Quantum Mechanics and Stellar Spectroscopy

http://apod.nasa.gov/apod/

The Electrical Force

Recall the electric force. Like gravity it is a “$1/r^2$” force/ That is:

$e = 4.803 \times 10^{-10}$ *esu*

$e^2 = 2.307 \times 10^{-19}$ dyne cm^2

$$\boxed{F_{elec} = \frac{Z_1 Z_2 e^2}{r^2}}$$

$Z_1 \leftrightarrow m$

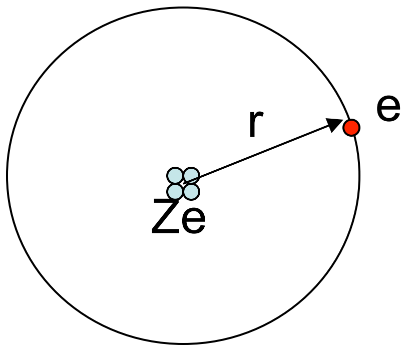
$Z_2 \leftrightarrow M$

$e^2 \leftrightarrow G$

where Z_1 and Z_2 are the (integer) numbers of electronic charges. Similarly, the electric potential energy is

$$\boxed{E_{elec} = -\frac{Z_1 Z_2 e^2}{r}}$$

Rutherford Atom (1911)



$$F_{elec} = F_{cent}$$

$$\frac{Ze^2}{r^2} = \frac{m_e v^2}{r} \Rightarrow \boxed{r = \frac{Ze^2}{m_e v^2}}$$

for a single electron

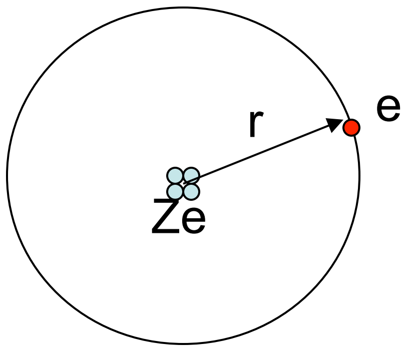
Z = 1,2,3,…
H, He, Li, etc

$$v = \sqrt{\frac{Ze^2}{rm_e}}$$

Protons in nucleus. Electrons orbit like planets. The neutron was not discovered until 1932 (Chadwick)

classically, any value of v or r is allowed. Much like planets.

Rutherford Atom (1911)



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$$F_{elec} = F_{cent}$$

$$\frac{Ze^2}{r^2} = \frac{m_e v^2}{r} \Rightarrow \boxed{r = \frac{Ze^2}{m_e v^2}}$$

Z = 1,2,3,…

$$v = \sqrt{\frac{Ze^2}{r m_e}}$$

classically, any value of v or r is allowed.

Total energy:

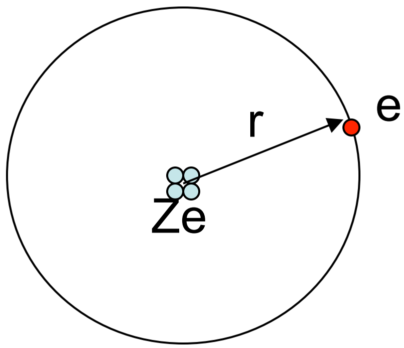
$$E_{tot} = KE + PE = \frac{m_e v^2}{2} - \frac{Ze^2}{r}$$

$$= \frac{Ze^2}{2r} - \frac{Ze^2}{r} = -\frac{Ze^2}{2r}$$

i.e., $\boxed{2KE = -PE}$ (if PE is negative)

Virial theorem still works for the electric force.

Rutherford Atom (1911)



$$E_{tot} = -\frac{Ze^2}{2r}$$

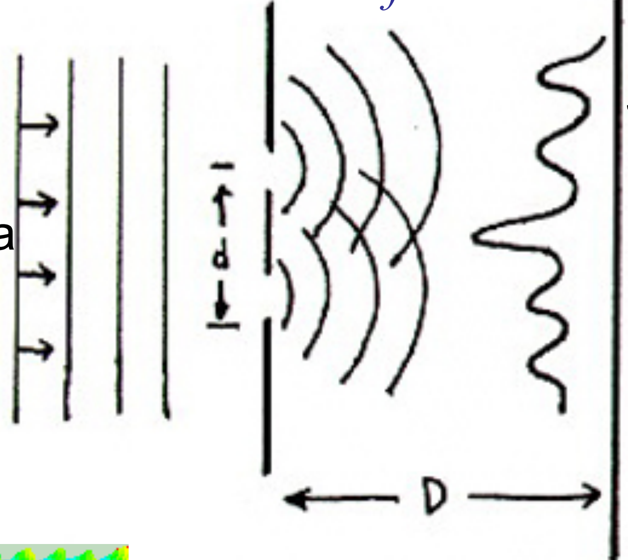
BUT,

As the electron moves in its classical orbit it is accelerated, and therefore emits radiation. Because energy is being radiated, the total energy of the system must decrease – become more negative. This means *r* must get smaller and *v* must increase. But smaller *r* and larger *v* also imply greater acceleration and radiation.

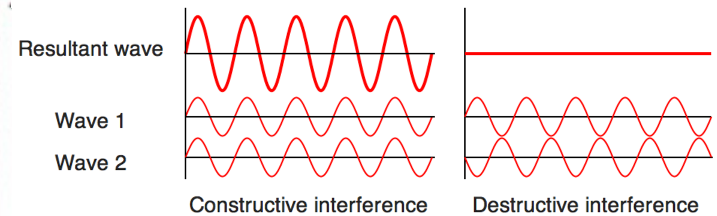
In approximately 10^{-6} s the electron spirals into the nucleus. Goodbye universe…

The solution lies in the wave-like property of the electron – and of all matter

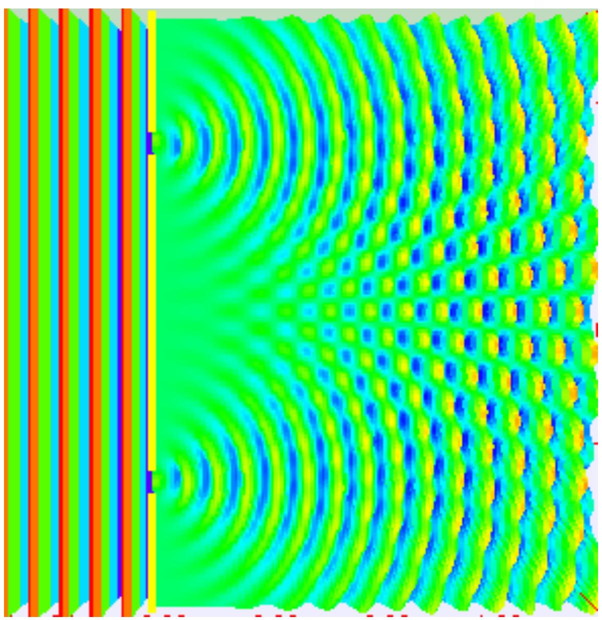
For wavelike phenomena e.g., light, “interference” is expected

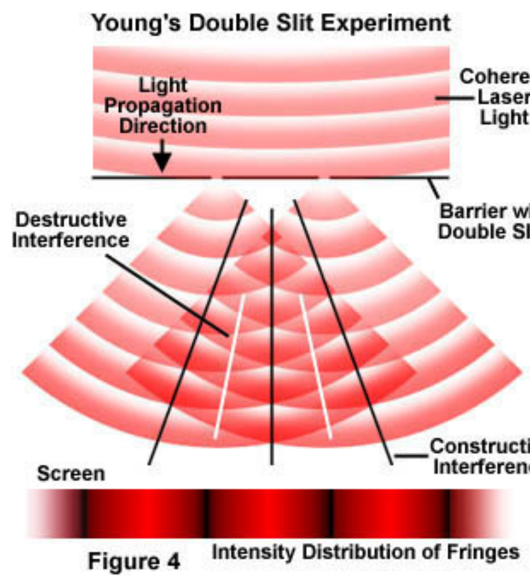


$$spacing = D\frac{\lambda}{d}$$

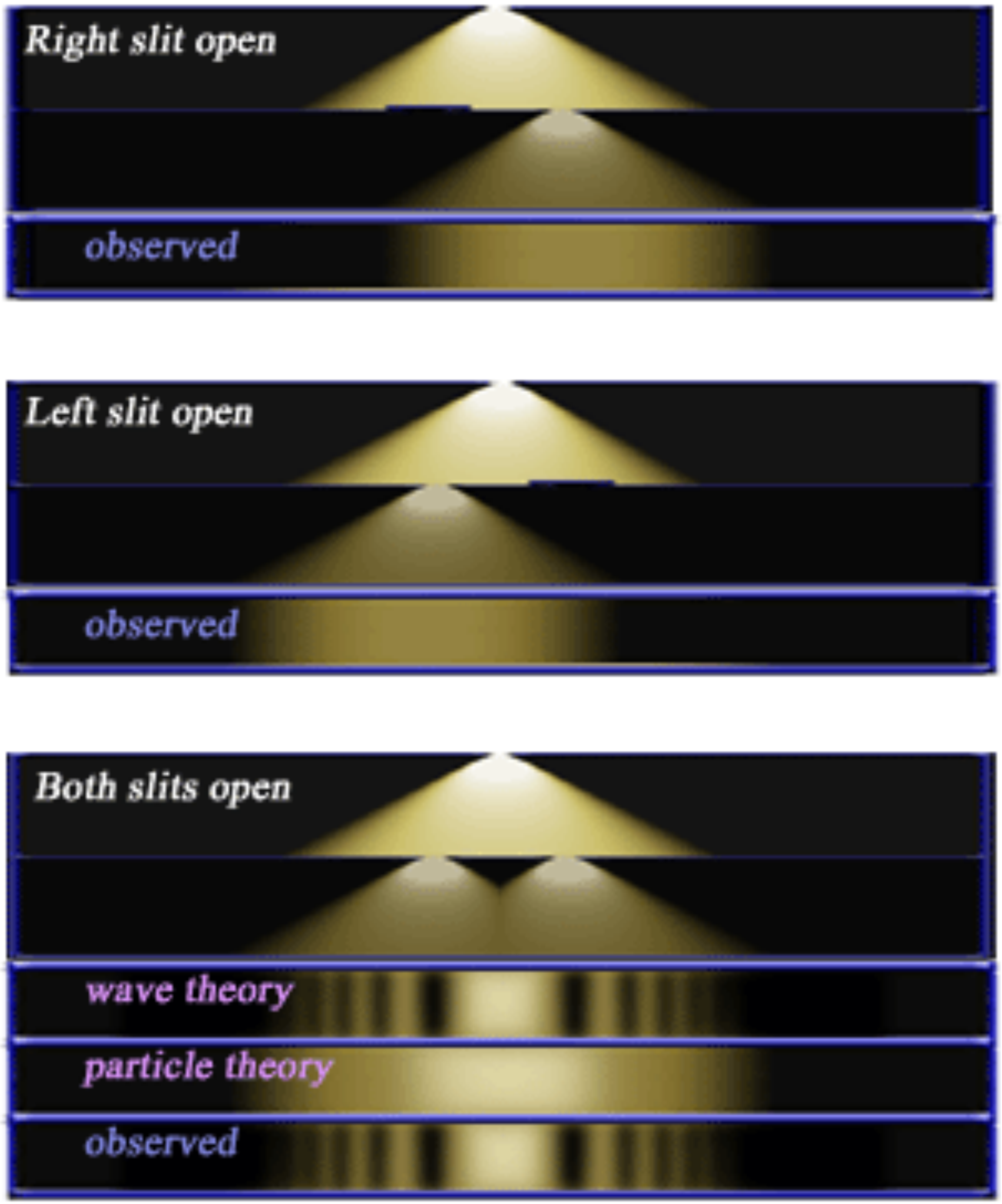


Thomas Young early 1800’s for light.

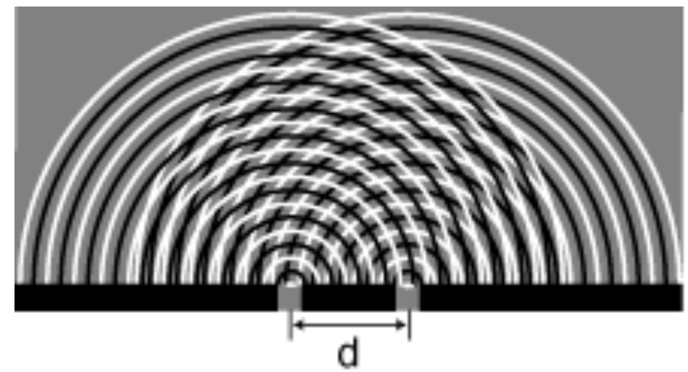


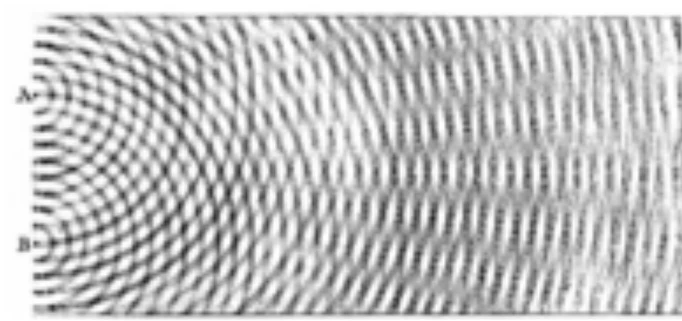


http://en.wikipedia.org/wiki/Interference_(wave_propagation) http://en.wikipedia.org/wiki/Double-slit_experiment



Young's experiment





Same basic result obtained using electrons!

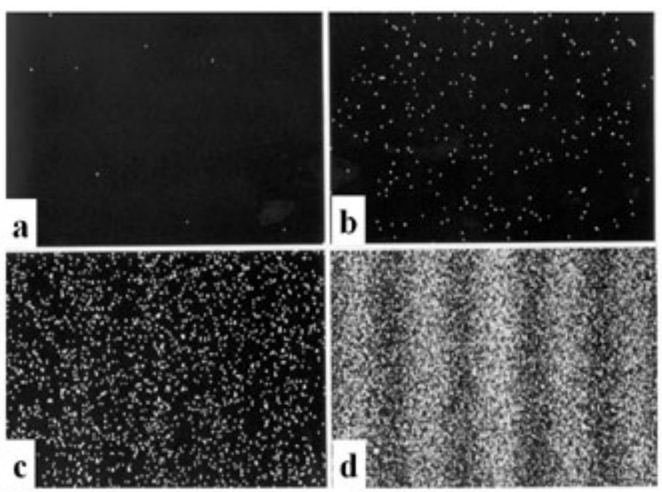
$$\lambda = \frac{h}{p}$$

$$h = 6.626 \times 10^{-27} \text{ erg sec}$$

$$= 6.626 \times 10^{-27} \ \frac{\text{gm cm}^2}{\text{sec}}$$

where p is the momentum of the electron, $m_e v$ which has units gm cm/sec

8e⁻ 270e⁻



Single-electron Build-up of Interference Pattern

2,000e⁻ 60,000e⁻

Hitachi labs (1989)

In 1924, Louis-Victor de Broglie formulated the DeBroglie hypothesis, claiming that all matter, not just light, has a wavelike nature. He related the wavelength (denoted as λ) and the momentum (denoted as p)

$$\boxed{\lambda = \frac{h}{p}}$$

A property of our universe

This is a little like the relation Planck had for photons

$$p = \frac{h\nu}{c} = \frac{h}{\lambda}$$

$$\lambda = \frac{h}{p}$$

http://en.wikipedia.org/wiki/Wave–particle_duality

Light and particles like the electron (and neutron and proton) all have wavelengths, and the shorter the wavelength, the higher the momentum p

This is also known as the Heisenberg Uncertainty Principle. The more accurately you locate a particle (λ), the more unbounded is its momentum

HEISENBERG UNCERTAINTY RELATION

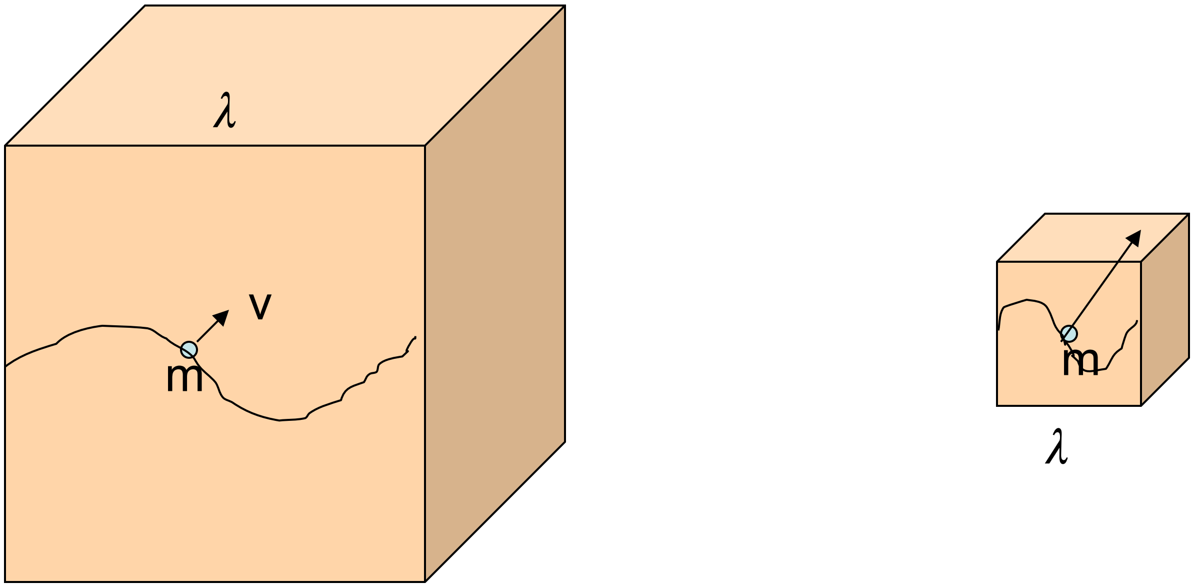
The condition that a particle cannot be localized to a region Δx smaller than its wavelength $\lambda = h/p$ implies

$$\lambda < \Delta x \;\Rightarrow\; p\,\Delta x > h \Rightarrow p > \frac{h}{\Delta x}$$

One cannot confine a particle to a region Δx without making its momentum increase

$p = \dfrac{h}{\Delta x}$ is the "degenerate" limit

Consider one electron in a contracting box



As you squeeze on the box, the particle in the box has to move faster. This is in addition to any thermal motion the particle may have

$$\boxed{\lambda = \frac{h}{p} = \frac{h}{mv}}$$

$$\lambda \downarrow \;\Rightarrow\; v \uparrow$$

The squeezing provides the energy to increase v

A little thought will show how this is going to solve our problem with the stability of matter (and also, later, the existence of white dwarfs)

As the electron is forced into a smaller and smaller volume, it must move faster. Ultimately this kinetic energy can support it against the electrical attraction of the nucleus.

