Lecture 11

Neutrino Losses and Advanced Stages of Stellar Evolution - I

The late stages (> helium burning) of evolution in massive stars are characterized by huge luminosities, carried away predominantly by neutrinos, and consequently by short time scales.

The nuclear physics can become quite complicated because of the presence of many species and occurrence of many reactions.

Woosley, Heger, and Weaver (2002)
Rev. Mod. Phys., **74**, 1016.

Thermal Neutrino Processes

Fowler and Hoyle, *ApJS*, **2**, 201, (1964)
Chiu in *Stellar Physics* p 259-282
Beaudet, Petrosian and Salpeter, *ApJ*, **150**, 978, (1967)
Itoh et al, *ApJ*, **339**, 354 (1989)
Itoh et al, *ApJS*, **102**, 411, (1996)

In nature, both of the following equations follow all the necessary conservation laws:

$$\nu_e + e^- \rightarrow \nu_e + e^-$$

$$e^+ + e^- \rightarrow \nu_e + \overline{\nu_e}$$

In about 1970 it was realized that *neutral currents* could lead to additional reactions and modifications of the rates of old ones. Where one had ν_e, one could now have ν_e, ν_μ, or ν_τ. (Dicus, Phys. Rev D, 6, 941, (1972)).

Thus, with a different coupling constant

$$e^+ + e^- \rightarrow \nu_\mu + \bar{\nu}_\mu \text{ or } \nu_\tau + \bar{\nu}_\tau$$

also

$$\nu + {}^Z A \rightarrow \nu + {}^Z A \quad \text{for } \nu_e, \nu_\mu, \nu_\tau$$ coherent process

Stellar Neutrino Energy Losses

(see Clayton p. 259ff, especially 272ff; Fowler and Hoyle 1964, eq. 3)

1) Pair annihilation - dominant in massive stars

$$kT \geq 10\% \ m_e c^2 \quad (T_9 > 0.5)$$

$$\boxed{e^+ + e^- \rightleftharpoons \textit{radiation} \rightarrow \nu_e + \bar{\nu}_e}$$

about 10^{-19} branch

$$\sigma_\pm = \frac{G_W^2 c}{3\pi v}\left(\frac{\hbar}{m_e c}\right)^2\left[\left(\frac{E}{m_e c^2}\right)^2 - 1\right]$$

E includes rest mass

$\hbar / m_e c$ = Compton wavelength/2π of e^- = 3.86×10^{-11} cm

$E^2 = m_e^2 c^4 + p^2 c^2 = \gamma^2 m_e^2 c^4$ $\quad p = mv$; v is the relative velocity of the pair ~ c

$G_W = 3.00 \times 10^{-12}$ (dimensionless)

$$\sigma_\pm = 1.42\left(\frac{c}{v}\right)\left[\left(\frac{E}{m_e c^2}\right)^2 - 1\right] \times 10^{-45}\ \text{cm}^2$$

Want energy loss per cm^3 per second. Integrate over thermal distribution of e$^+$ and e$^-$ velocities. These have, in general, a Fermi-Dirac distribution.

$$\boxed{P_\pm = n_{e^+} n_{e^-} < \sigma v E >}$$

E = total energy including rest mass

$$n_\pm = \frac{1}{\pi^2}\left(\frac{m_e c}{\hbar}\right)^3 \int_{-\infty}^{+\infty} \frac{W(W^2-1)\,dW}{\exp(\theta W \pm \phi)-1}$$

Fermi Integral

-

$$\theta = \frac{m_e c^2}{kT} = 5.93/T_9 \qquad W = \frac{E}{m_e c^2} \quad \text{c/m energy}$$

ϕ = Chemical potential/kT

(determined by the condition that

$$n_- - n_+ = n_e\ (\text{matter}) = \rho N_A Y_e)$$

Clayton (Chap 4) and Lang in Astrophysical Formulae give some approximations (not corrected for neutral currents)

$$(NDNR) \quad \boxed{P_\pm \approx 4.9\times10^{18}\ T_9^3 \exp(\underbrace{-11.86/T_9}_{2m_e c^2/kT})\ \text{erg cm}^{-3}\ \text{s}^{-1}} \quad T_9 < 2$$

$$(NDR) \quad \boxed{P_\pm \approx 4.6\times10^{15}\ T_9^9\ \text{erg cm}^{-3}\ \text{s}^{-1}} \quad T_9 > 3, \text{ but not too bad at } T_9 > 2$$

(better is 3.2×10^{15})

Note origin of T^9 :

