIIA	He Helium 4.0026
2 Li Lithium 6.941	Be Beryllium 9.0122											B Boron 10.811	C Carbon 12.011	N Nitrogen 14.007	O Oxygen 15.999	F Fluorine 18.998	Ne Neon 20.179	
3 Na Sodium 22.990	Mg Magnesium 24.305	3 IIIB	4 IVB	5 VB	6 VIB	7 VIIB	8 VIII	9 VIII	10 VIII	11 IB	12 IIB	Al Aluminium 26.982	Si Silicon 28.086	P Phosphorus 30.974	S Sulphur 32.065	Cl Chlorine 35.453	Ar Argon 39.948	
4 K Potassium 39.098	Ca Calcium 40.078	Sc Scandium 44.956	Ti Titanium 47.867	V Vanadium 50.942	Cr Chromium 51.996	Mn Manganese 54.938	Fe Iron 55.845	Co Cobalt 58.933	Ni Nickel 58.693	Cu Copper 63.546	Zn Zinc 65.39	Ga Gallium 69.723	Ge Germanium 72.64	As Arsenic 74.922	Se Selenium 78.96	Br Bromine 79.904	Kr Krypton 83.80	
5 Rb Rubidium 85.468	Sr Strontium 87.62	Y Yttrium 88.906	Zr Zirconium 91.224	Nb Niobium 92.906	Mo Molybdenum 95.94	Tc Technetium (98)	Ru Ruthenium 101.07	Rh Rhodium 102.91	Pd Palladium 106.42	Ag Silver 107.87	Cd Cadmium 112.41	In Indium 114.82	Sn Tin 118.71	Sb Antimony 121.76	Te Tellurium 127.60	I Iodine 126.90	Xe Xenon 131.29	
6 Cs Cesium 132.91	Ba Barium 137.33	La Lanthanide	Hf Hafnium 178.49	Ta Tantalum 180.95	W Tungsten 183.84	Re Rhenium 186.21	Os Osmium 190.23	Ir Iridium 192.22	Pt Platinum 195.08	Au Gold 196.97	Hg Mercury 200.59	Tl Thallium 204.38	Pb Lead 207.2	Bi Bismuth 208.98	Po Polonium (209)	At Astatine (210)	Rn Radon (222)	
7 Fr Francium (223)	Ra Radium (226)	Ac Actinide	Rf Rutherfordium (261)	Db Dubnium (262)	Sg Seaborgium (266)	Bh Bohrium (264)	Hs Hassium (277)	Mt Meitnerium (268)	Uun Ununnilium (281)	Uuu Ununquadium (283)	Uub Ununhexium (285)	Uut Ununseptium (284)	Uuq Ununquadium (289)	Uup Ununpentium (288)	Uuh Ununhexium (291)	Uus Ununseptium (291)	Uuo Ununoctium (294)	

Electron Shells

1	K	2	S	P	D	F
2	L	8	2	6		
3	M	18	2	6	10	
4	N	32	2	6	10	14
5	O	32	2	6	10	14
6	P	18	2	6	10	
7	Q	8	2	6		
8	R	2	2			

Lanthanide

57 La Lanthanum 138.91 2-8-18-18-9-2	58 Ce Cerium 140.12 2-8-18-20-8-2	59 Pr Praseodymium 140.91 2-8-18-21-8-2	60 Nd Neodymium 144.24 2-8-18-22-8-2	61 Pm Promethium (145) 2-8-18-23-8-2	62 Sm Samarium 150.36 2-8-18-24-8-2	63 Eu Europium 151.96 2-8-18-25-8-2	64 Gd Gadolinium 157.25 2-8-18-25-9-2	65 Tb Terbium 158.93 2-8-18-27-8-2	66 Dy Dysprosium 162.50 2-8-18-28-8-2	67 Ho Holmium 164.93 2-8-18-29-8-2	68 Er Erbium 167.26 2-8-18-30-8-2	69 Tm Thulium 168.93 2-8-18-31-8-2	70 Yb Ytterbium 173.04 2-8-18-32-8-2	71 Lu Lutetium 174.97 2-8-18-32-9-2
--	---	---	--	--	---	---	---	--	---	--	---	--	--	---