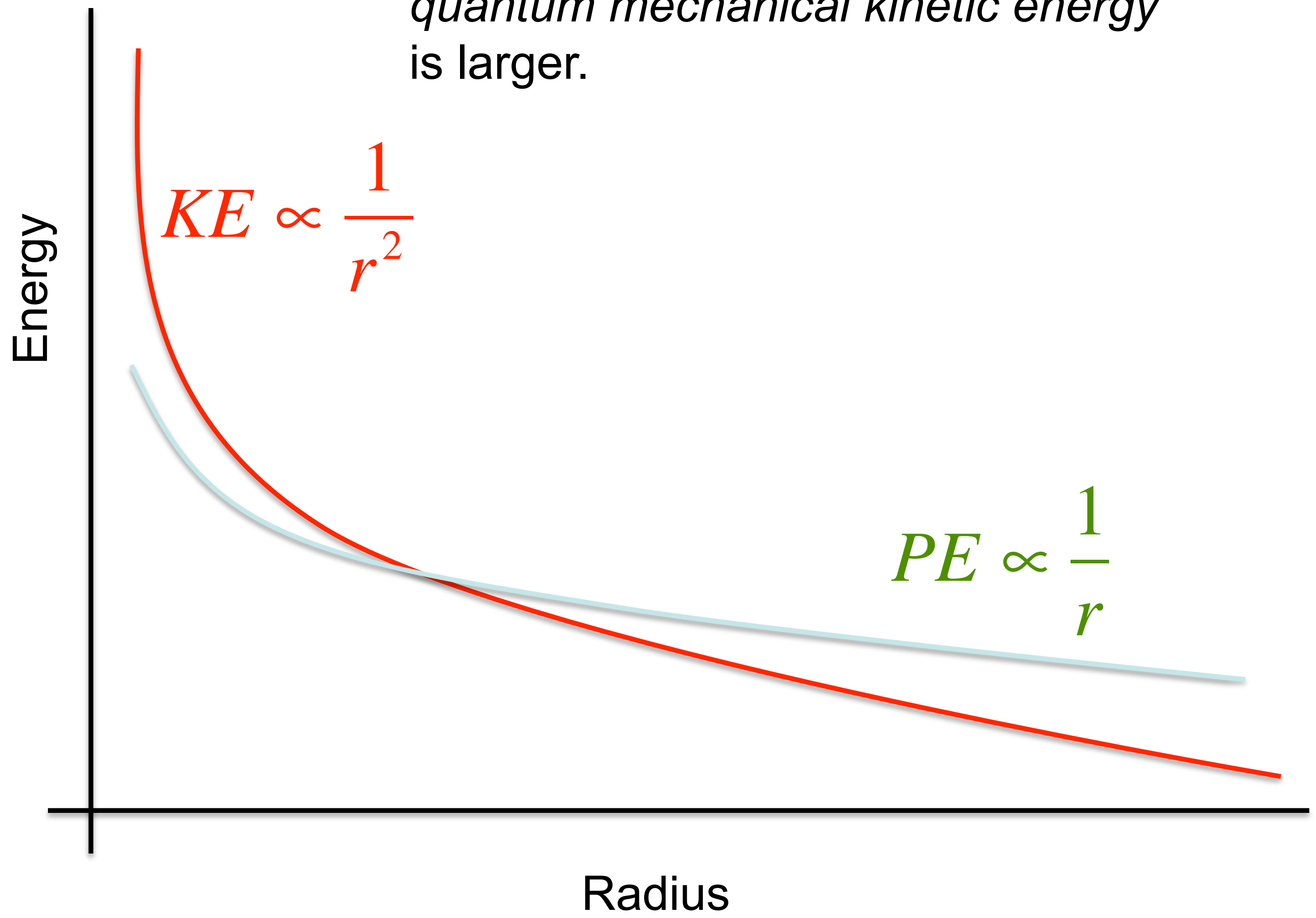
$$\text{Since } p = \frac{h}{\lambda} \Rightarrow KE = \frac{1}{2} m_e v^2 = \frac{p^2}{2m_e} \propto \frac{1}{\lambda^2} \sim \frac{1}{r^2}$$

$$\text{but } PE = -\frac{Ze^2}{r} \propto \frac{1}{r}$$

The kinetic energy increases quadratically with 1/r, the electrical potential, only linearly.

There comes a minimum radius where the electron cannot radiate because the sum of its potential and kinetic energies has reached a minimum.

At large distance electrical attraction dominates. At short distances the *quantum mechanical kinetic energy* is larger.

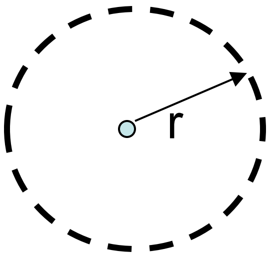


Ground state of the hydrogen atom – Neils Bohr (1913)

(lowest possible energy state)

Must fit the wavelength of the electron inside a circle of radius r, the average distance between the electron and the proton.

$$\lambda = 2\pi r$$



$e^- + p$

$$= \frac{h}{p}$$

The 2π here is rather arbitrary but gives the right answer and omits deeper discussion of "wave functions"

$$\therefore p = \frac{h}{\lambda} = \frac{h}{2\pi r}$$

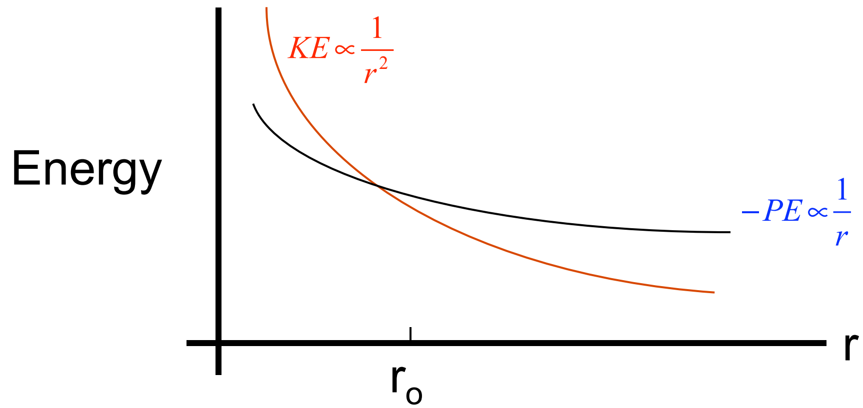
$$KE = \frac{mv^2}{2} = \frac{m^2v^2}{2m} = \frac{p^2}{2m} = \frac{h^2}{2m(4\pi^2 r^2)}$$

Assuming $Z_1 = Z_2 = 1$

$$PE = -\frac{e^2}{r}$$

as before

Note that PE goes as 1/r and KE goes as $1/r^2$



For a single electron bound to a single proton, i.e., hydrogen.

At r_o

$$KE = -\frac{1}{2} PE$$

$$i.e.,\ p = \frac{h}{2\pi r} \text{ and } \mathrm{KE} = \frac{1}{2} m_e v^2 = \frac{(m_e v)^2}{2m_e}$$

$$= \frac{p^2}{2m_e} = \frac{h^2}{8\pi^2 r^2 m_e}$$

$$\frac{h^2}{2m_e(4\pi^2 r_0^2)} = \frac{Ze^2}{2r_o}$$

$$\frac{h^2}{2m_e Ze^2} = \frac{4\pi^2 r_0^2}{2r_o}$$

$$\boxed{r_o = \frac{h^2}{4\pi^2 Z e^2 m_e}}$$

Energy would have to be provided to the electron to make it move any closer to the proton (because it would have to move faster), more energy than e^2/r can give.

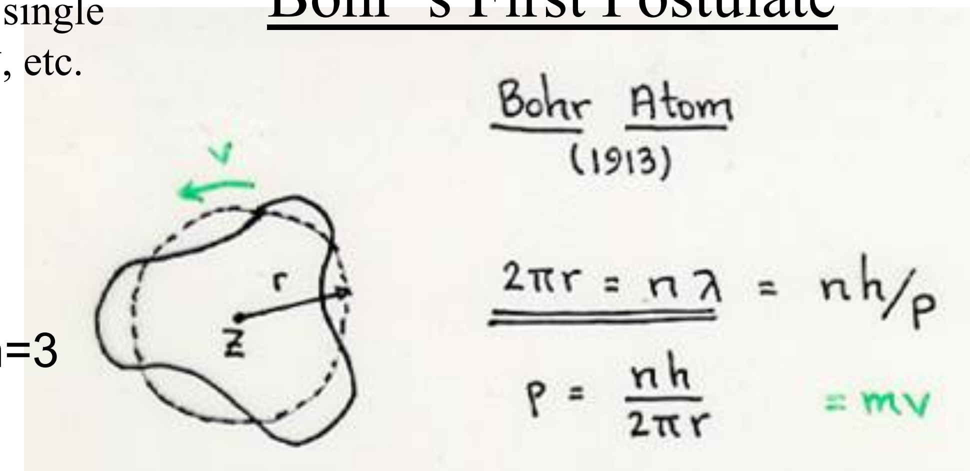
For Z=1 (hydrogen) $r_o = 0.529189379\,\mathring{A} = 5.29189379\,x\,10^{-9}$ cm

This is the (average) radius of the "ground state" of the hydrogen atom, 0.529189… A. It is permanently stable. There is no state with lower energy to make a transition to.

However, there also exist "excited states" of atoms that have a transitory existence.

Bohr's First Postulate

For atoms with a single electron – H, He^+, etc.



n=3

The only possible states of the electron are those for which

$$mvr = \frac{nh}{2\pi}$$

Solve as before:

$$r = \frac{n^2h^2}{4\pi^2 Z e^2 m_e} = 0.53\,\frac{n^2}{Z}\ \text{Angstroms}$$

$$E_{tot} = -\frac{Ze^2}{2r} = -\frac{2\pi^2 Z^2 e^4 m_e}{n^2 h^2}$$

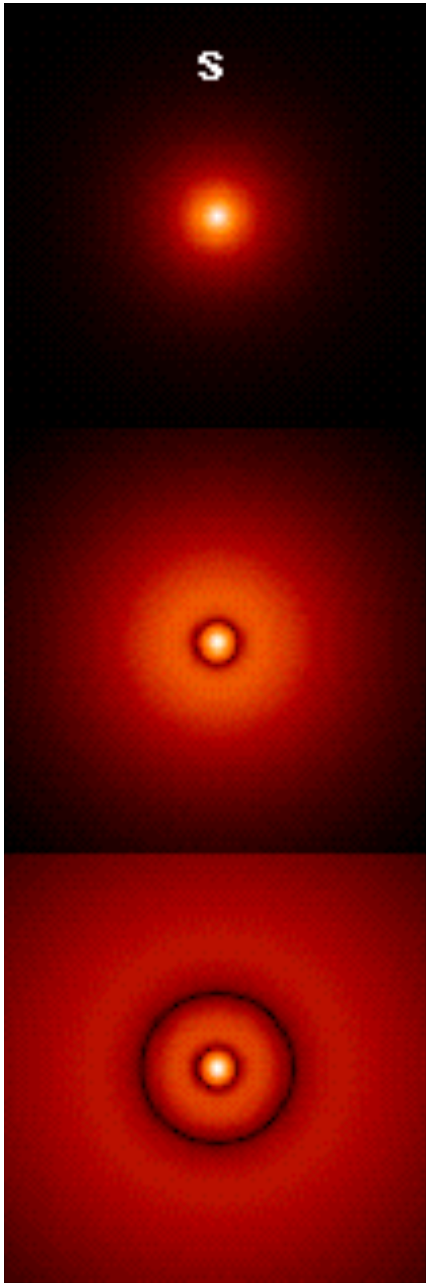
$$\boxed{E_{tot} = -13.6\ \text{eV}\left(\frac{Z^2}{n^2}\right)}$$

$$1\text{eV} \equiv 1.602\times 10^{-12}\ \text{erg}$$

n = 1 is the "ground state"

For atoms with only a single electron.

For hydrogen Z = 1



*

n = 1

In the full quantum mechanical solution the electron is described by a "wave function" that gives its probability for being found at any particular distance from the nucleus.

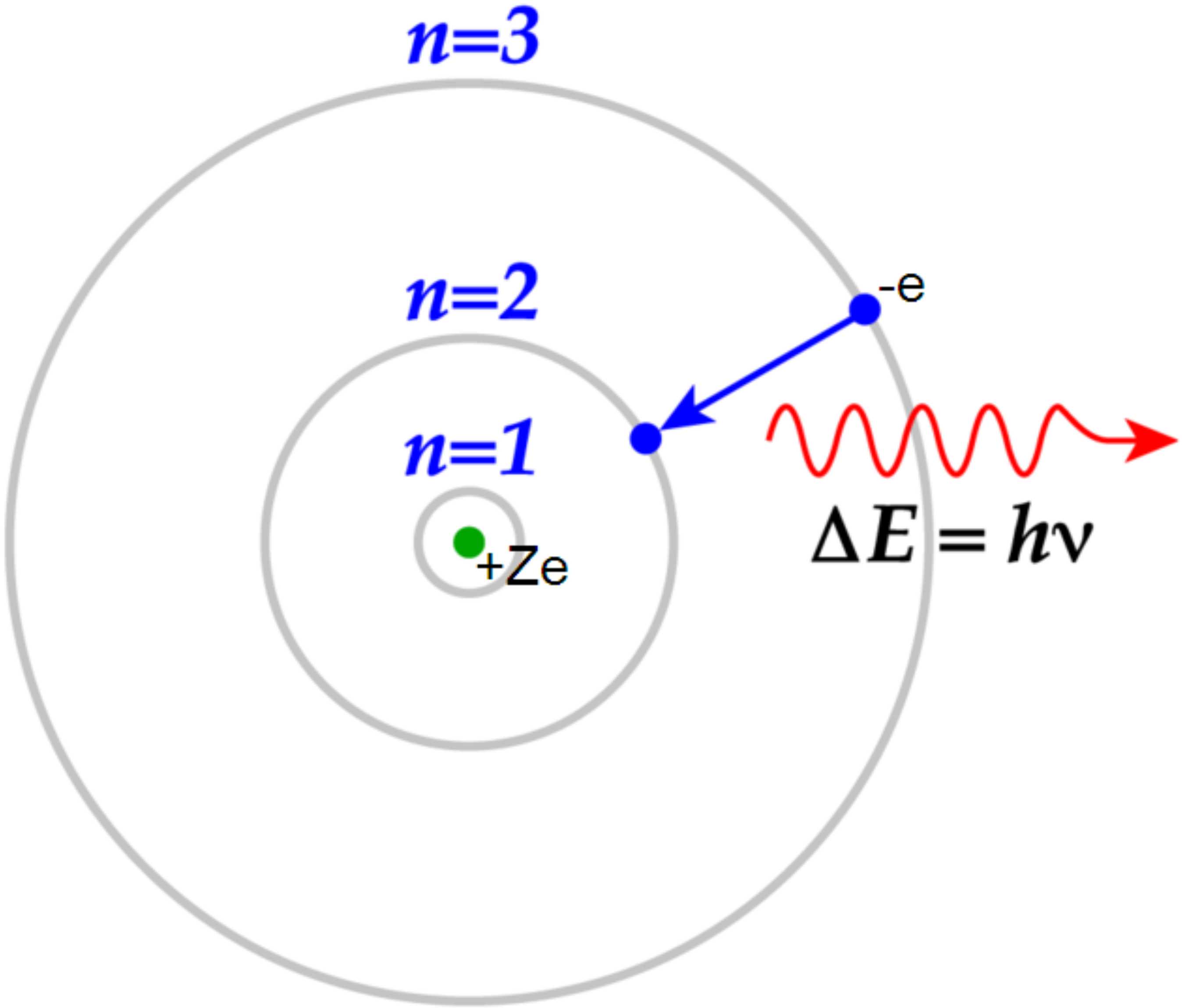
n=2

In the simplest case these distributions are spherical.

n=3

The radius in the Bohr model is the average radius but the energy is precise.