If $n_\pm$ is relativistic, $n_\pm \propto T^3$ (like radiation)

$$\sigma \propto \frac{E^2}{v} \propto \frac{(kT)^2}{v} \qquad \text{(2 pages back)}$$

energy carried per reaction ~ kT

$$P_\pm \approx \underset{n_+ n_-}{\left(T^6\right)}\underset{\sigma v}{\left(T^2\right)}\underset{E}{\left(T\right)} = T^9$$

v cancels 1/v in σ

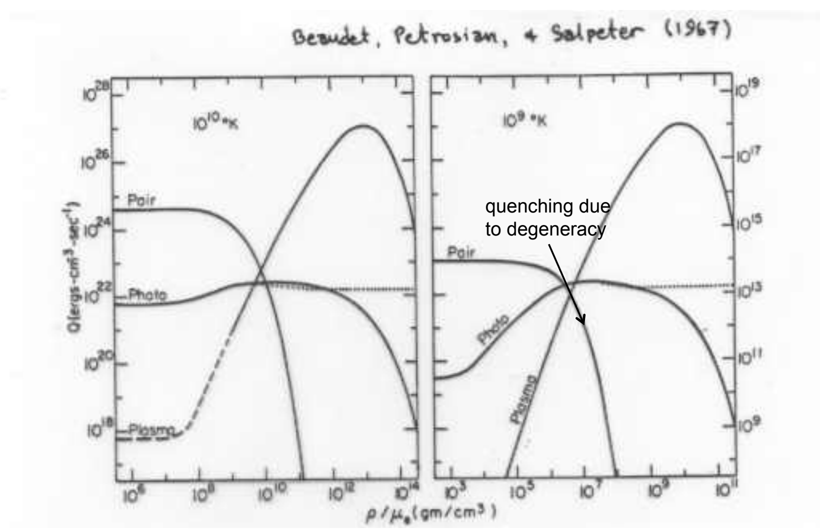
These formulae are very crude; for more accurate results use subroutine neut01.f on the class website.

More frequently we use the energy loss rate per **gram** per second

$$\varepsilon_\nu = \frac{P_\nu}{\rho} \ \text{erg gm}^{-1}\ \text{s}^{-1}$$

In the non-degenerate limit ε_ν from pair annihilation declines as ρ^{-1}.

In degenerate situations, the filling of phase space suppresses the creation of electron-positron pairs and the loss rate plummets. Usually pair annihilation neutrino emission dominates other processes when the matter is non-degenerate. This includes most of the advanced stages of stellar evolution (especially when electron capture on nuclei is negligible).



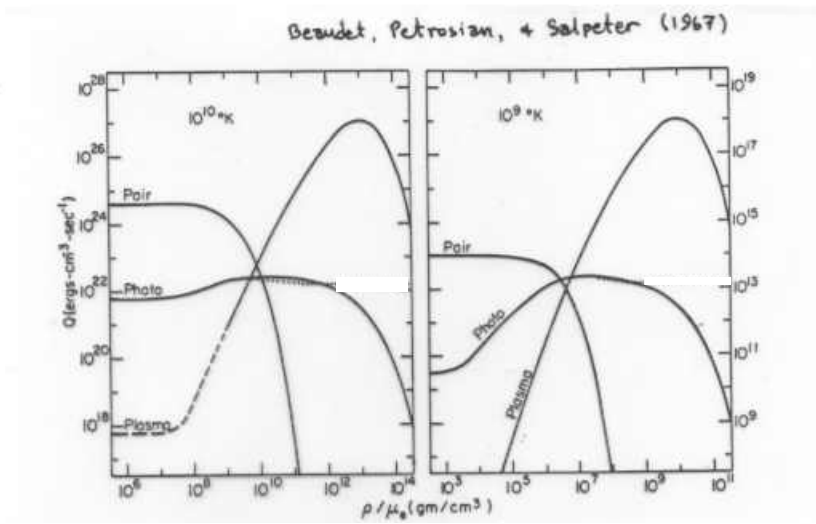
note that units here are erg cm^{-3} s^{-1} not erg g^{-1} s^{-1}

2) Photoneutrino process: (Clayton p. 280)

$$\boxed{e^- + \gamma \rightarrow e^- + \nu + \bar{\nu}}$$

Analogue of Compton scattering with the outgoing photon replaced by a neutrino pair. The electron absorbs the extra momentum. This process can be marginally significant during helium and carbon burning.

When non-degenerate and non-relativistic P_{photo} is proportional to the density (because it depends on the electron abundance) and $\varepsilon_{\nu,photo}$ is independent of the density. At high density, degeneracy blocks the phase space for the outgoing electron. Also for relativistic electrons the density dependence is weaker.



relativistic ~non-relativistic

Usually pair production (hi T) or plasma losses (hi ρ) more important

3) Plasma Neutrino Process: (Clayton 275ff)

This process is important at high densities where the plasma frequency is high and $\hbar\omega_{\text{plasma}}$ can be comparable to kT. This limits its applicability to essentially white dwarfs, and to a lesser extent, the evolved cores of massive stars. It is favored in degenerate environments.