Actinide

89 Ac Actinium (227) -18-32-18-9-2	90 Th Thorium 232.04 -18-32-18-10-2	91 Pa Protactinium 231.04 -18-32-20-9-2	92 U Uranium 238.03 -18-32-21-9-2	93 Np Neptunium (237) -18-32-23-8-2	94 Pu Plutonium (244) -18-32-24-8-2	95 Am Americium (243) -18-32-25-8-2	96 Cm Curium (247) -18-32-25-9-2	97 Bk Berkelium (247) -18-32-27-8-2	98 Cf Californium (251) -18-32-28-8-2	99 Es Einsteinium (252) -18-32-29-8-2	100 Fm Fermium (257) -18-32-30-8-2	101 Md Mendelevium (258) -18-32-31-8-2	102 No Nobelium (259) -18-32-32-8-2	103 Lr Lawrencium (262) -18-32-32-9-2
--	---	---	---	---	---	---	--	---	---	---	--	--	---	---

Chemical Evolution III

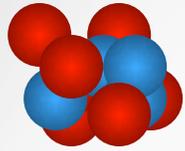
- What about the *trans-Fe* elements?
- Equilibrium fusion reactions of light elements don't proceed past Fe because of Fe's location at the peak of the curve of binding energy.
- However, in certain circumstances, supernovae for example, non-equilibrium reactions can build elements beyond Fe in the Periodic Table. Many of these are radioactive, but some are stable.



Neutron Capture Elements

There are two principle paths to building the elements heavier than Fe. Both use the addition of neutrons to existing 'seed' nuclei (neutrons have no charge so are much easier to add to positively-charged nuclei).

- S-process (slow addition of neutrons)
- R-process (rapid addition of neutrons)

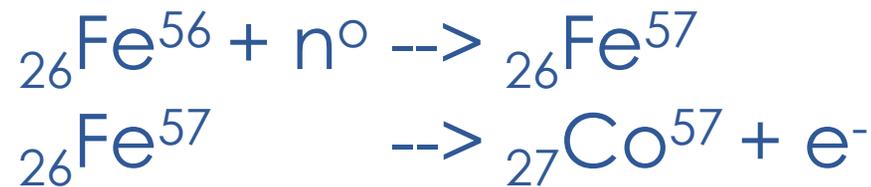


The S-process

- The S-process stands for the Slow addition of neutrons to nuclei. The addition of a n^0 produces heavier isotope of a particular element. However, if an electron is emitted (this is called beta-decay), the nucleus moves one step up the periodic table.

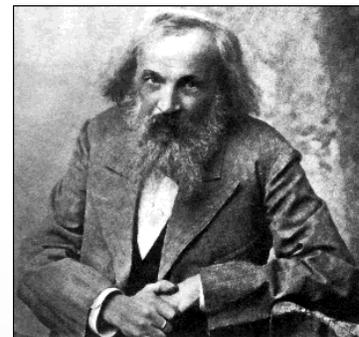
S-Process

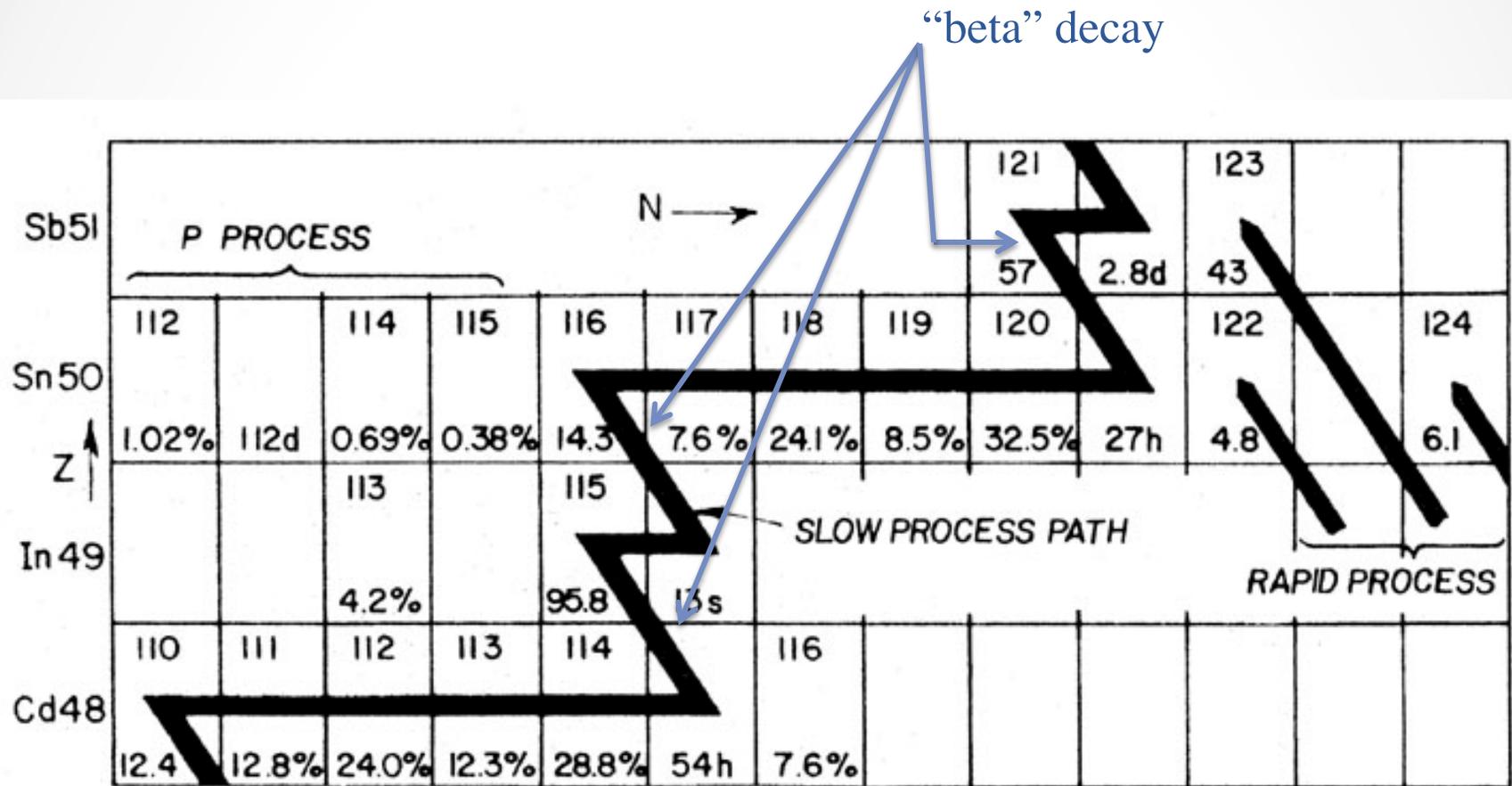
- `Slow' here means that rate of n° captures is low compared to the beta-decay rate.
- It really is slow, sometimes 100's of years go by between neutron captures.