Bohr's Second Postulate



Only the "ground state", n = 1, is permanently stable

Bohr's Second Postulate

Radiation in the form of a single quantum (photon) is Emitted (or absorbed) as the electron makes a transition From one state to another. The energy in the photon is the Difference between the energies of the two states.

emission

$$\mathrm{E}_m \rightarrow \mathrm{E}_n + h\nu$$

absorption

$$\mathrm{E}_n + h\nu \rightarrow \mathrm{E}_m \qquad m > n$$

$$h\nu = \frac{hc}{\lambda} = E_m - E_n$$

$$\frac{1}{\lambda} = \frac{E_m - E_n}{hc} = \frac{2\pi^2 Z^2 e^4 m_e}{h^3 c}\left(\frac{1}{n^2} - \frac{1}{m^2}\right)$$

$$\frac{1}{\lambda_{mn}} = 1.097 \times 10^5 \; Z^2 \left(\frac{1}{n^2} - \frac{1}{m^2}\right) \; \mathrm{cm}^{-1}$$

$$\boxed{\lambda_{mn} = \frac{911.6\ \mathrm{A}}{Z^2}\left(\frac{1}{n^2} - \frac{1}{m^2}\right)^{-1}}$$

(for atoms with only one electron)

$$\lambda_{mn} = \frac{911.6\ \text{A}}{Z^2}\left(\frac{1}{n^2} - \frac{1}{m^2}\right)^{-1}$$

E.g.,

$m = 2, \quad n = 1, \quad Z = 1$

$$\lambda = 911.6\ \mathring{\text{A}}\left(\frac{1}{1^2} - \frac{1}{2^2}\right)^{-1} = 911.6\left(\frac{3}{4}\right)^{-1}$$

$$= 911.6\left(\frac{4}{3}\right) = 1216\ \mathring{\text{A}}$$

$m = 3, n = 1, Z = 1$

$$\lambda = 911.6\left(\frac{1}{1^2} - \frac{1}{3^2}\right)^{-1} = 911.6\left(\frac{8}{9}\right)^{-1}$$

$$= 911.6\left(\frac{9}{8}\right) = 1026\ \mathring{\text{A}}$$

Lines that start or end on n=1 are called the "Lyman" series. All are between 911.6 and 1216 A.

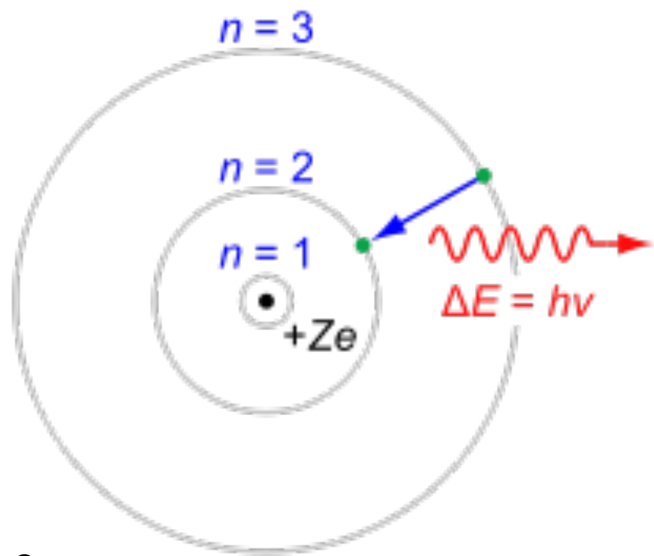
$m = 3, n = 2, Z = 1$

$$\lambda = 911.6\left(\frac{1}{2^2} - \frac{1}{3^2}\right)^{-1} = 911.6\left(\frac{1}{4} - \frac{1}{9}\right)^{-1}$$

$$= 911.6\left(\frac{5}{36}\right)^{-1} = 911.6\left(\frac{36}{5}\right) = 6563\ \mathring{\text{A}}$$

Lines that start or end on n=2 are called the "Balmer" series. All are between 3646 and 6564 A.

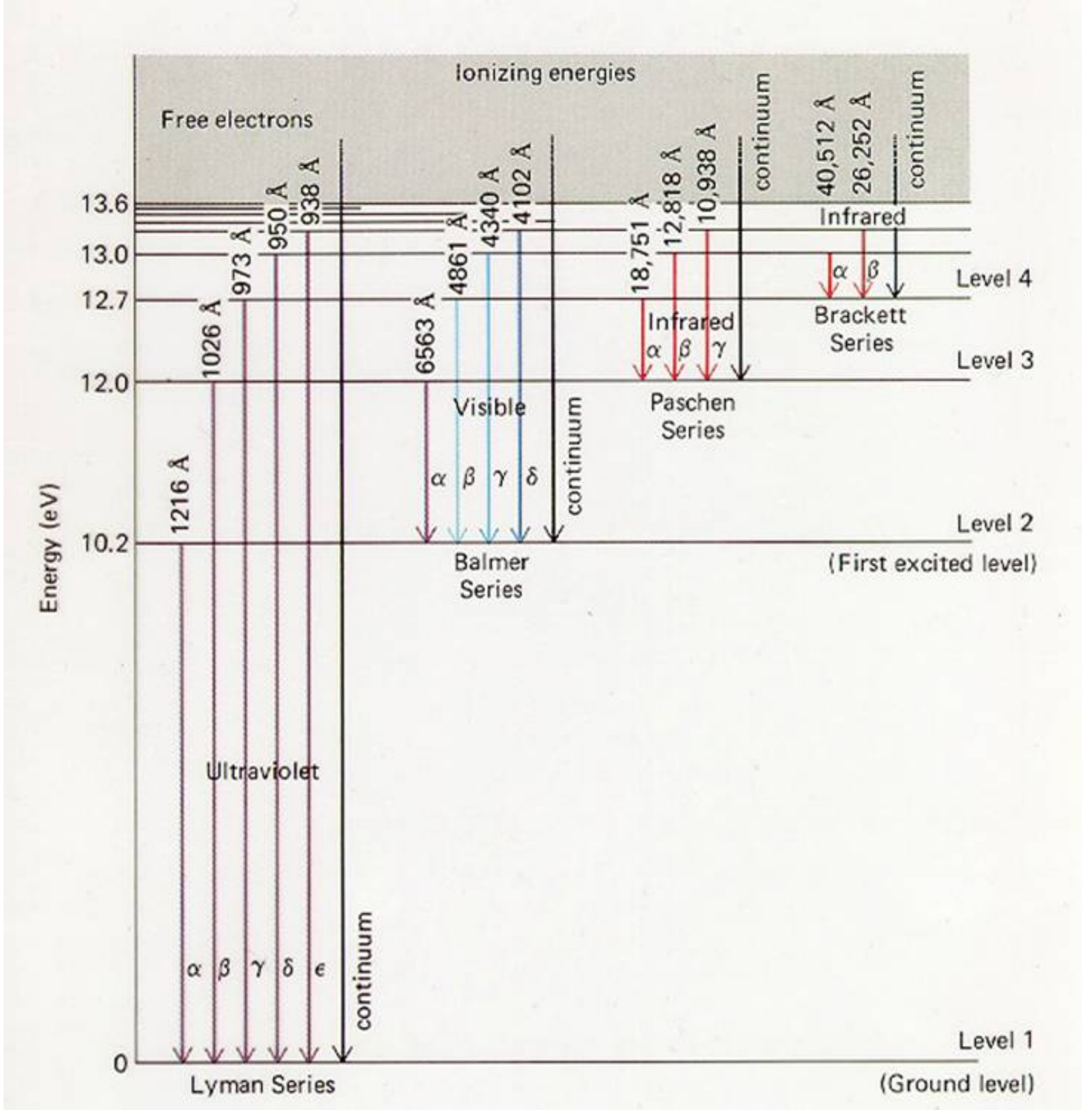
BALMER SERIES



$5 \rightarrow 2$ $\quad$ $4 \rightarrow 2$ $\quad$ $3 \rightarrow 2$

H_γ $\quad$ H_β $\quad$ H_α

Hydrogen emission line spectrum
Balmer series

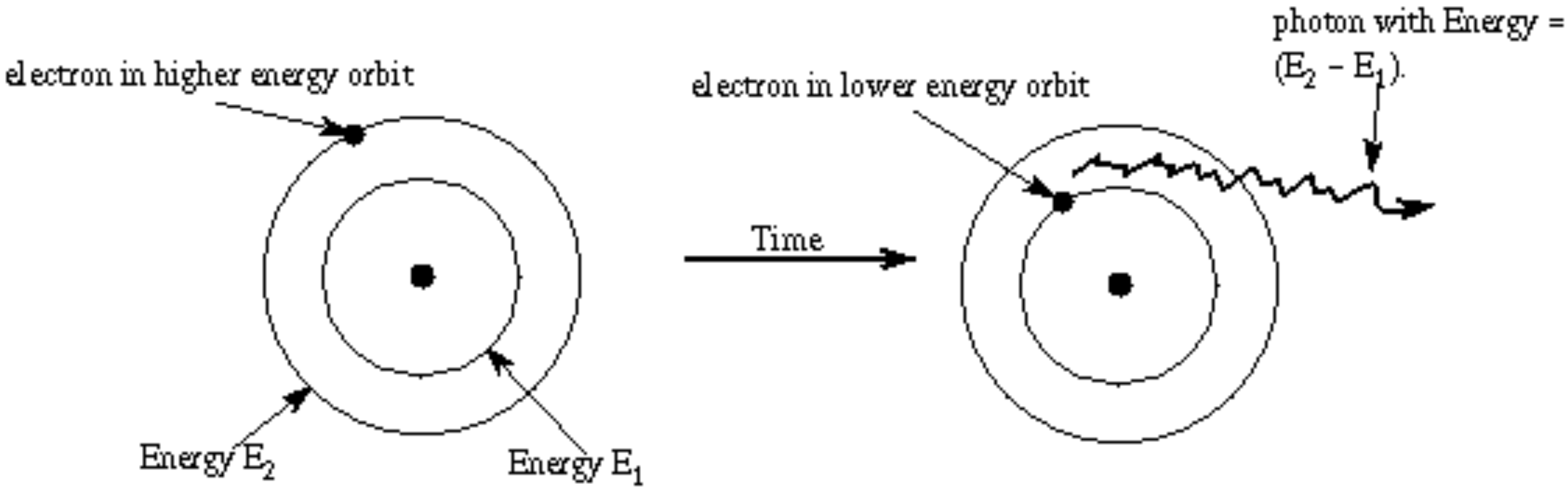


$H_{\alpha,\beta,\gamma,.....}$

$Ly_{\alpha,\beta,\gamma,...}$

Adjusting the energy of each state in hydrogen by adding 13.6 eV (so that the ground state becomes zero), one gets a diagram where the energies of the transitions can be read off easily.

Emission line

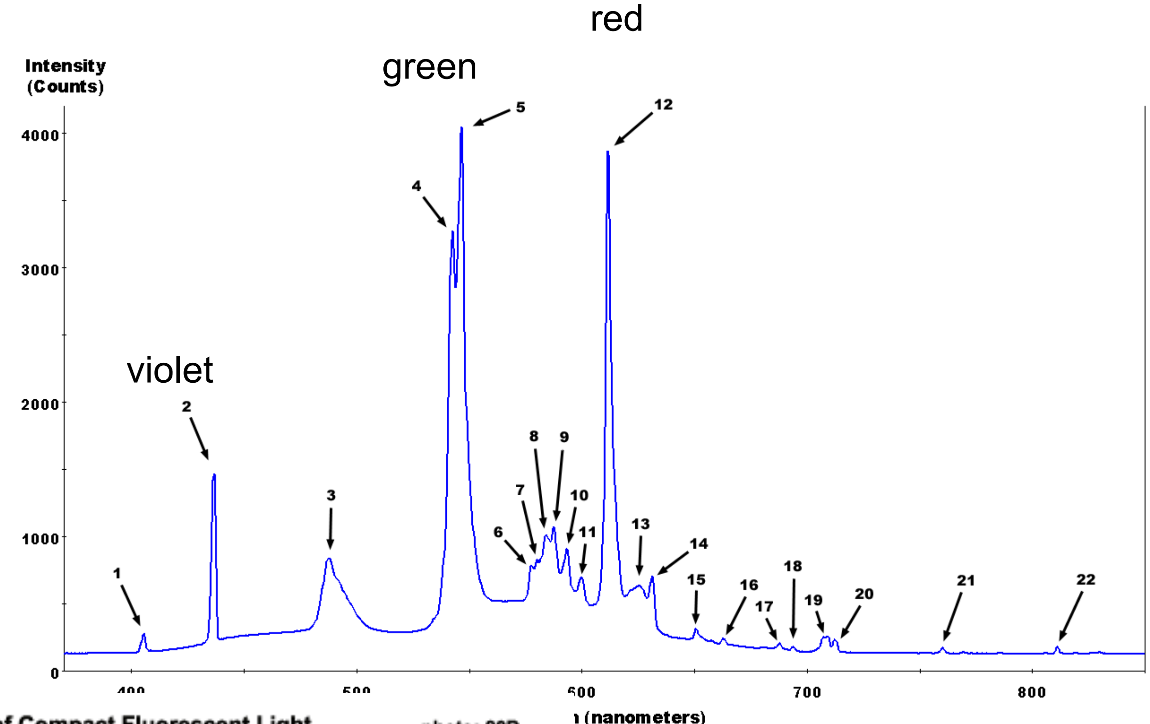


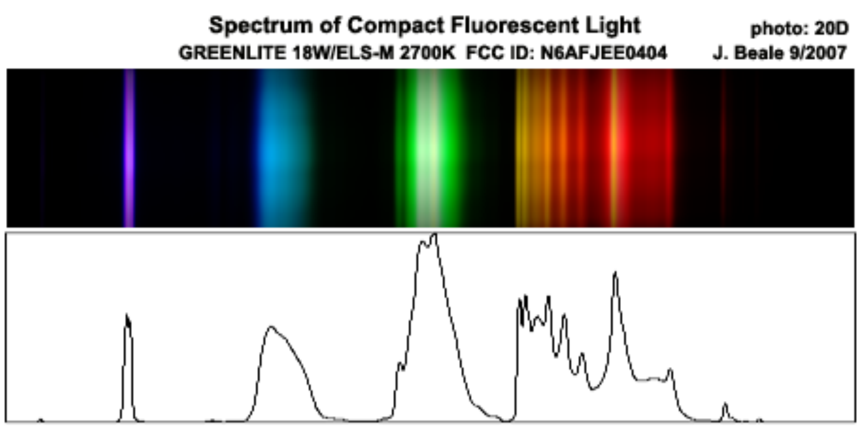
Atom was hit/bumped by another atom and gained some energy = (E_2 – E_1). Electron in higher energy orbit (E_2).

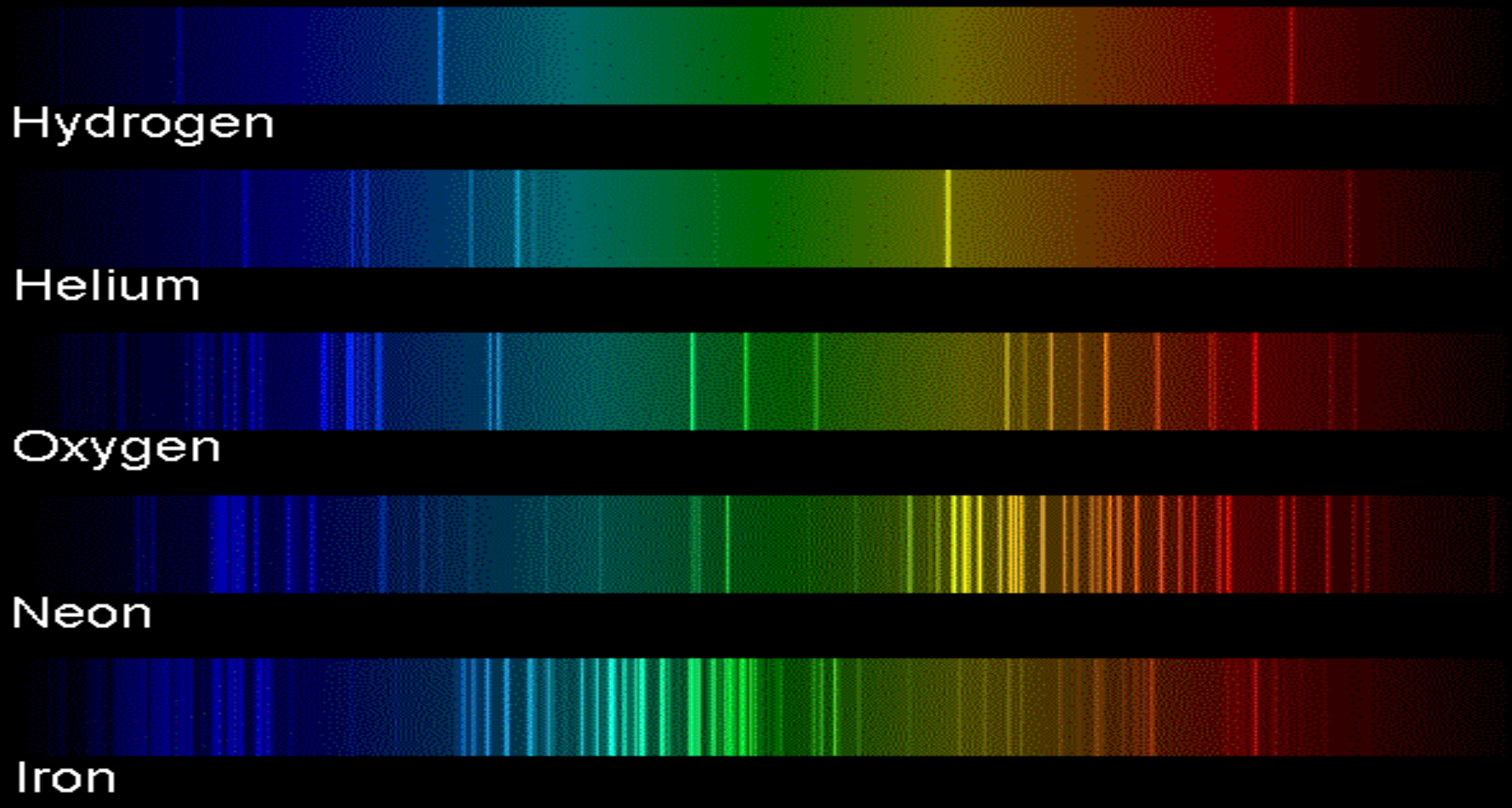
Emission line produced!