A "plasmon" is a quantized collective charge oscillation in an ionized gas. For our purposes it behaves like a photon with rest mass.

Consider a neutral plasma, consisting of a gas of positively charged ions and negatively charged electrons. If one displaces by a tiny amount all of the electrons with respect to the ions, the Coulomb force pulls back, acting as a restoring force.

If the electrons are cold it is possible to show that the plasma oscillates at the plasma frequency.

ND $$\omega_p = \left(\frac{4\pi n_e e^2}{m_e}\right)^{1/2} = 5.6\times10^4\, n_e^{1/2}$$ increases with density

D $$\omega_p = \left(\frac{4\pi n_e e^2}{m_e}\right)^{1/2}\left[1+\underbrace{\left(\frac{\hbar}{m_e c}\right)^2\left(3\pi^2 n_e\right)^{2/3}}_{2\varepsilon_F/m_e c^2}\right]^{-1/2}$$

suppression for degeneracy

$$F = m_e a = \frac{4\pi r^3 n_e e}{3r^2} e$$

$$\tau = \left(\frac{2r}{a}\right)^{1/2} = \left(\frac{2r\,m3}{4\pi r\, n_e e^2}\right)^{1/2} = \frac{2\pi}{\omega}$$

$$\Rightarrow \omega \propto \left(\frac{n_e e^2}{m}\right)^{1/2}$$

A photon of any energy in a vacuum cannot decay into e^+ and e^- because such a decay would not simultaneously satisfy the conservation of energy and momentum (e.g., a photon that had energy just equal to 2 electron masses, $h\nu = 2\ m_e c^2$, would also have momentum $h\nu/c = 2m_e c$, but the electron and positron that are created, at threshold, would have no kinetic energy, hence no momentum. Such a decay is only allowed when the photon couples to matter that can absorb the excess momentum.

The common case is a γ-ray of over 1.02 MeV passing near a nucleus, but the photon can also acquire an effective mass by propagating through a plasma.

$$\gamma_{plasmon} \rightarrow e^+ + e^-$$

In a plasma the dispersion relation between the wave number, $K = 2\pi/\lambda$, and the frequency is not linear. For a photon of frequency ω

$$\omega = (K^2c^2 + \omega_p^2)^{1/2}$$

below the plasma frequency, a wave cannot propagate.

and the photon energy

$$E_\gamma = \hbar\omega = \left(K^2\hbar^2c^2 + \frac{(\hbar\omega_p)^2}{c^4}c^4\right)^{1/2}$$

analogous to $E^2 = p^2c^2 + m^2c^4$ if $m_{eff} = \frac{\hbar\omega_p}{c^2}$

$$p = \frac{h}{\lambda}$$

An electromagnetic wave propagating through a plasma has an excess energy above that implied by its momentum. This excess energy is available for decay.

For moderate values of temperature and density, raising the density implies more energy in the oscillations and raising the temperature excites more oscillations. Hence the loss rate increases with temperature and density.

However, once the density becomes so high that, for a given temperature $\hbar\omega_p > kT$, raising the density still further freezes out the oscillations. The thermal plasma no longer has enough energy to exite them. The loss rate plummets exponentially.

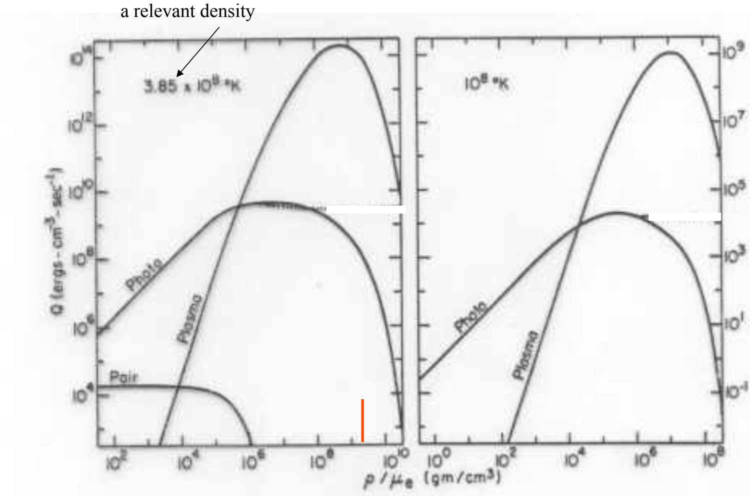
a) $\hbar\omega_p \le kT$

$$P_{plasma} \approx 7.4\times10^{21}\left(\frac{\hbar\omega_p}{m_ec^2}\right)^6\left(\frac{kT}{m_ec^2}\right)^3 \text{ erg cm}^{-3}\text{ s}^{-1} \propto \rho^3 T^3$$