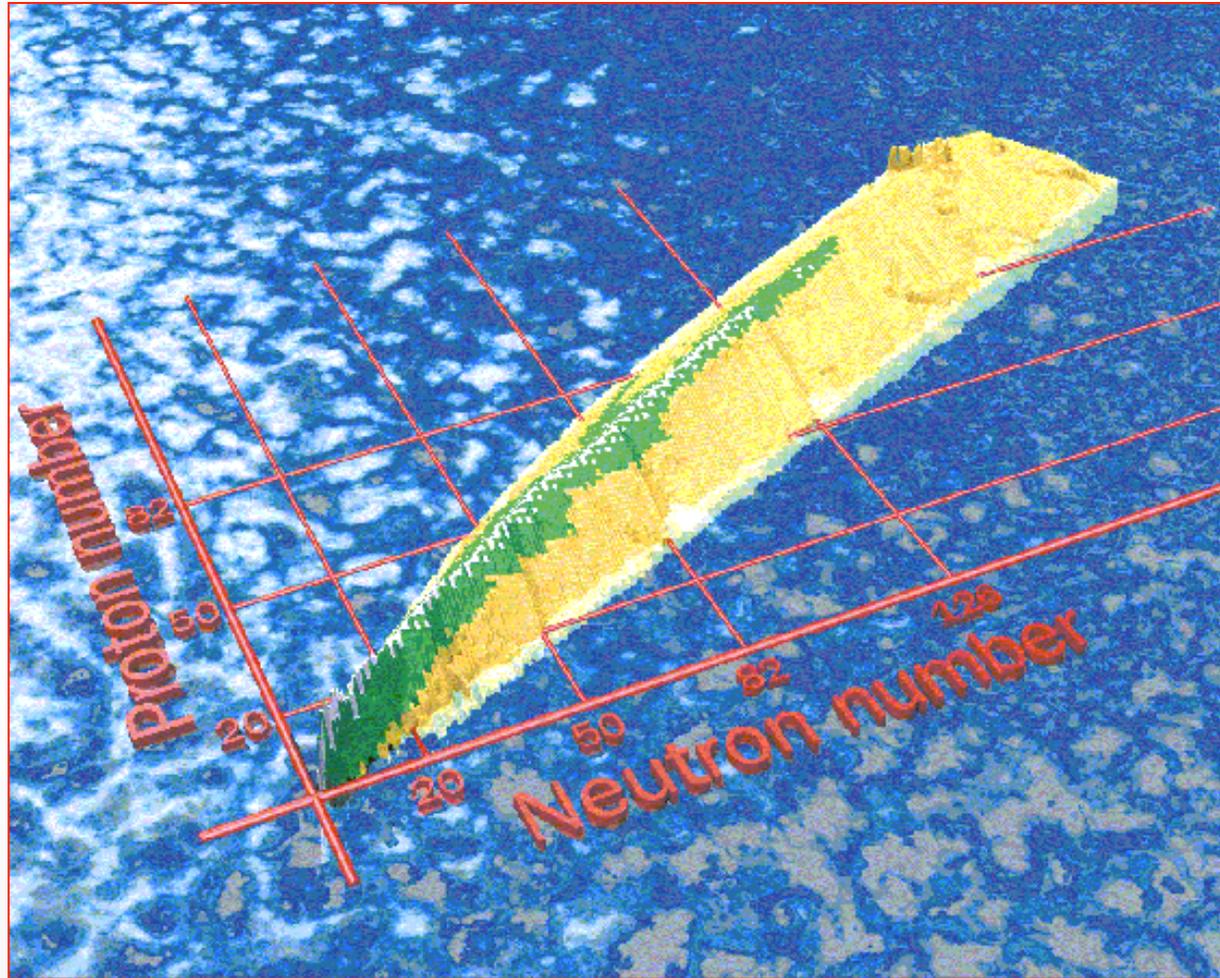
Here a neutron changed into a proton by emitting an electron