Fluorescent Light Fixture

Peak number	Wavelength of peak (nm)	Species producing peak
1	405.4	mercury
2	436.6	mercury
3	487.7	terbium from Tb^{3+}
4	542.4	terbium from Tb^{3+}
5	546.5	mercury
6	577.7	possibly mercury
7	580.2	mercury or europium in $Eu^{+3}:Y_2O_3$ or terbium likely Tb^{3+}
8	584.0	possibly terbium from Tb^{3+}
9	587.6	likely europium in $Eu^{+3}:Y_2O_3$
10	593.4	likely europium in $Eu^{+3}:Y_2O_3$
11	599.7	likely europium in $Eu^{+3}:Y_2O_3$
12	611.6	europium in $Eu^{+3}:Y_2O_3$
13	625.7	likely terbium from Tb^{3+}
14	631.1	likely europium in $Eu^{+3}:Y_2O_3$
15	650.8	likely europium in $Eu^{+3}:Y_2O_3$
16	662.6	





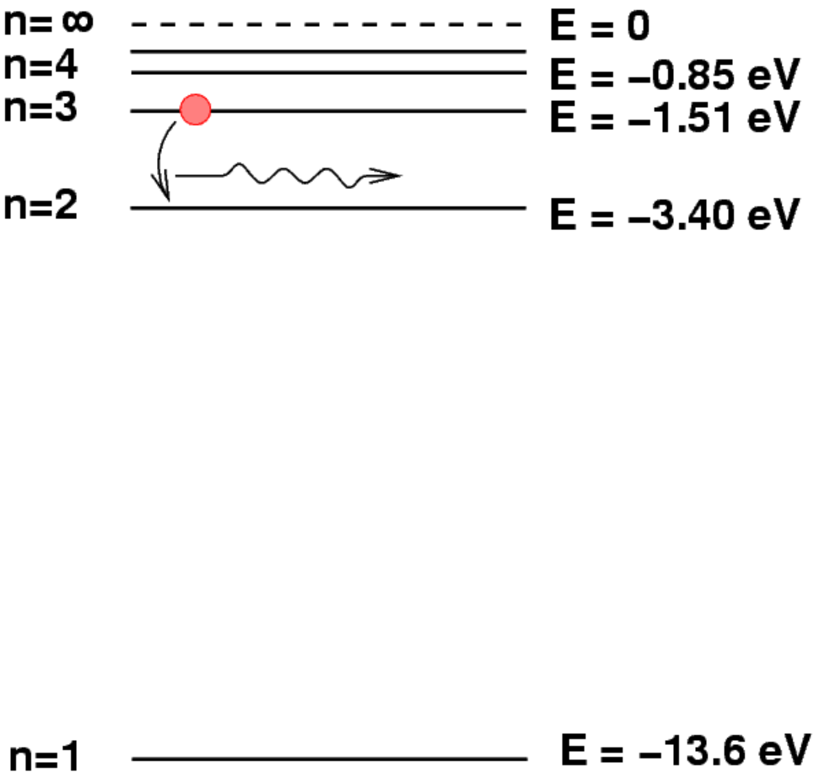

Hydrogen
Helium
Oxygen
Neon
Iron

How are excited states populated?

- Absorb a photon of the right energy
- Collisions
- Ionization - recombination

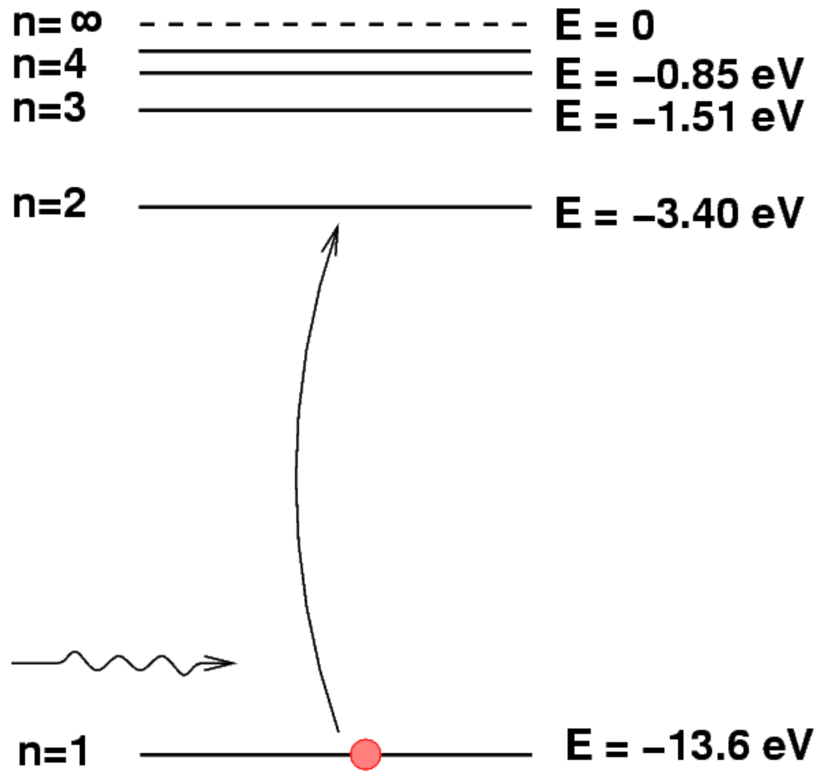
http://spiff.rit.edu/classes/phys301/lectures/spec_lines/Atoms_Nav.swf

Emission

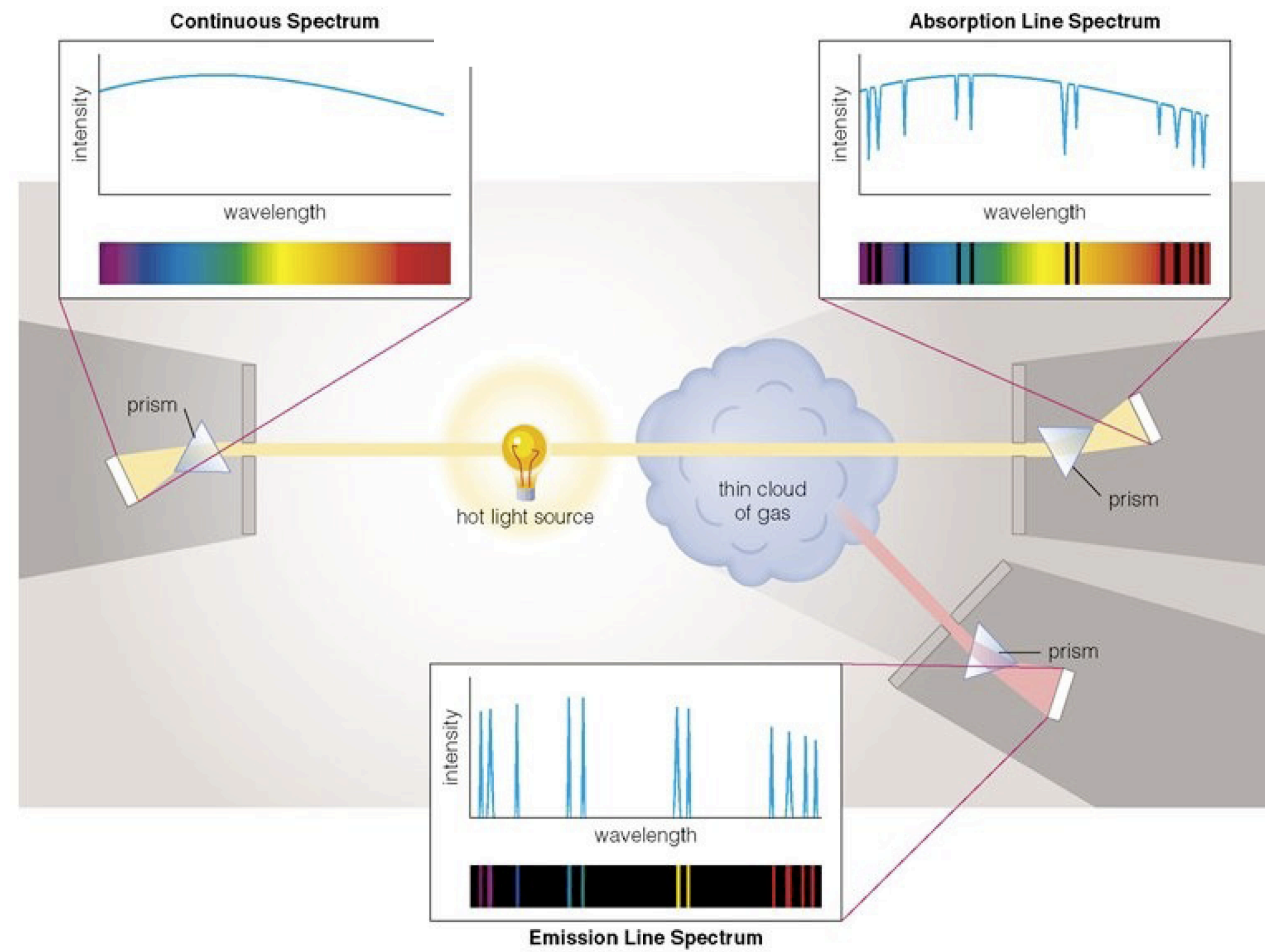


Emission – H-alpha

Absorption

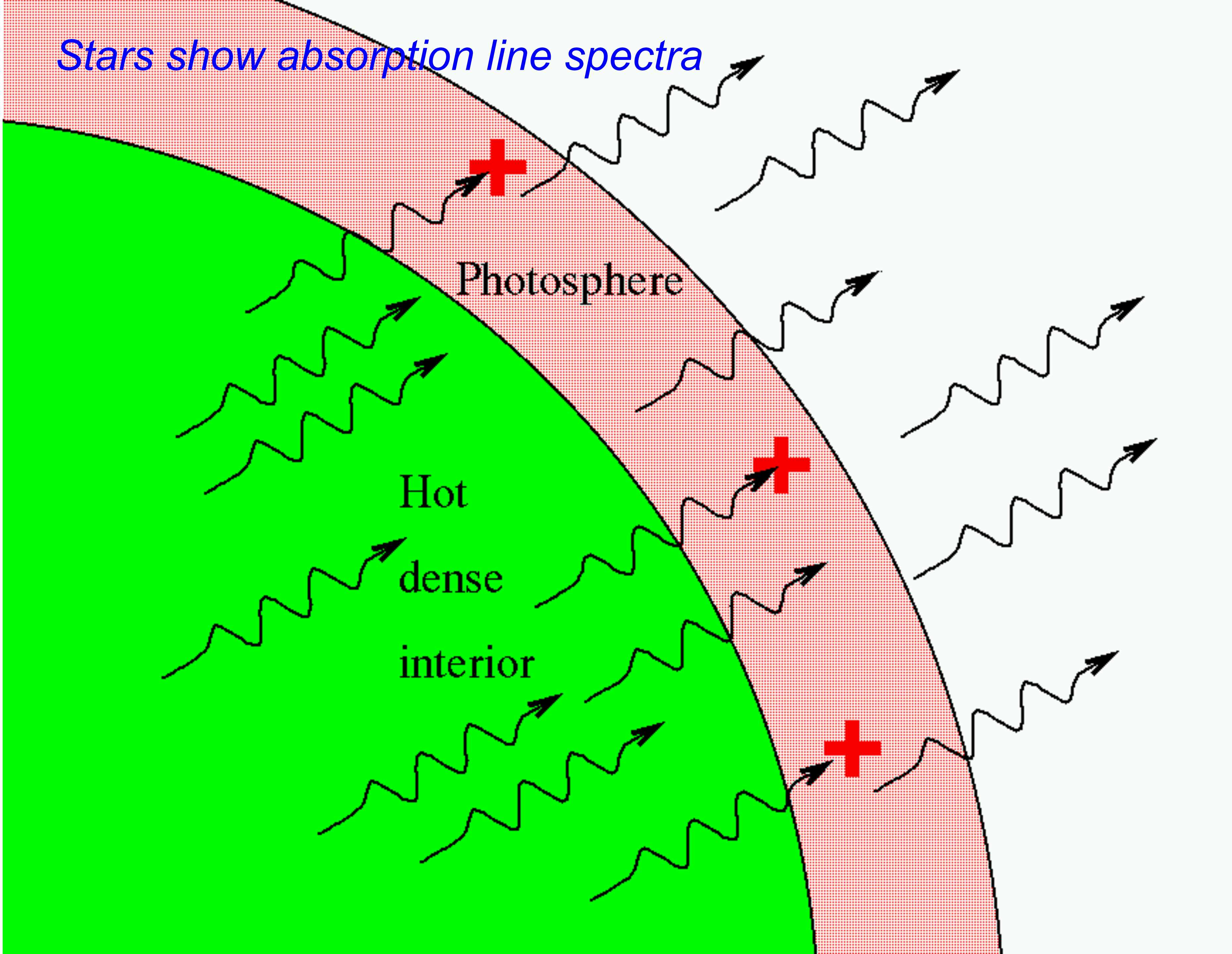


Absorption – Ly-alpha

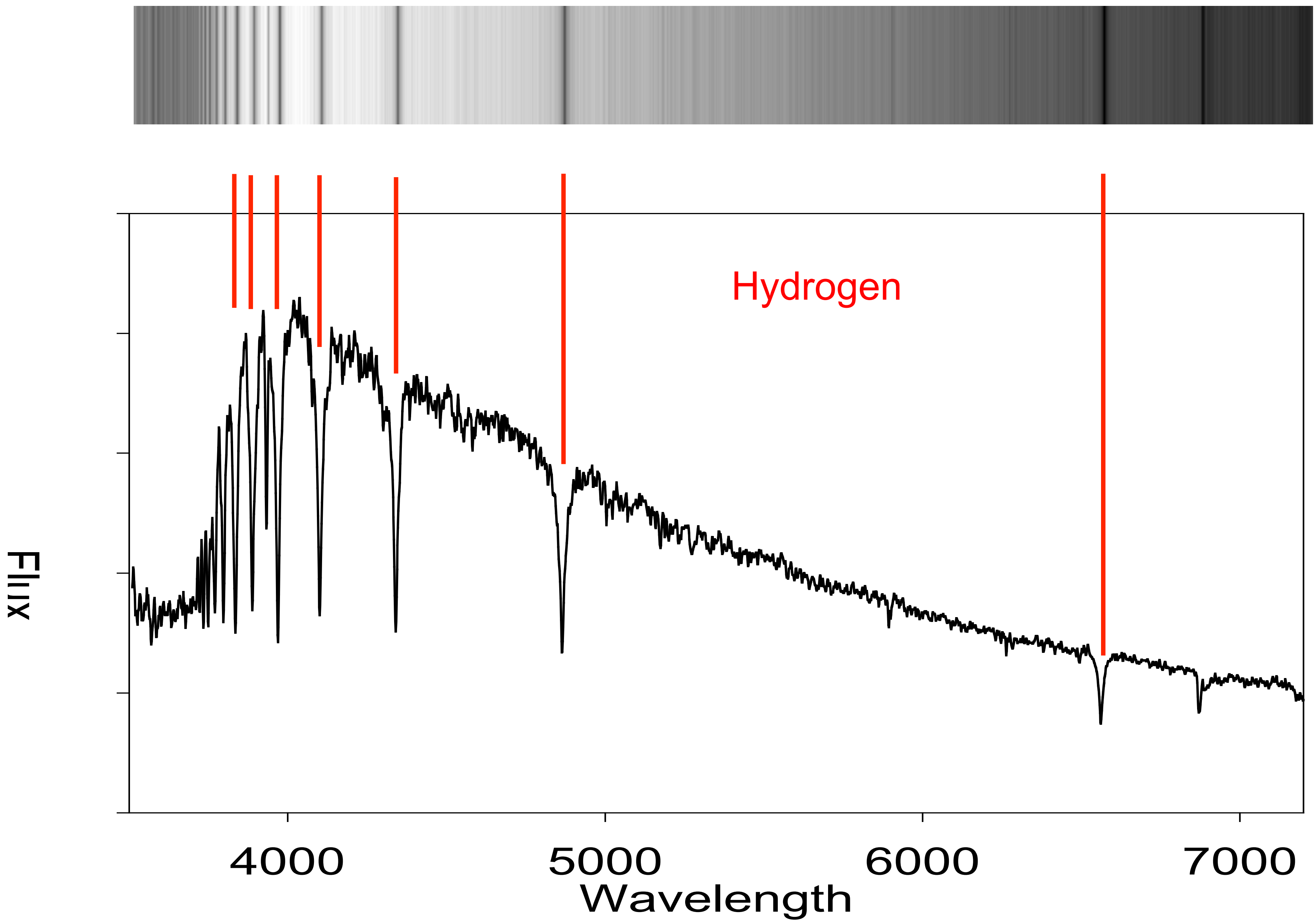


Continuous Spectrum
intensity
wavelength
Absorption Line Spectrum
intensity
wavelength
prism
hot light source
thin cloud
of gas
prism
prism
intensity
wavelength
Emission Line Spectrum