b) $\hbar\omega_p >> kT$

$$P_{plasma} \approx 3.3\times10^{21}\left(\frac{\hbar\omega_p}{m_ec^2}\right)^{7.5}\left(\frac{kT}{m_ec^2}\right)^{3/2} \exp(-\hbar\omega_p/kT) \text{ erg cm}^{-3}\text{ s}^{-1}$$

This is a relevant temperature for Type Ia supernovae and the red line a relevant density



4) <u>Ordinary weak interactions</u> – neutrinos from the decay of unstable nuclei

- Beta-decay
- Electron capture
- Positron emission

Electron capture – and to a lesser extent beta-decay can be very important in the final stages of stellar evolution – especially during silicon burning and core collapse.

Typically these are included by studying each nucleus individually, its excited state distribution, distribution of weak strength, etc. The results are then published as fitting functions at f(T,r).

Fuller, Fowler, & Newman, *ApJS*, **48**, 27 (1982a)
ApJ, **252**, 715, (1982b)
ApJ, **293**, 1, (1985)
Oda et al, *Atom. Data and Nuc. Data Tables*, **56**, 231, (1996)
Langanke & Martinez-Pinedo, *Nuc Phys A*, **673**, 481 (2000)

The late stages of stellar evolution are accelerated by (pair) neutrino losses.

Advanced Nuclear Burning Stages
(e.g., 20 solar masses)

Fuel	Main Product	Secondary Products	Temp (10^9 K)	Time (yr)
H	He	^{14}N	0.02	10^7
He	C,O	^{18}O,^{22}Ne s- process	0.2	10^6
C	Ne, Mg	Na	0.8	10^3
Ne	O, Mg	Al, P	1.5	3
O	Si, S	Cl, Ar K, Ca	2.0	0.8
Si	Fe	Ti, V, Cr Mn, Co, Ni	3.5	1 week

<u>Carbon Burning</u>

Approximate initial conditions:

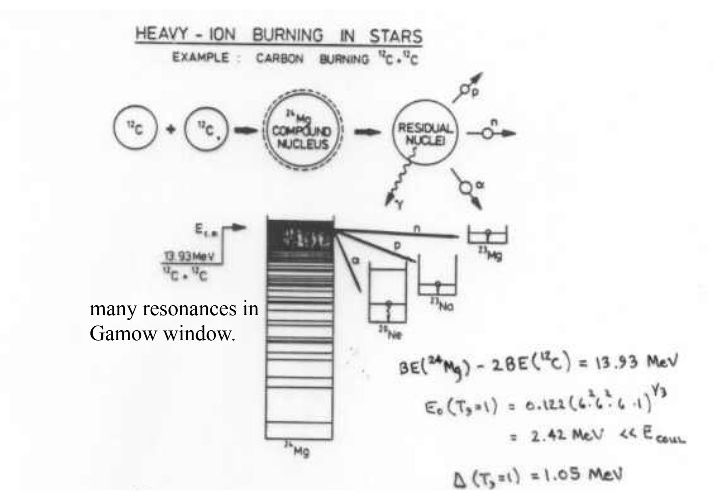
As we shall see, the temperature at which carbon burns in a massive star is determined by a state of *balanced power* between neutrino losses by the pair process and nuclear energy generation. This gives 8 x 10^8 K for carbon core burning. Burning in a shell is usually a little hotter at each step, about 1.0 x 10^9K for carbon burning.

Assuming that T^3/ρ scaling persists at the center, and that helium burned at 2 x 10^8 K and 1000 gm cm^{-3}, this implies a carbon burning density around a few x 10^5 gm cm^{-3}.

The initial composition is the ashes of helium burning, chiefly C and O in an approximate 1 : 4 ratio (less carbon in more massive stars).