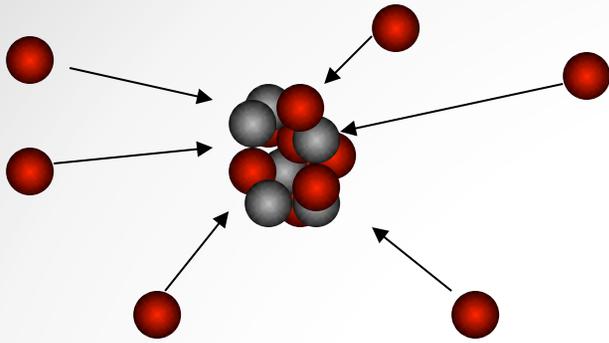
- The S-process can produce elements up to #83 - Bismuth. There are peaks in the Solar System abundance of heavy elements at ^{38}Sr , ^{56}Ba and ^{82}Pb . These are easily understood in the context of the S-process and 'magic' numbers of neutrons.
- The site of the S-process is AGB stars during and between shell flashes. The n° source is a by-product of $\text{C}^{13} + \text{He}^4 \rightarrow \text{O}^{16}$
- ^{43}Tc is an s-process nucleus and proof that it is in operation in AGB stars.





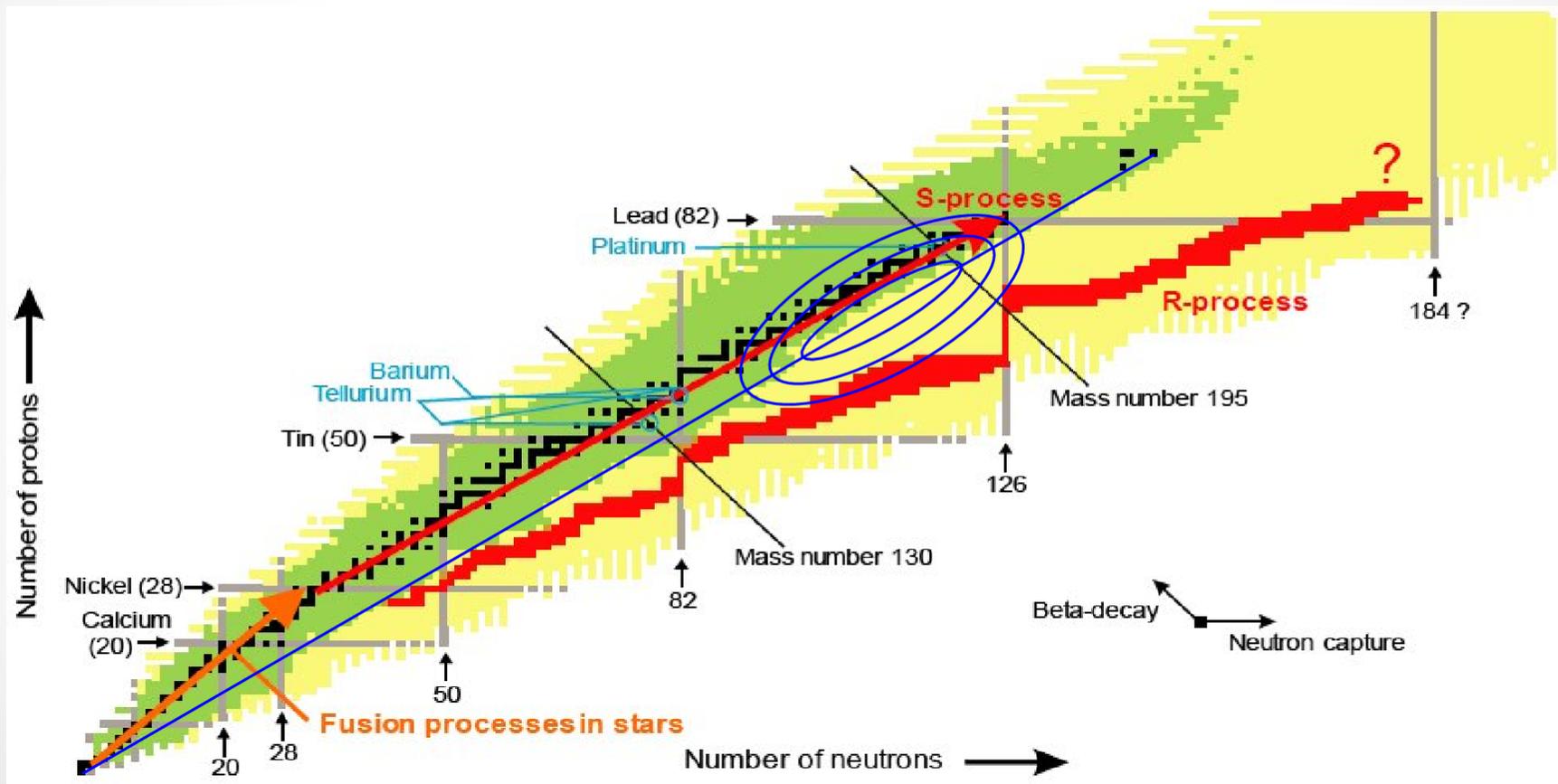
Number of neutrons





The R-process

- The R-process is the Rapid addition of neutrons to existing nuclei. Rapid here means that many neutrons are added before a beta-decay occurs.
- First build up a VERY heavy isotope, then as beta-decays occur you march up in atomic number and produce the REALLY HEAVY STUFF.



The R-process

- For this to happen need a big burst of neutrons. The most promising place with the right conditions is in a SNII explosion right above the collapsed core.
- We see an overabundance of R-process elements in the oldest stars. As the early chemical enrichment of the Galaxy was through SNII, this is evidence of SNII as the source of r-process elements