Stars show absorption line spectra

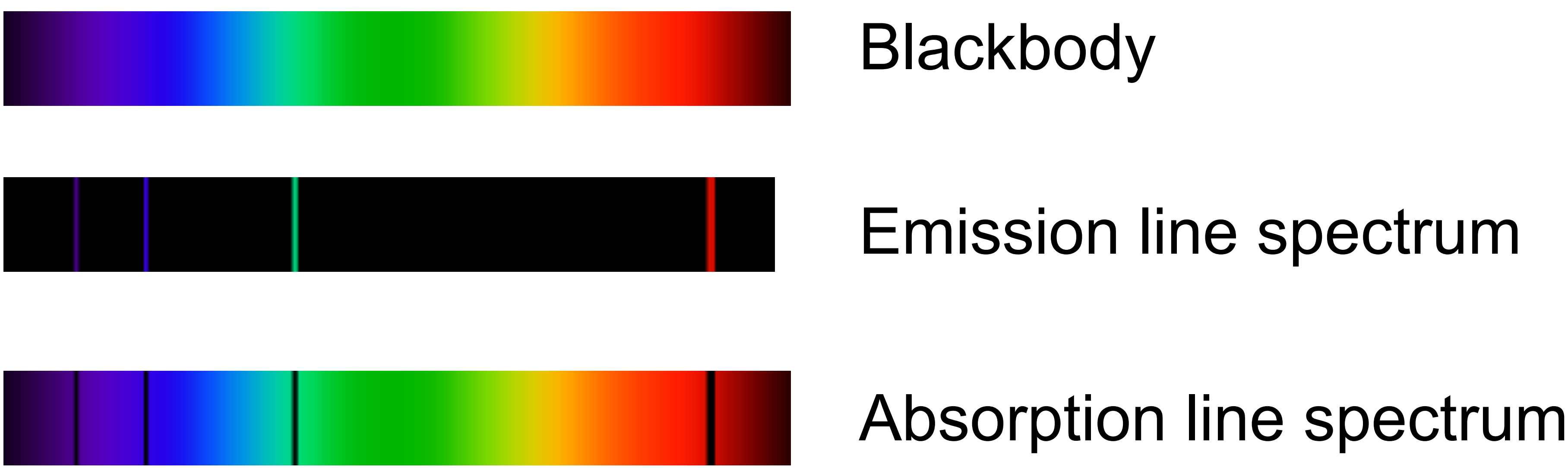


Absorption Line Spectrum (not the sun)

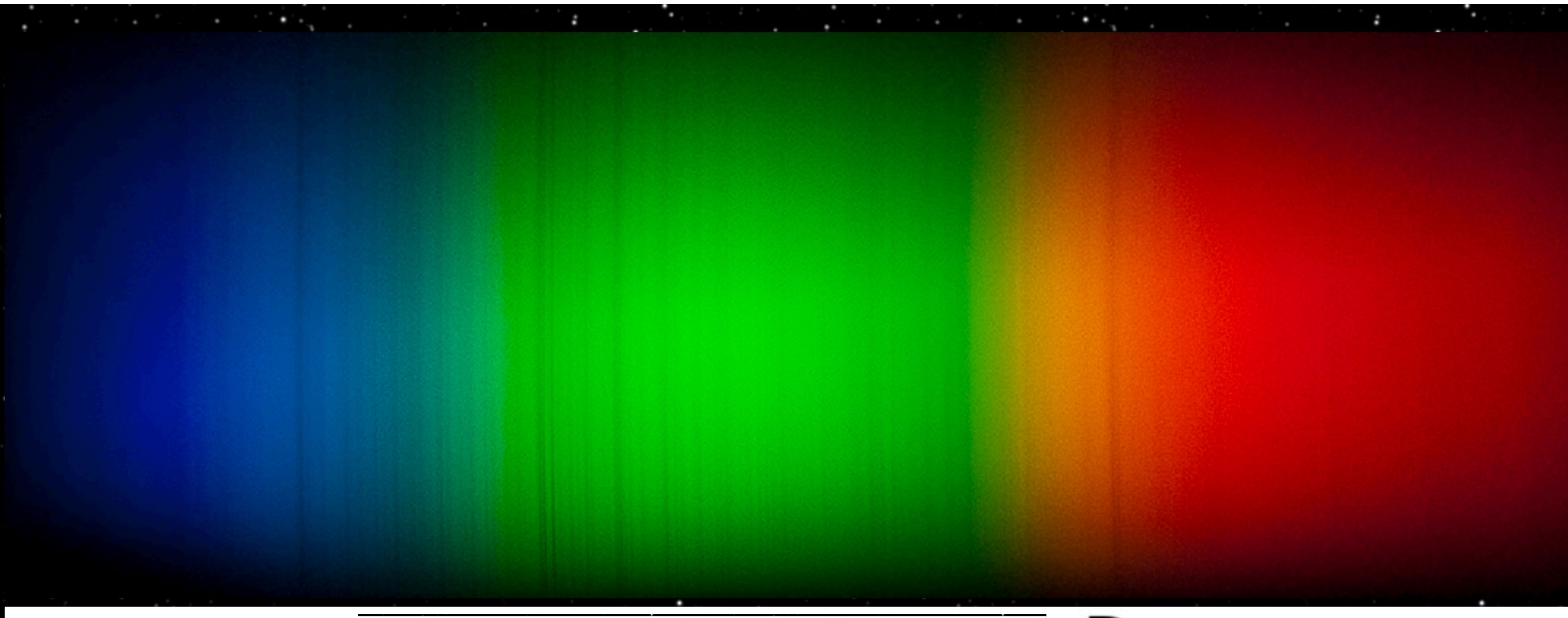


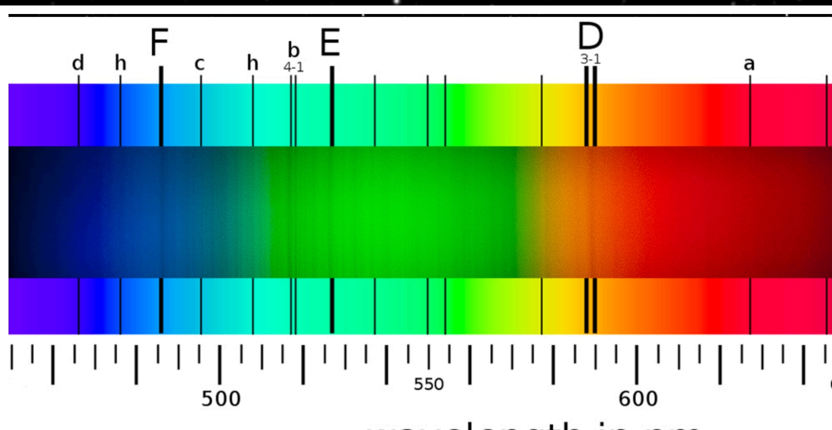
When we examine the spectra of stars, with a few exceptions to be discussed later, we see blackbody spectra with a superposition of *absorption* lines.

The identity and intensity of the “spectral lines” that are present reflect the temperature, density and composition of the stellar photosphere.



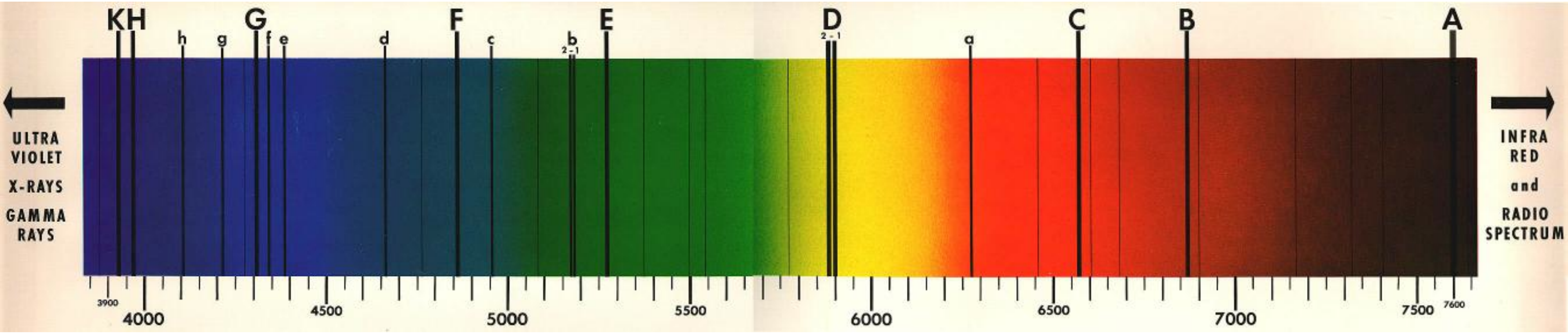
The sun through a low resolution spectrograph





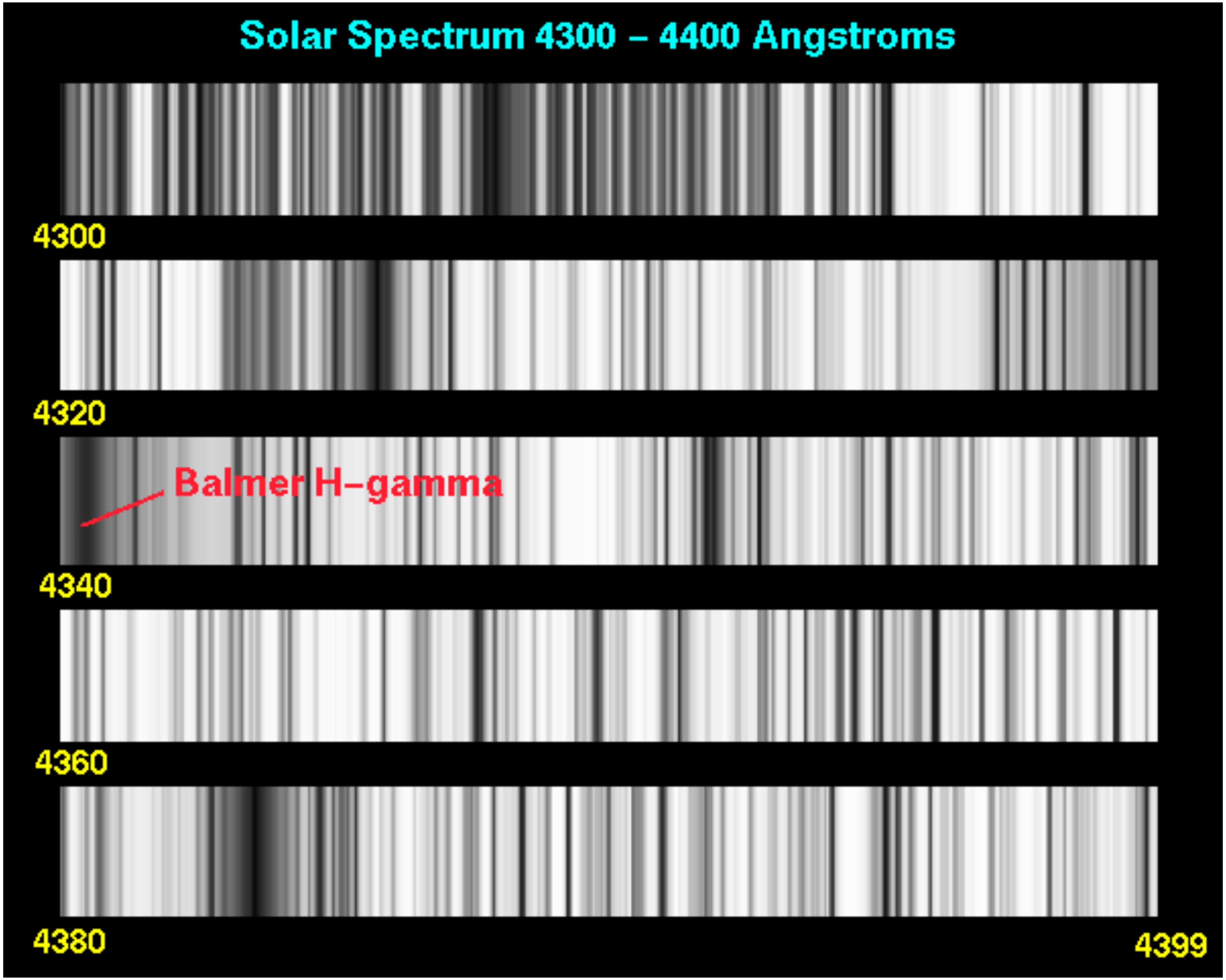
The solar spectrum

C = Balmer alpha
F = Balmer beta
f = Balmer gamma
B = oxygen
D = sodium
E = iron
H, K = singly ionized calcium
others = Fe, Mg, Na, etc.



Wollaton (1802) discovered dark lines in the solar spectrum. Fraunhaufer rediscovered them (1817) and studied the systematics

(Part of) the high resolution solar spectrum

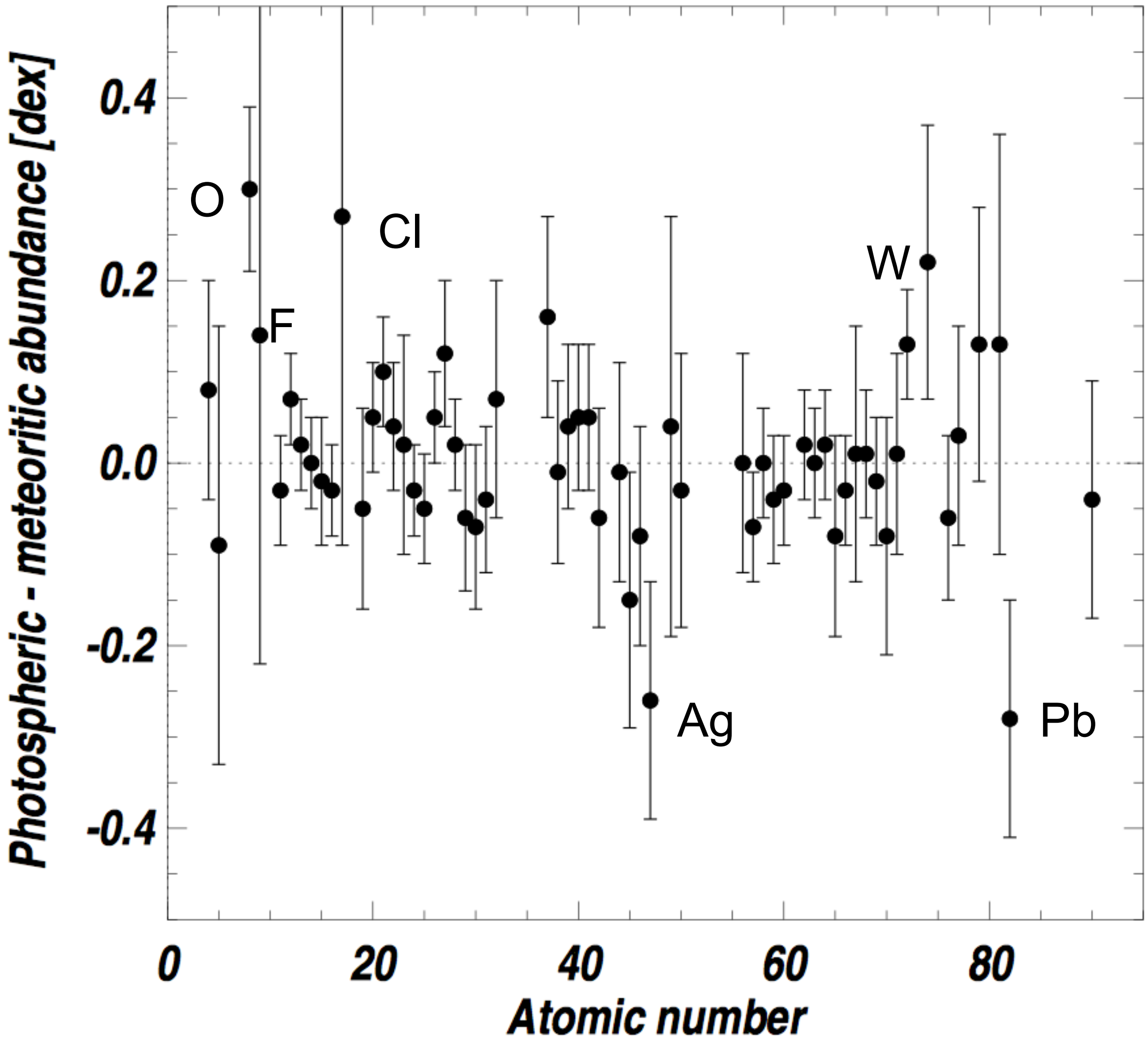


Almost every element has been observed spectroscopically in the sun and has an accurate abundance determination. The rest, except for noble gases and “volatile” elements, have an accurate determination from primitive meteorites (carbonaceous chondrites)

Table 1: Element abundances in the present-day solar photosphere. Also given are the corresponding values for CI carbonaceous chondrites (Lodders, Palme & Gail 2009). Indirect photospheric estimates have been used for the noble gases (Sect. 3.9).