There are also many other elements present in trace amounts:

- ^{22}Ne, ^{25,26}Mg from the processing of CNO elements in He-burning
- The light s-process
- Traces of other heavy elements present in the star since birth
- Up to ~1% ^{20}Ne from ^{16}O(α,γ)^{20}Ne during He-burning



Measured to about 2.5 MeV and S-factor is overall smooth but shows poorly understood broad “structures” at the factor of 2 level. See Rolfs and Rodney, p 419 ff - alpha cluster? Not seen in ^{16}O + ^{16}O

Principal nuclear reaction

$$\begin{aligned} {}^{12}\mathrm{C} + {}^{12}\mathrm{C} \rightarrow {}^{24}\mathrm{Mg}^* &\rightarrow {}^{23}\mathrm{Mg} + n - 2.62\ \mathrm{MeV} \\ &\rightarrow {}^{20}\mathrm{Ne} + \alpha + 4.62\ \mathrm{MeV} \\ &\rightarrow {}^{23}\mathrm{Na} + p + 2.24\ \mathrm{MeV} \end{aligned}$$

^{12}C + ^{12}C → ^{23}Mg + n

T_9	B_n	T_9	B_n	T_9	B_n
0.7	1.6(-5)	1.0	1.1(-3)	2	0.024
0.8	1.1(-4)	1.2	4.0(-3)	3	0.042
0.9	4.0(-4)	1.4	8.8(-3)	5	0.054

$B_P = 0.45 - B_n/2 \qquad B_\alpha = 0.55 - B_n/2$

Dayras, Switkowsky, & Woosley, Nuc Phys A, 279, 70, (1979)

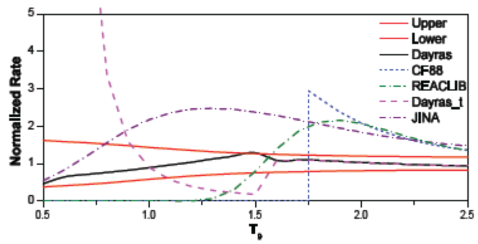
Many important secondary reactions:

^{20}Ne(α,γ)^{24}Mg	^{23}Na(α,p)^{26}Mg	^{26}Mg(p,γ)^{27}Al
^{23}Na(p,γ)^{24}Mg	^{23}Na(p,α)^{20}Ne	^{25}Mg(p,γ)^{26}Al
^{22}Ne(α,n)^{25}Mg	^{25}Mg(α,n)^{28}Si	^{23}Mg(n,p)^{23}Na
^{25}Mg(n,γ)^{26}Mg	^{23}Mg(e$^+\nu$)^{23}Na	^{21}Na(e$^+\nu$)^{21}Ne
^{20}Ne(p,g)^{21}Na	^{21}Ne(p,g)^{22}Na	^{22}Na(e$^+\nu$)^{22}Ne

and dozens (hundreds?) more

Bucher et al (ApJ, in press 2015) [I am a coauthor]

$$\beta_n = 0.11954 \exp\left[-\left(\frac{0.16446}{T_9{}^3} + \frac{2.57495}{T_9{}^2} + \frac{1.94145}{T_9}\right)\right]$$
$$(T_9 \le 1.5)$$
$$= 0.2212\,[1 - \exp(-0.13597\,T_9 + 0.158)]$$
$$(1.5 \le T_9 \le 2.5)$$
$$= 0.048811\,[1 - \exp(-2.1124\,T_9 + 3.8791)]$$
$$(2.5 \le T_9 \le 5.0)$$
$$= 0.04875 \quad (T_9 > 5.0)$$



G. R. CAUGHLAN and W. A. FOWLER   Thermonuclear Reaction Rates
1988

TABLE III. Thermonuclear Reaction Rates $1 \le Z \le 14$
See page 291 for Explanation of Tables
TABLE III 27

T9	C12C12	C12016	016016	T9
0.110	7.42E-50		0.00E+00	0.110
0.120	1.04E-47		0.00E+00	0.120
0.130	8.72E-46		0.00F:00	0.130
0.140	4.71E-44		0.00E+00	0.140
0.150	1.77E-42		0.00E+00	0.150
0.160	4.86E-41		0.00E+00	0.160
0.180	1.71E-38		0.00E+00	0.180
0.200	2.66E-36		0.00E+00	0.200
0.250	6.52E-32		0.00E+00	0.250
0.300	1.44E-28		0.00E+00	0.300
0.350	6.66E-26		1.60E-47	0.350
0.400	1.05E-23		6.18E-44	0.400
0.450	7.49E-22		6.61E-41	0.450
0.500	2.94E-20	3.86E-28	2.67E-38	0.500
0.600	1.23E-17	8.15E-25	5.21E-34	0.600
0.700	1.51E-15	3.62E-22	1.36E-30	0.700
0.800	7.89E-14	5.46E-20	8.87E-28	0.800
0.900	2.21E-12	3.75E-18	2.08E-25	0.900
1.000	3.86E-11	1.42E-16	2.27E-23	1.000
1.250	1.16E-08	2.02E-13	2.60E-19	1.250
1.500	8.66E-07	5.02E-11	3.04E-16	1.500
1.750	2.64E-05	4.03E-09	8.15E-14	1.750
2.000	4.33E-04	1.49E-07	7.81E-12	2.000
2.500	3.34E-02	4.16E-05	9.09E-09	2.500
3.000	8.64E-01	2.82E-03	1.73E-06	3.000
3.500	1.10E+01	7.22E-02	1.01E-04	3.500
4.000	8.58E+01	8.92E-01	2.61E-03	4.000
5.000	1.91E+03	3.10E+01	3.33E-01	5.000
6.000	1.73E+04	3.34E+02	9.93E+00	6.000
7.000	8.53E+04	1.88E+03	1.15E+02	7.000
8.000	2.66E+05	7.04E+03	6.82E+02	8.000
9.000	5.76E+05	1.99E+04	2.50E+03	9.000
10.000	9.13E+05	4.59E+04	6.30E+03	10.000