R-process



- If we look at the Crab Nebula or other SNIa remnants we don't see r-process elements.
- We DO see regions of enhanced O, Si, Ne and He which appear to reflect the 'onion skin' structure of the massive star progenitor.

Solar Composition by Mass

H	78.4%	Big Bang	Low-mass stars
He.....	19.8%		
O	0.8%	High-mass stars	
C	0.3%		
N	0.2%		
Ne	0.2%		
Si	0.04%		
Fe	0.04%		
Gold.....	0.000000009%		
	(\$2.1 x 10 ²⁴ at \$300/ounce)		

We are star dust

- This is literally true



The Origin of the Solar System Elements

1 H	big bang fusion 					cosmic ray fission 					2 He						
3 Li	4 Be	merging neutron stars 					exploding massive stars 					5 B	6 C	7 N	8 O	9 F	10 Ne
11 Na	12 Mg	dying low mass stars 					exploding white dwarfs 					13 Al	14 Si	15 P	16 S	17 Cl	18 Ar
19 K	20 Ca	21 Sc	22 Ti	23 V	24 Cr	25 Mn	26 Fe	27 Co	28 Ni	29 Cu	30 Zn	31 Ga	32 Ge	33 As	34 Se	35 Br	36 Kr
37 Rb	38 Sr	39 Y	40 Zr	41 Nb	42 Mo	43 Tc	44 Ru	45 Rh	46 Pd	47 Ag	48 Cd	49 In	50 Sn	51 Sb	52 Te	53 I	54 Xe
55 Cs	56 Ba	72 Hf	73 Ta	74 W	75 Re	76 Os	77 Ir	78 Pt	79 Au	80 Hg	81 Tl	82 Pb	83 Bi	84 Po	85 At	86 Rn	
87 Fr	88 Ra																
		57 La	58 Ce	59 Pr	60 Nd	61 Pm	62 Sm	63 Eu	64 Gd	65 Tb	66 Dy	67 Ho	68 Er	69 Tm	70 Yb	71 Lu	
		89 Ac	90 Th	91 Pa	92 U												