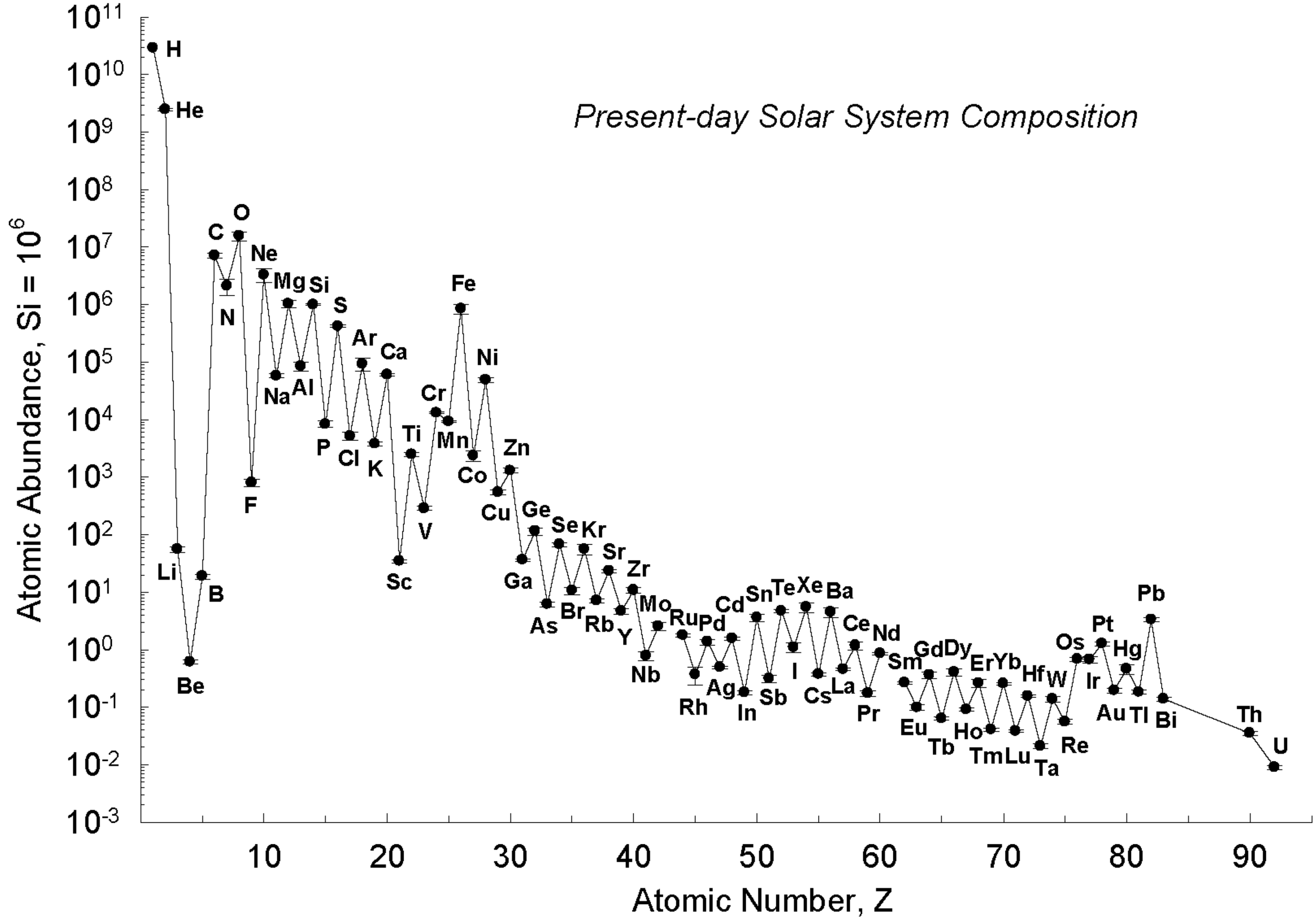
	Elem.	Photosphere	Meteorites		Elem.	Photosphere	Meteorites
1	H	12.00	8.22 ± 0.04	44	Ru	1.75 ± 0.08	1.76 ± 0.03
2	He	$[10.93 \pm 0.01]$	1.29	45	Rh	0.91 ± 0.10	1.06 ± 0.04
3	Li	1.05 ± 0.10	3.26 ± 0.05	46	Pd	1.57 ± 0.10	1.65 ± 0.02
4	Be	1.38 ± 0.09	1.30 ± 0.03	47	Ag	0.94 ± 0.10	1.20 ± 0.02
5	B	2.70 ± 0.20	2.79 ± 0.04	48	Cd		1.71 ± 0.03
6	C	8.43 ± 0.05	7.39 ± 0.04	49	In	0.80 ± 0.20	0.76 ± 0.03
7	N	7.83 ± 0.05	6.26 ± 0.06	50	Sn	2.04 ± 0.10	2.07 ± 0.06
8	O	8.69 ± 0.05	8.40 ± 0.04	51	Sb		1.01 ± 0.06
9	F	4.56 ± 0.30	4.42 ± 0.06	52	Te		2.18 ± 0.03
10	Ne	$[7.93 \pm 0.10]$	-1.12	53	I		1.55 ± 0.08
11	Na	6.24 ± 0.04	6.27 ± 0.02	54	Xe	$[2.24 \pm 0.06]$	-1.95
12	Mg	7.60 ± 0.04	7.53 ± 0.01	55	Cs		1.08 ± 0.02
13	Al	6.45 ± 0.03	6.43 ± 0.01	56	Ba	2.18 ± 0.09	2.18 ± 0.03
14	Si	7.51 ± 0.03	7.51 ± 0.01	57	La	1.10 ± 0.04	1.17 ± 0.02
15	P	5.41 ± 0.03	5.43 ± 0.04	58	Ce	1.58 ± 0.04	1.58 ± 0.02
16	S	7.12 ± 0.03	7.15 ± 0.02	59	Pr	0.72 ± 0.04	0.76 ± 0.03
17	Cl	5.50 ± 0.30	5.23 ± 0.06	60	Nd	1.42 ± 0.04	1.45 ± 0.02
18	Ar	$[6.40 \pm 0.13]$	-0.50	62	Sm	0.96 ± 0.04	0.94 ± 0.02
19	K	5.03 ± 0.09	5.08 ± 0.02	63	Eu	0.52 ± 0.04	0.51 ± 0.02
20	Ca	6.34 ± 0.04	6.29 ± 0.02	64	Gd	1.07 ± 0.04	1.05 ± 0.02

20	Ca	6.34 ± 0.04	6.29 ± 0.02	64	Gd	1.07 ± 0.04	1.05 ± 0.02
21	Sc	3.15 ± 0.04	3.05 ± 0.02	65	Tb	0.30 ± 0.10	0.32 ± 0.03
22	Ti	4.95 ± 0.05	4.91 ± 0.03	66	Dy	1.10 ± 0.04	1.13 ± 0.02
23	V	3.93 ± 0.08	3.96 ± 0.02	67	Ho	0.48 ± 0.11	0.47 ± 0.03
24	Cr	5.64 ± 0.04	5.64 ± 0.01	68	Er	0.92 ± 0.05	0.92 ± 0.02
25	Mn	5.43 ± 0.05	5.48 ± 0.01	69	Tm	0.10 ± 0.04	0.12 ± 0.03
26	Fe	7.50 ± 0.04	7.45 ± 0.01	70	Yb	0.84 ± 0.11	0.92 ± 0.02
27	Co	4.99 ± 0.07	4.87 ± 0.01	71	Lu	0.10 ± 0.09	0.09 ± 0.02
28	Ni	6.22 ± 0.04	6.20 ± 0.01	72	Hf	0.85 ± 0.04	0.71 ± 0.02
29	Cu	4.19 ± 0.04	4.25 ± 0.04	73	Ta		-0.12 ± 0.04
30	Zn	4.56 ± 0.05	4.63 ± 0.04	74	W	0.85 ± 0.12	0.65 ± 0.04
31	Ga	3.04 ± 0.09	3.08 ± 0.02	75	Re		0.26 ± 0.04
32	Ge	3.65 ± 0.10	3.58 ± 0.04	76	Os	1.40 ± 0.08	1.35 ± 0.03
33	As		2.30 ± 0.04	77	Ir	1.38 ± 0.07	1.32 ± 0.02
34	Se		3.34 ± 0.03	78	Pt		1.62 ± 0.03
35	Br		2.54 ± 0.06	79	Au	0.92 ± 0.10	0.80 ± 0.04
36	Kr	$[3.25 \pm 0.06]$	-2.27	80	Hg		1.17 ± 0.08
37	Rb	2.52 ± 0.10	2.36 ± 0.03	81	Tl	0.90 ± 0.20	0.77 ± 0.03
38	Sr	2.87 ± 0.07	2.88 ± 0.03	82	Pb	1.75 ± 0.10	2.04 ± 0.03
39	Y	2.21 ± 0.05	2.17 ± 0.04	83	Bi		0.65 ± 0.04
40	Zr	2.58 ± 0.04	2.53 ± 0.04	90	Th	0.02 ± 0.10	0.06 ± 0.03
41	Nb	1.46 ± 0.04	1.41 ± 0.04	92	U		-0.54 ± 0.03
42	Mo	1.88 ± 0.08	1.94 ± 0.04				



Asplund et al (2009; ARAA)

Figure 7: Difference between the logarithmic abundances determined from the solar photosphere and the CI carbonaceous chondrites as a function of atomic number. With a few exceptions the agreement is excellent. Note that due to depletion in the Sun and meteorites, the data points for Li, C, N and the noble gases fall outside the range of the figure.



Present-day Solar System Composition
Atomic Abundance, Si = 10^6
Atomic Number, Z
10^{11}
10^{10}
10^9
10^8
10^7
10^6
10^5
10^4
10^3
10^2
10^1
10^0
10^{-1}
10^{-2}
10^{-3}
10
20
30
40
50
60
70
80
90
H
He
Li
Be
B
C
N
O
F
Ne
Na
Mg
Al
Si
P
S
Cl
Ar
K
Ca
Sc
Ti
V
Cr
Mn
Fe
Co
Ni
Cu
Zn
Ga
Ge
As
Se
Br
Kr
Rb
Sr
Y
Zr
Nb
Mo
Ru
Rh
Pd
Ag
Cd
In
Sn
Sb
Te
I
Xe
Cs
Ba
La
Ce
Pr
Nd
Sm
Eu
Gd
Tb
Dy
Ho
Er
Tm
Yb
Lu
Hf
Ta
W
Re
Os
Ir
Pt
Au
Hg
Tl
Pb
Bi
Th
U

Ionization

As the temperature in a gas is raised, electrons will be removed by collisions and interactions with light. The gas comes *ionized*.

The degree of ionization depends on the atom considered and the temperature.

Notation: Ionization stages

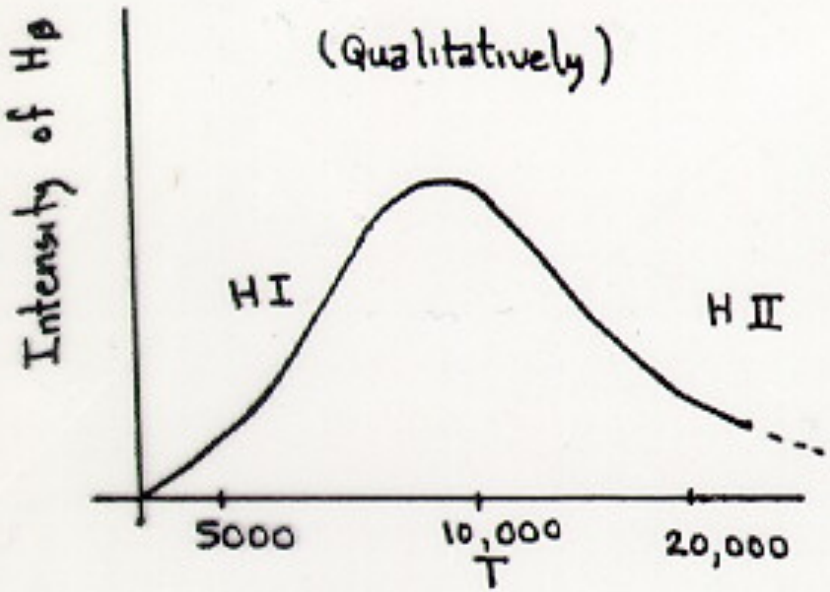
H I	neutral hydrogen	1 p	1 e
H II	ionized hydrogen	1 p	0 e
He I	neutral helium	2 p	2 e
He II	singly ionized helium	2 p	1 e
He III	doubly ionized helium	2 p	0 e
C I	neutral carbon	6 p	6 e
C II	C^{+}	6 p	5 e
C III	C^{++}	6 p	4 e

etc.

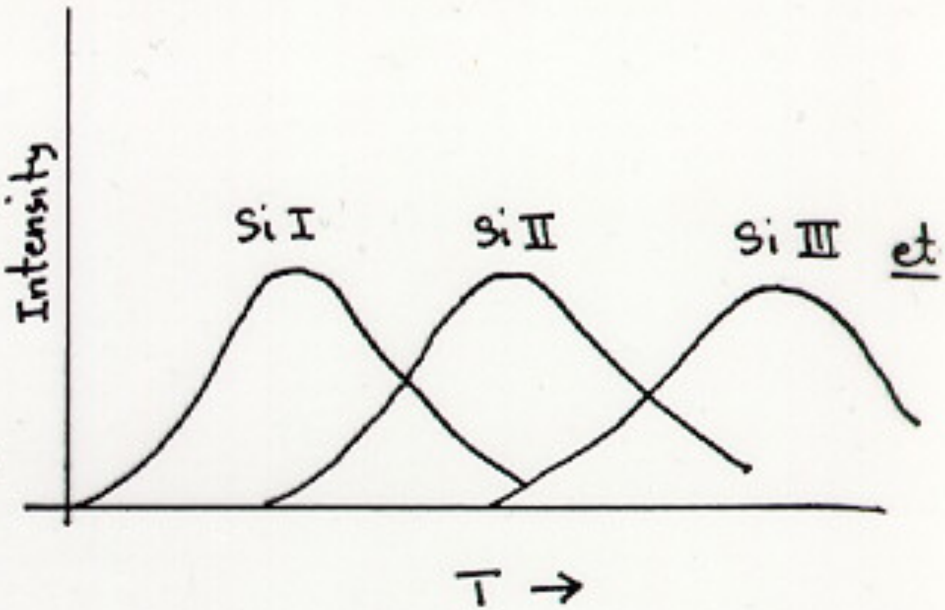
The ionization energy is the energy required to remove a single electron from a given ion. The excitation energy is the energy required to excite an electron from the ground state to the first excited state.

Ion	Excitation energy (eV)	Ionization energy (eV)
H I	10.2	13.6
He I	20.9	24.5
He II	40.8	54.4
rare Li I	1.8	5.4
Ne I	16.6	21.5
Na I	2.1	5.1
Mg I	2.7	7.6
Ca I	1.9	6.1

Li is He plus one proton, Na is Ne plus 1 proton, Ca is Ar plus 2 protons.
The noble gases have closed electron shells and are very stable.