$$\tau = 4.248\left(\frac{Z_1^2 Z_2^2 \hat{A}}{T_9}\right)^{1/3}$$
$$= 4.248\left(\frac{(36)(36)\left(\frac{12\bullet 12}{24}\right)}{1}\right)^{1/3}$$
$$= 84.16\ T_9^{-1/3}$$

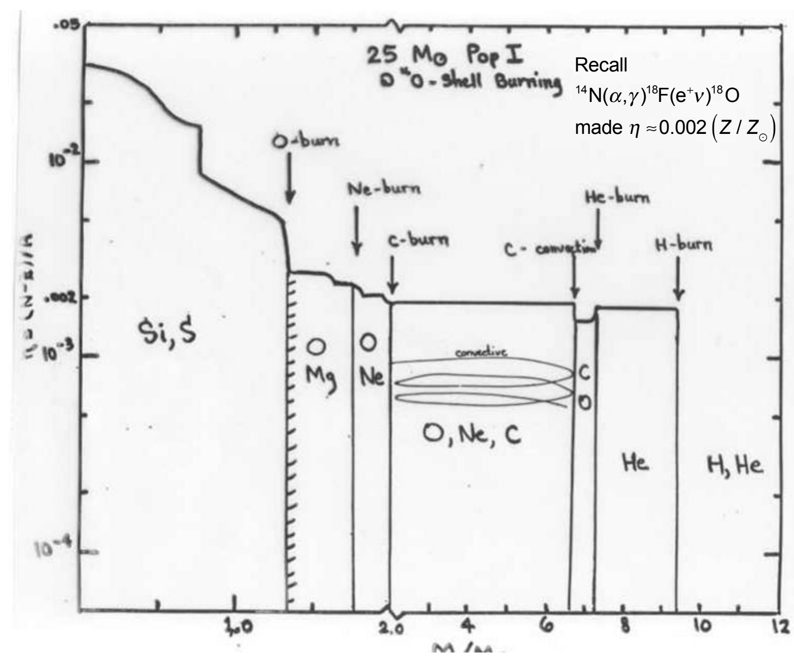
$$\lambda \propto T^n \qquad n = \frac{\tau - 2}{3}$$
$$= 27 \text{ at } T_9 = 1$$
$$= 30 \text{ at } T_9 = 0.8$$

$$\boxed{\lambda \approx 3.9\times 10^{-11}\, T_9^{29}\ \mathrm{cm^3\,gm^{-1}\,s^{-1}}}$$

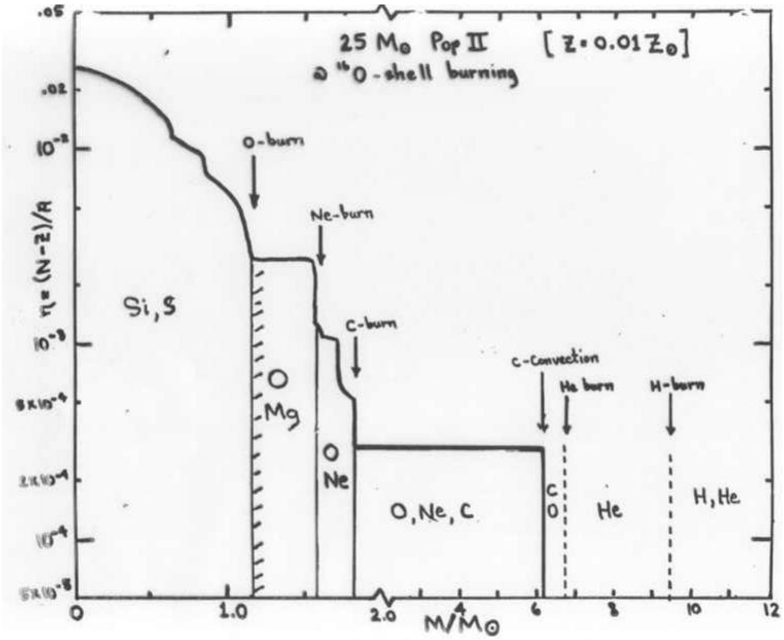
There are also some important weak interactions that can change the neutron excess η.