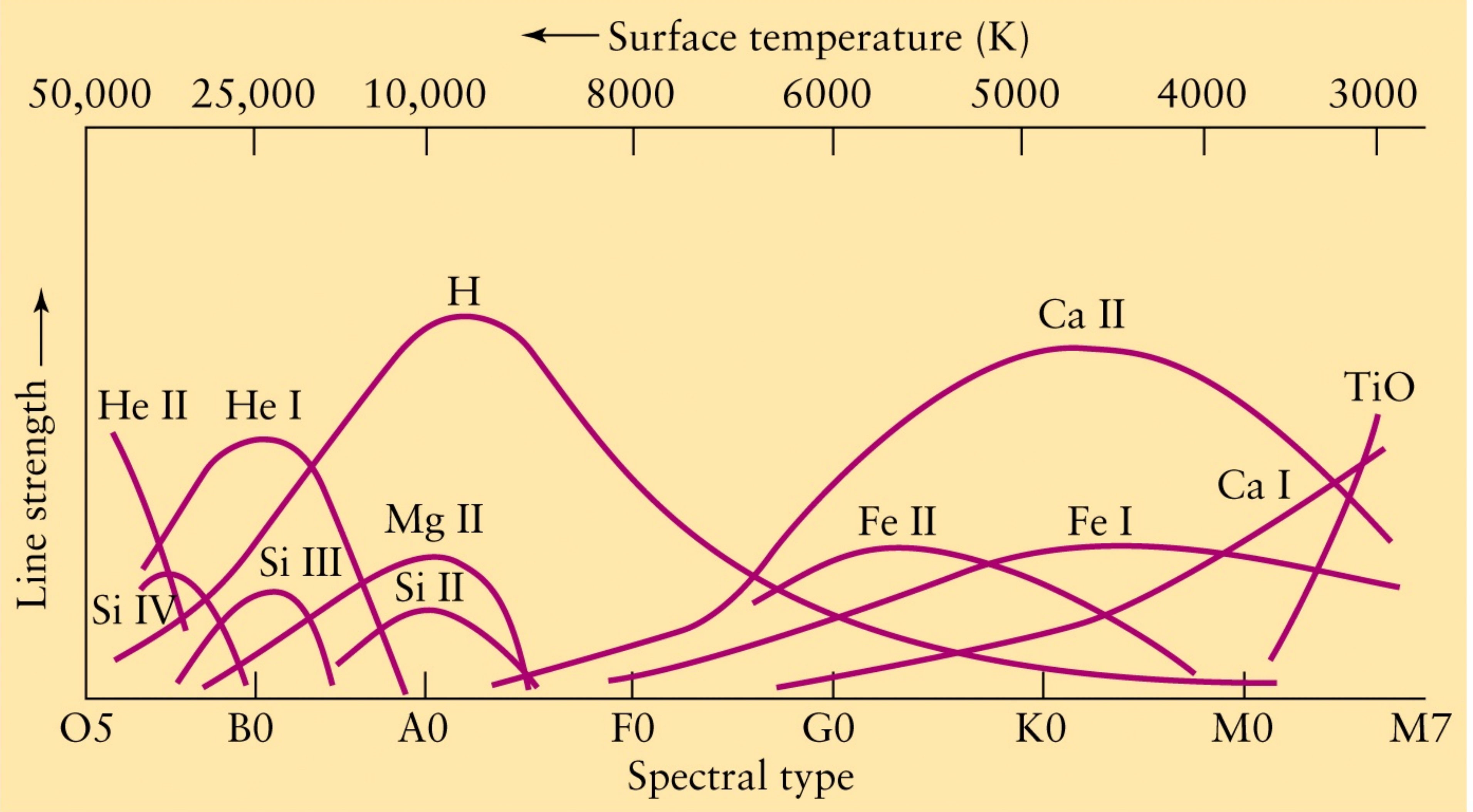


Intensity of Hβ
(Qualitatively)
H I
H II
5000
10,000
T
20,000

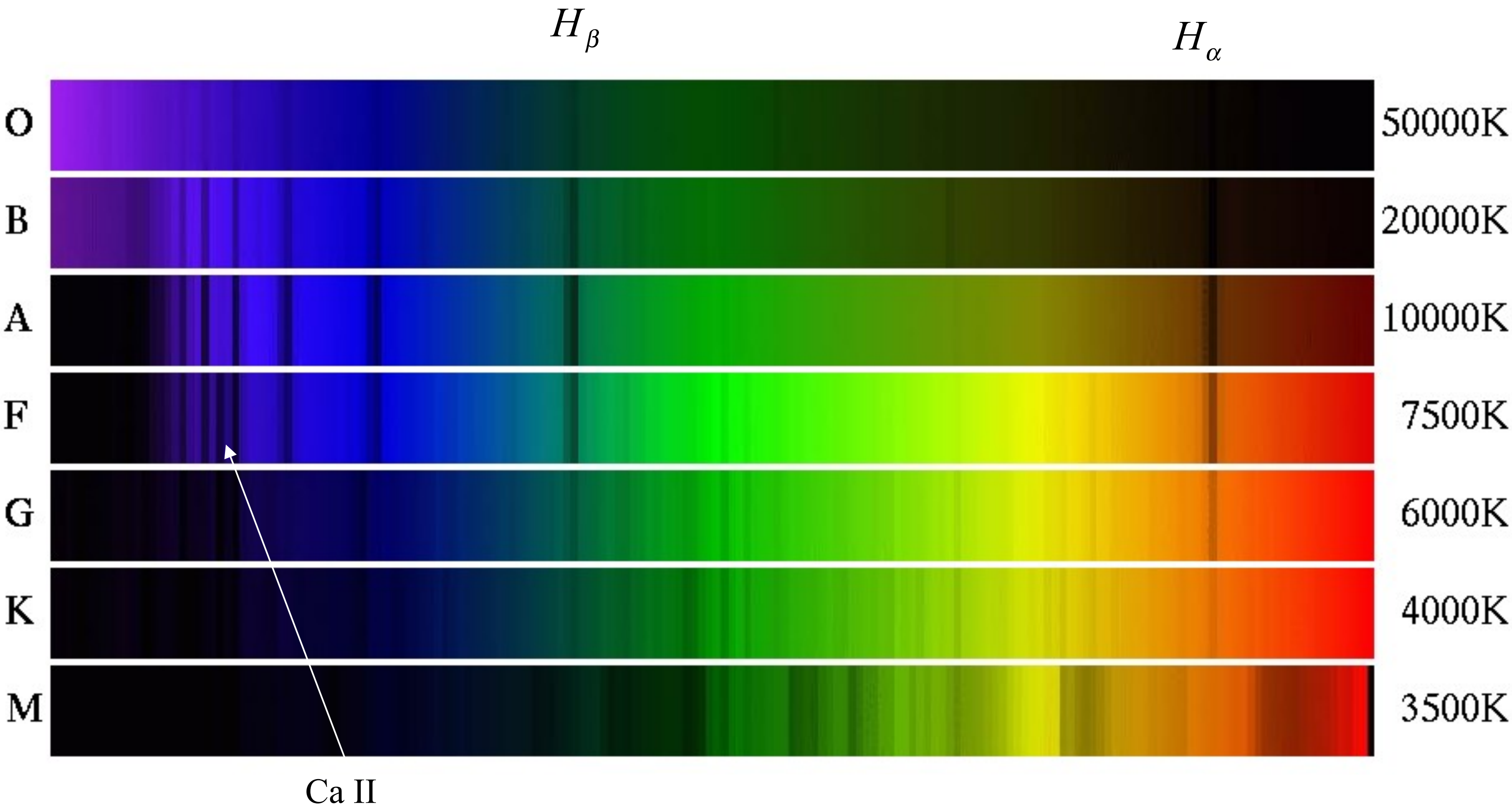


Intensity
Si I
Si II
Si III
et
T →

Some of the stronger lines in stars



Spectral Sequence



The Henry Draper Spectral Sequence

Spectral Type	Principal Characteristics	Spectral Criteria
O	Hottest blue stars Relatively few lines He II dominates	Strong He II lines—in absorption, sometimes emission. He I lines weak, but increasing in strength from O5 to O9. Hydrogen Balmer lines prominent, but weak compared to later types. Lines of Si IV, O III, N III, and C III.
B	Hot blue stars More lines He I dominates	He I lines dominate, with maximum strength at B2; He II lines virtually absent. Hydrogen lines strengthening from B0 to B9. Also Mg II and Si II lines.
A	Blue stars Ionized metal lines Hydrogen dominates	The hydrogen lines reach maximum strength at A0. Lines of ionized metals (Fe II, Si II, Mg II) at maximum strength near A5. Ca II lines strengthening. The lines of neutral metals are appearing weakly.

F	White stars Hydrogen lines declining Neutral metal lines increasing	The hydrogen lines are weakening rapidly, while the H and K lines of Ca II strengthen. Neutral metal (Fe I and Cr I) lines gaining on ionized metal lines by late F.
G	Yellow stars Many metal lines Ca II lines dominate	The hydrogen lines are very weak. The Ca II H and K lines reach maximum strength near G2. Neutral metal (Fe I, Mn I, Ca I) lines strengthening, while ionized metal lines diminish. The molecular G-band of CH becomes strong.
K	Reddish stars Molecular bands appear Neutral metal lines dominate	The hydrogen lines are almost gone. The Ca lines are strong. Neutral metal lines are very prominent. By late K the molecular bands of TiO begin to appear.
M	Coolest red stars Neutral metal lines strong Molecular bands dominate	The neutral metal lines are very strong. Molecular bands are prominent, with the TiO bands dominating the spectrum by M5. Vanadium oxide (VO) bands appear.

- Cannon further refined the spectral classification system by dividing the classes into numbered subclasses:

- For example, A was divided into

  A0 A1 A2 A3 ... A9

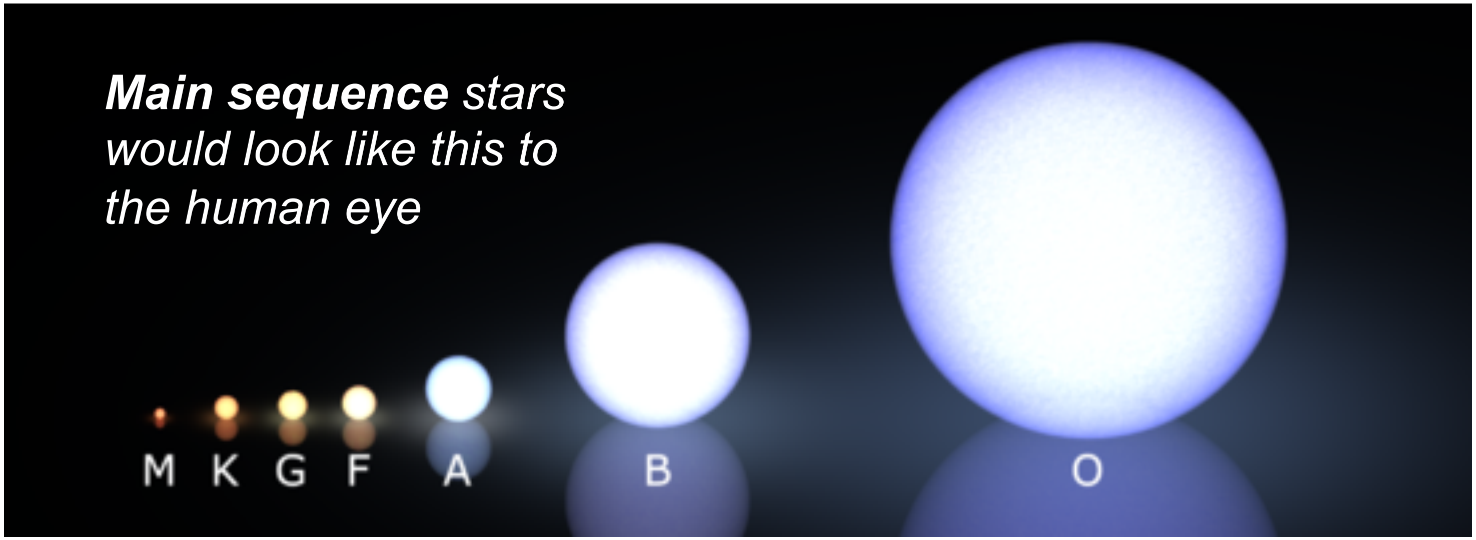
  A0 is hotter than A9

  B9 comes before A0

  F1 comes after A9

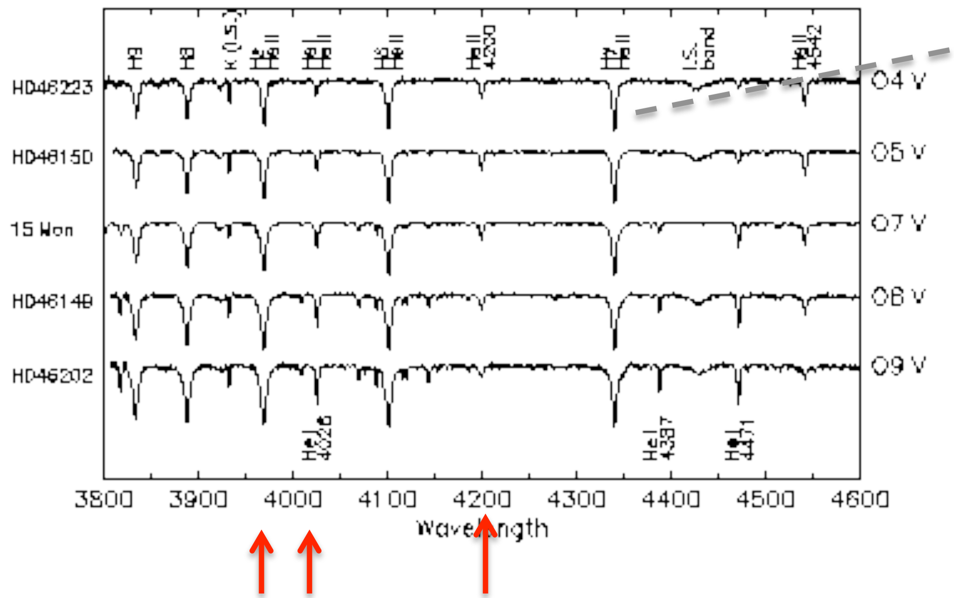
OBAFGKM

			Fraction MS stars solar neighborhood
O	> 25,000 K	Delta Orionis	1/3,000,000
B	11,000 – 25,000	Pleiades brightest	1/800
A	7500 – 11,000	Sirius	1/160
F	6000 – 7500	Canopus	1/133
G	5000 – 6000	Sun	1/13
K	3500 – 5000	Arcturus	1/8
M	< 3500	Proxima Centauri	3/4



http://en.wikipedia.org/wiki/Stellar_classification

Main Sequence O4 – O9



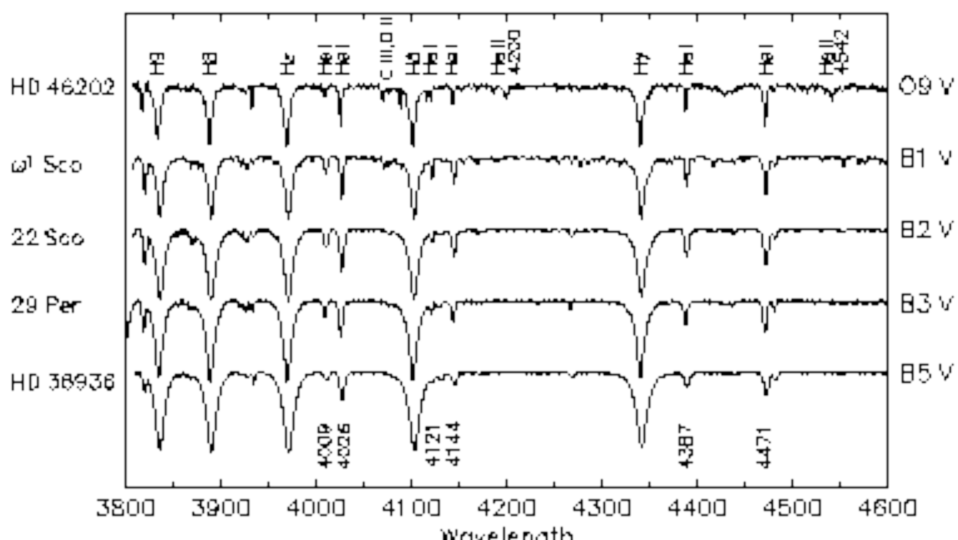
10 → 4 in He II

has same wavelength as

5 → 2 in H I

He II present,
He I increasing
from O4 to O9
H prominent

Main Sequence O9 – B5



He I lines dominate
H increasing in strength

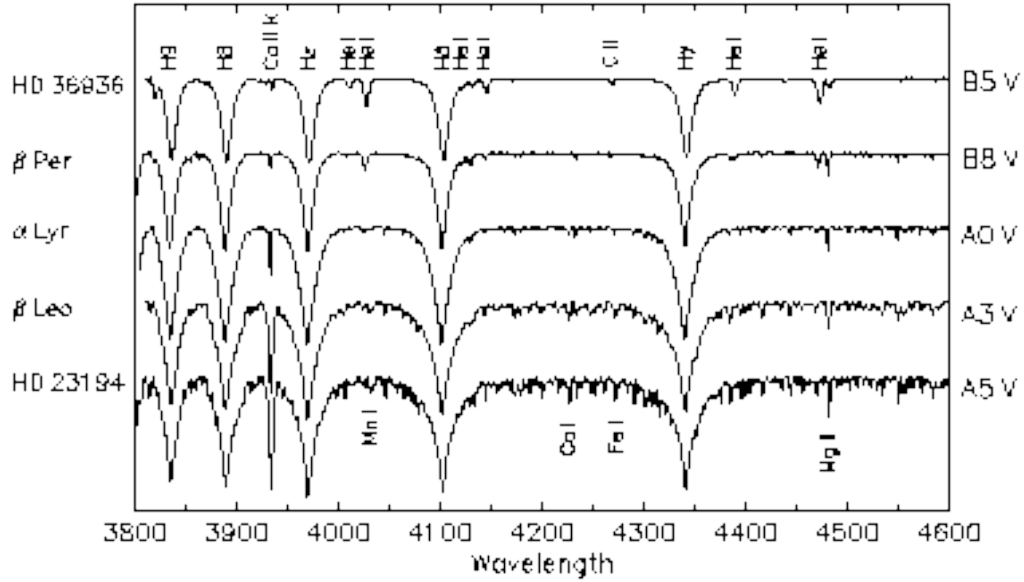
http://nedwww.ipac.caltech.edu/level5/Gray/Gray_contents.html

Balmer Series

Transition	**3 -> 2**	**4 -> 2**	**5 -> 2**	**6 -> 2**	**7-> 2**
Name	H_α	H_β	H_γ	H_δ	H_ε
Wavelength	6563	4861	4341	4102	3970
Color	Red	Blue-green	Violet	Violet	Ultra-violet