- The neutron branch of $^{12}C + {}^{12}C$ itself makes ^{23}Mg.
  At lower temperature this decays by $^{23}Mg(e^+\nu)^{23}Na$.
  At higher temperature it is destroyed by $^{23}Mg(n,p)^{23}Na$.
  The former changes η; the latter does not, so there is some temperature, hence mass dependence of the result.
- $^{20}Ne(p,\gamma)^{21}Na(e^+\nu)^{21}Ne$
- $^{21}Ne(p,\gamma)^{22}Na(e^+\nu)^{22}Ne$

Together these reactions can add - a little - to the neutron excess that was created in helium burning by $^{14}N(\alpha,\gamma)^{18}F(e^+\nu)^{18}O$ or, in stars of low metallicity they can create a neutron excess where none existed before.



Recall
$^{14}N(\alpha,\gamma)^{18}F(e^+\nu)^{18}O$
made $\eta \approx 0.002\,(Z/Z_\odot)$



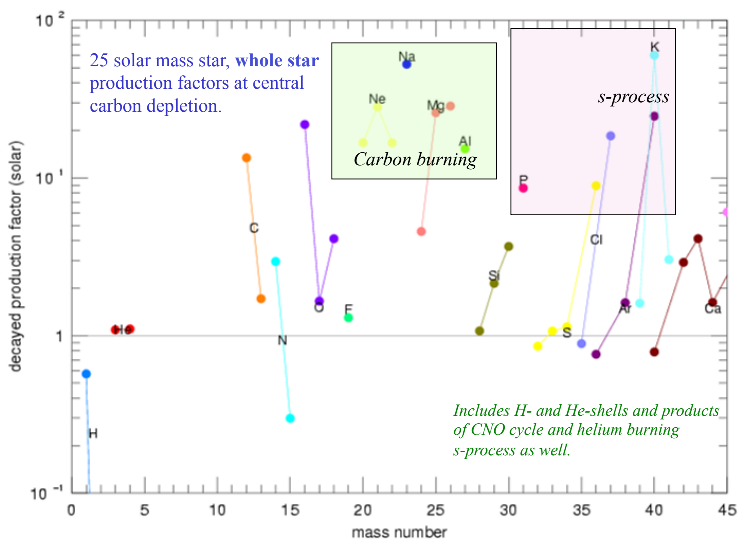
Principal Nucleosynthesis in carbon burning:

^{20,21}Ne, ^{23}Na, ^{24,25,26}Mg, $^{(26),27}$Al, and to a lesser extent, ^{29,30}Si, ^{31}P

The ^{16}O initially present at carbon ignition essentially survives unscathed. There are also residual products from helium burning – the s-process, and further out in the star H- and He-burning continue.

A typical composition going into neon burning – major abundances only would be

70% ^{16}O, 25% ^{20}Ne, 5% ^{24}Mg, and traces of heavier elements



Several carbon shell burning episodes remain

Model s25a28

D. Energy Generation

Suppose we make ^{20}Ne and ^{24}Mg in a 3:1 ratio (approximately solar)

$$7\left(^{12}\mathrm{C}\right) \rightarrow 3\left(^{20}\mathrm{Ne}\right) + {}^{24}\mathrm{Mg}$$

$$\varepsilon_{nuc} = 9.65\times 10^{17}\sum\left(\frac{dY_i}{dt}\right)BE_i \text{ erg g}^{-1}\text{ s}^{-1}$$

$$\frac{dY(^{20}\mathrm{Ne})}{dt} = -\frac{3}{7}\frac{dY(^{12}\mathrm{C})}{dt} \qquad \frac{dY(^{24}\mathrm{Mg})}{dt} = -\frac{1}{7}\frac{dY(^{12}\mathrm{C})}{dt}$$

$$\frac{dY(^{12}\mathrm{C})}{dt} = -\,2\rho Y^2(^{12}\mathrm{C})\lambda_{12,12}/2$$

$$\varepsilon_{nuc} = 9.65\times 10^{17}\left[-\frac{3}{7}(160.646) - \frac{1}{7}(198.258) + 1(92.160)\right]\frac{dY(^{12}\mathrm{C})}{dt}$$

$$= -(9.65\times10^{17})(5.01)\frac{dY(^{12}\mathrm{C})}{dt}$$

$$\boxed{\varepsilon_{nuc} \approx 4.84\times10^{18}\,\rho Y^2(^{12}C)\lambda_{12,12} \text{ erg g}^{-1}\text{ s}^{-1}}$$

where $\lambda_{12,12}$ was given a few pages back.