Main Sequence B5 — A5

H lines reach maximum strength. Ca II growing. Fe II, Si II, Mg II reach

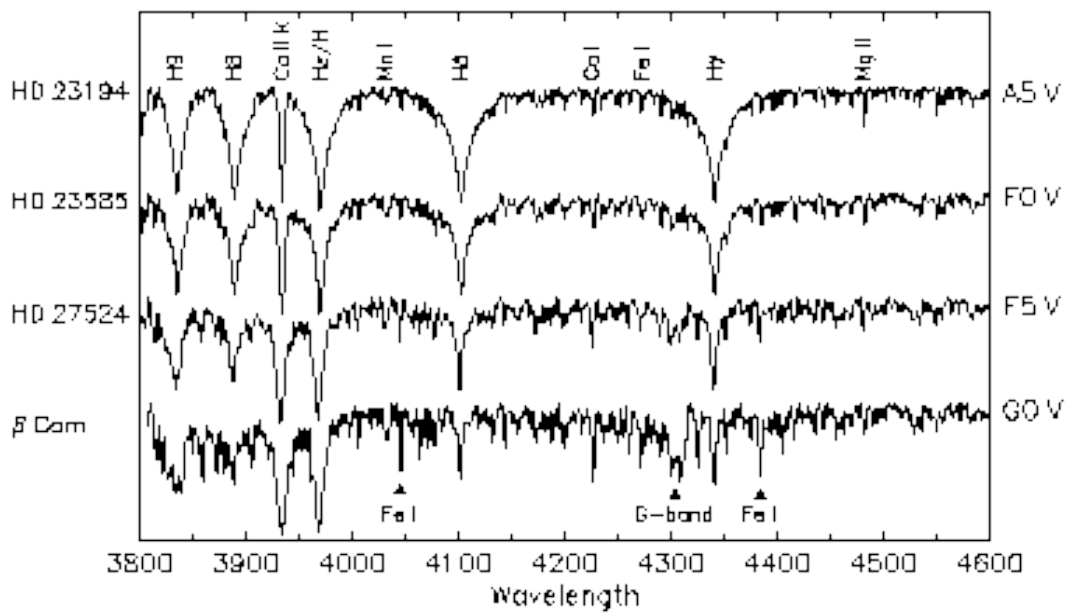


$H_\gamma = 4341\,A$

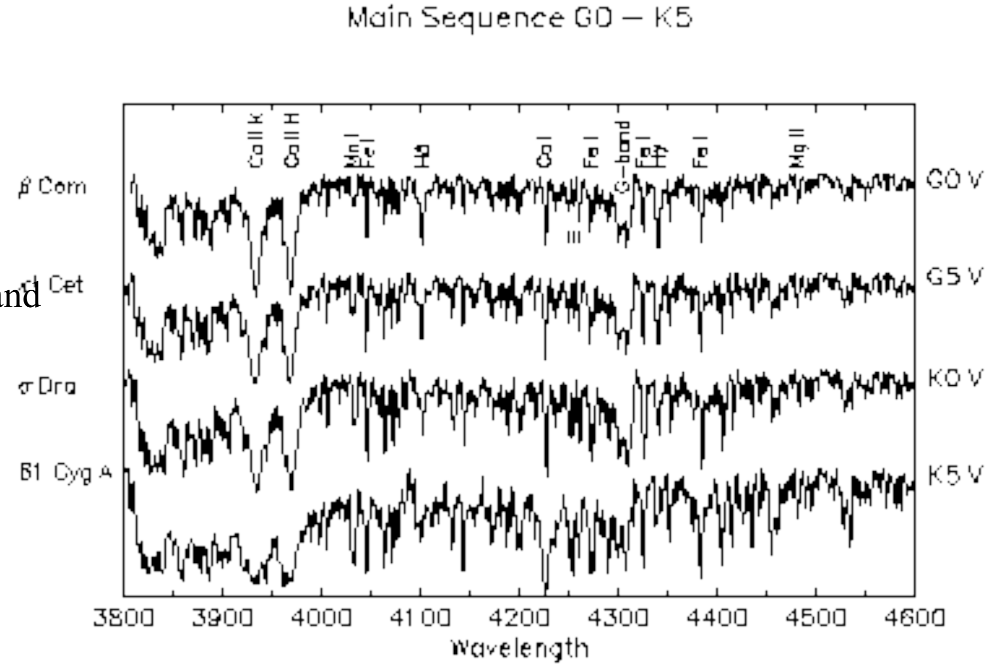
$H_\delta = 4102\ A$

Main Sequence A5 — G0

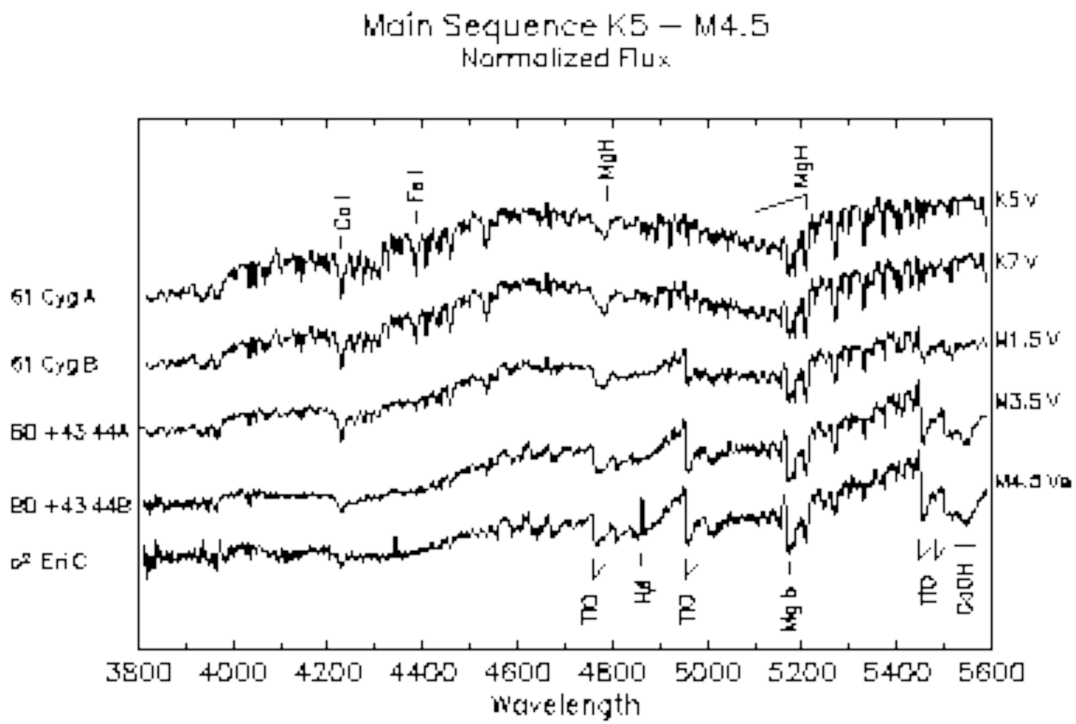
H lines start to decrease in strength. Ca II strong. Fe I growing in strength. Mg II decreasing.



Ca II lines strongest, H lines weak, neutral metal lines strong. G-band of CH is strong.



H lines weak. Lines of neutral metals present but weakening. Major characteristic is bands from molecules like TiO and MgH



DISTINGUISHING MAIN SEQUENCE STARS FROM RED GIANTS OF THE SAME COLOR

The surface gravity

$$g = \frac{GM}{R^2}$$

of a star is clearly larger for a smaller radius (if M is constant)

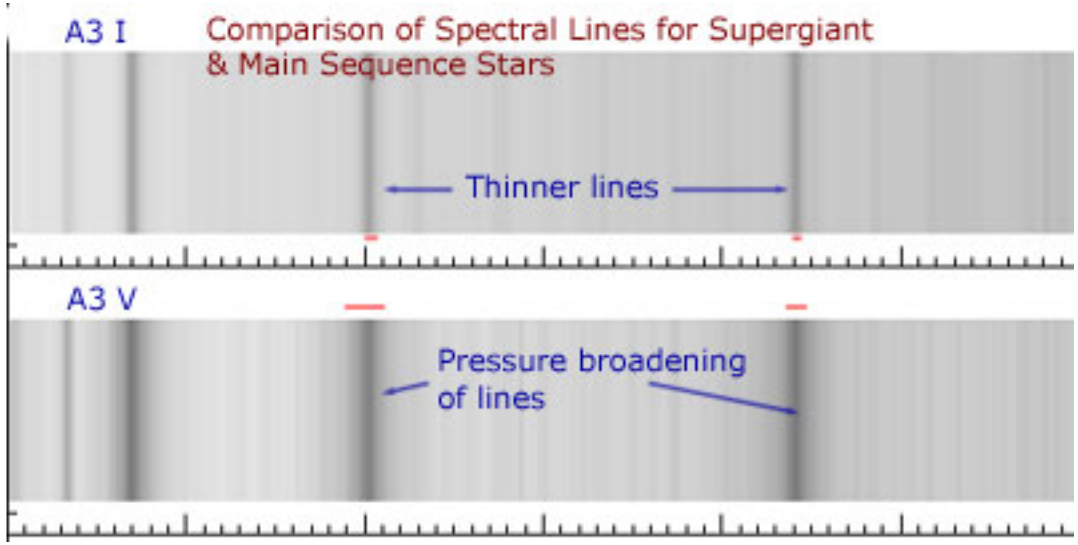
To support itself against this higher gravity, a the stellar photosphere must have a larger pressure. As we shall see later for an ideal gas

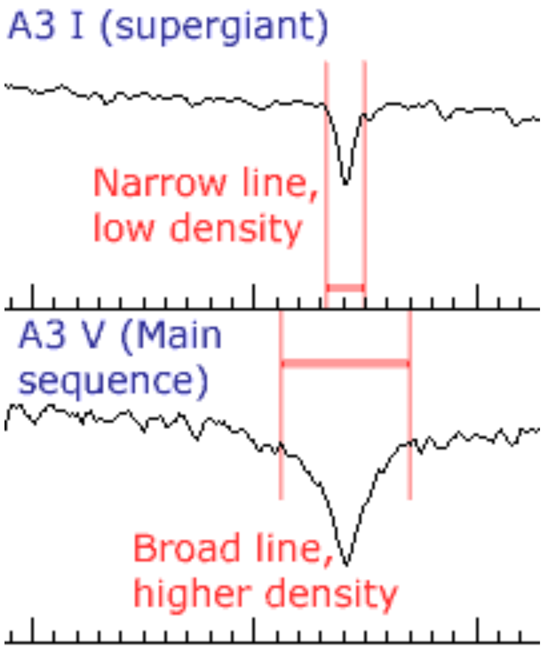
$$P = n\,k\,T$$

where n is the number density and T is the temperature. If two stars have the same temperature, T, the one with the higher pressure (smaller radius) will have the larger n, i.e., its atoms will be more closely crowded together. This has two effects:

1) At a greater density (and the same T) a gas is less ionized

2) If the density is high, the electrons in one atom "feel" the presence of other nearby nuclei. This makes their binding energy less certain. This spreading of the energy level is called "Stark broadening"

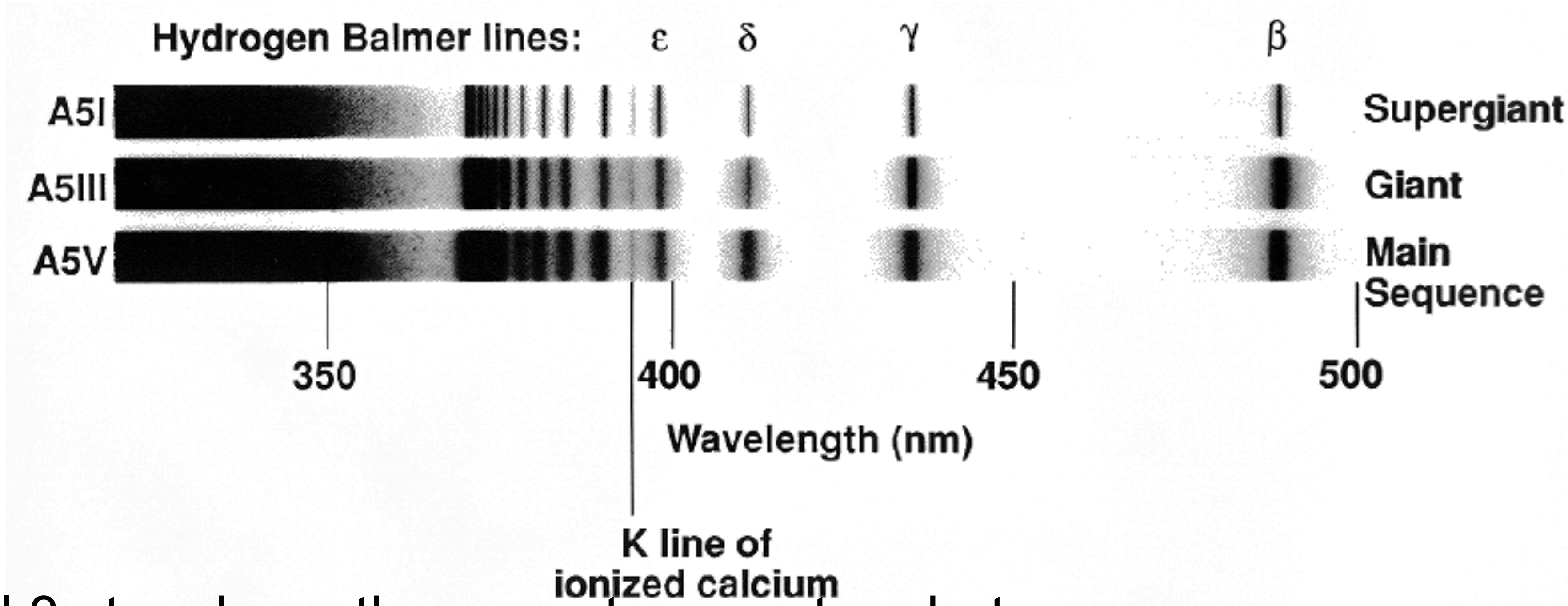




Note: Surface gravity on the main sequence is higher for *lower* mass stars

$$R \propto M^{0.65}$$

$$\frac{GM}{R^2} \text{ decreases with increasing M}$$



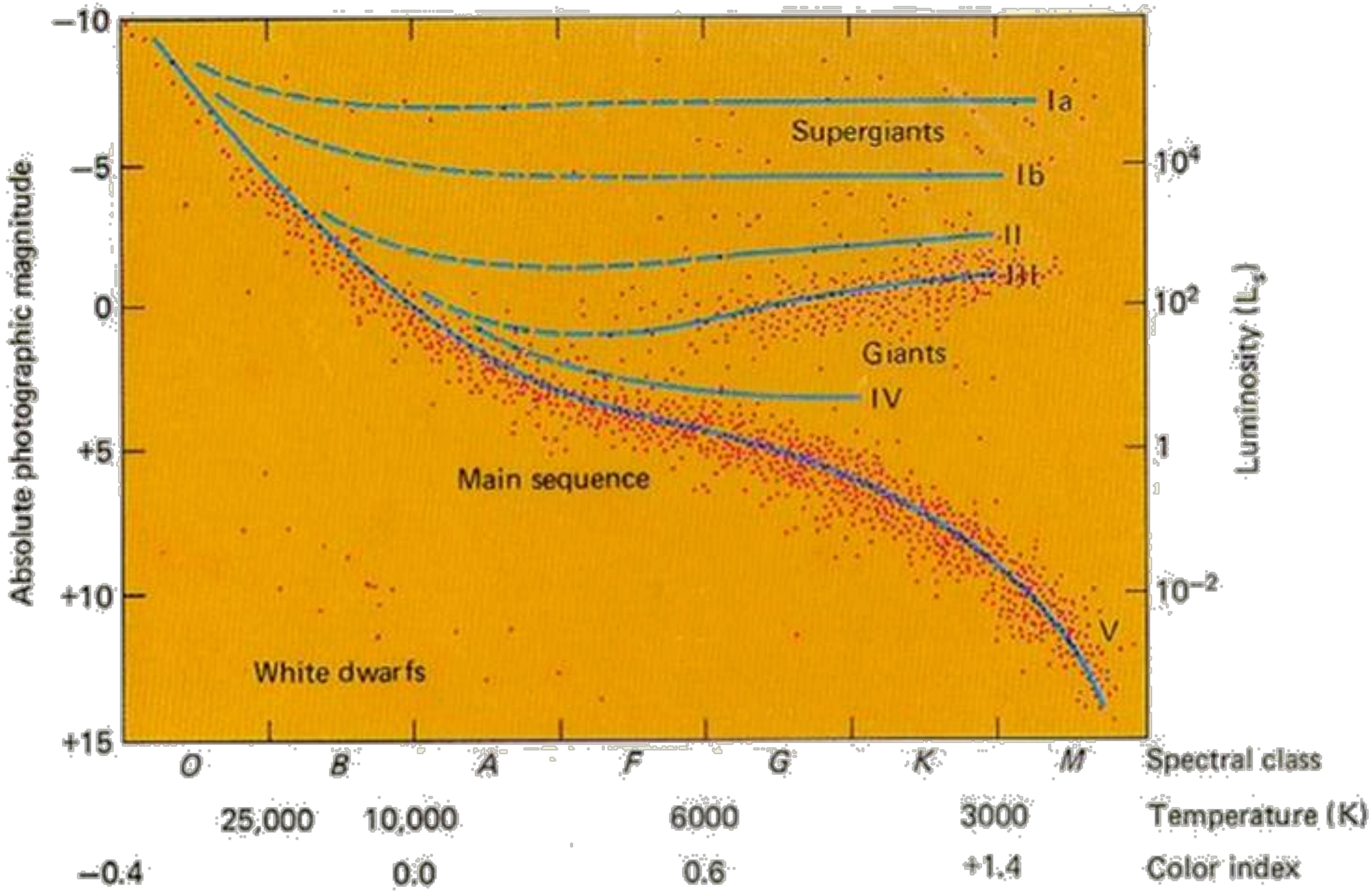
All 3 stars have the same temperature but,

- The supergiants have the narrowest absorption lines
- Small Main-Sequence stars have the broadest lines
- Giants are intermediate in line width and radius

- In 1943, Morgan & Keenan added the *Luminosity Class* as a second classification parameter:
  - Ia = Bright Supergiants
  - Ib = Supergiants
  - II = Bright Giants
  - III = Giants
  - IV = Subgiants
  - V = Main sequence

And so the sun is a G2-V star

Luminosity Classes



71. Abell/Morrison/Wolff: EXPLORATION OF THE UNIVERSE, 6/E