The total energy released during carbon burning is

$$q_{nuc} = 9.65 \times 10^{17} \sum \Delta Y_i (BE_i)$$

$$\Delta Y_{12} = \frac{1}{12} \Delta X_{12}$$

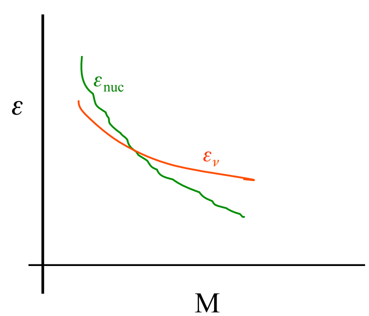
$$\Delta Y_{20} = -\frac{3}{7} \Delta Y_{12} \qquad\qquad q_{nuc} = \frac{9.65 \times 10^{17}}{12} (5.01)\ \Delta X_{12}$$

$$\Delta Y_{24} = -\frac{1}{7} \Delta Y_{12}$$

$$\boxed{q_{nuc} = 4.03 \times 10^{17}\ \Delta X_{12}\ \text{erg g}^{-1}}$$

Since $\Delta X_{12} \ll 1$, this is significantly less than helium burning

For carbon burning u = 2 s = 30
neutrino losses u = 0 s ~ 16



$$\frac{\langle \varepsilon_\nu \rangle}{\langle \varepsilon_{nuc} \rangle} = 1 = \left(\frac{\varepsilon_{\nu o}}{16^{3/2}} \right) / \left(\frac{\varepsilon_{nuc0}}{36^{3/2}} \right)$$

$$\Rightarrow \varepsilon_{nuc0} = 3.4\ \varepsilon_{\nu 0}$$

E. Balanced Power

Averaged over the burning region, which is highly centrally concentrated

$$\langle \varepsilon_{nuc} \rangle \approx \langle \varepsilon_\nu \rangle \quad \text{since} \quad L_\nu = \int \varepsilon_\nu\, dM \gg L_\gamma$$

Neutrino losses in carbon burning are due to pair annihilation. Near $T_9 = 1$ the non-relativistic, non-degenerate formula applies and ε_ν is approximately proportional to T^{16} (at $\rho \sim 10^5$ gm cm^{-3})

Fowler and Hoyle (1964) showed that averaged over an n = 3 polytrope a density and temperature sensitive function has an average:

$$\langle \varepsilon \rangle = \frac{\int \varepsilon\, dM}{\int dM} = \varepsilon_o \frac{3.2}{(3u+s)^{3/2}}$$

where ε_o is the central value of ε, and $\varepsilon \propto \rho^{u-1} T^s$

$\rho = 2 \times 10^5$
$X(^{12}C) = 0.2$

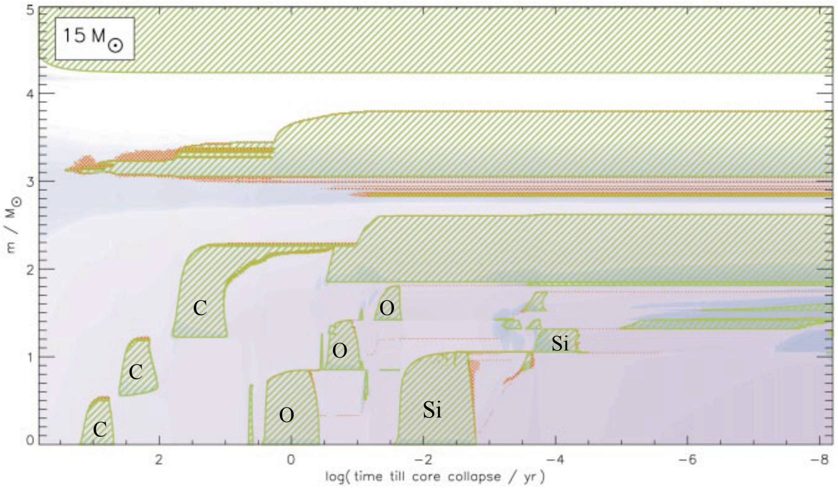
T_9	ε_{nuc} (from formula)	$\varepsilon_{\nu\,total}$ (from sneut01)
0.6	3.4(3)	1.7(5)
0.7	4.0(5)	1.0(6)
0.8	2.2(7)	9.6(6)
0.9	6.0(8)	7.4(7)
1.0	1.0(10)	4.2(8)

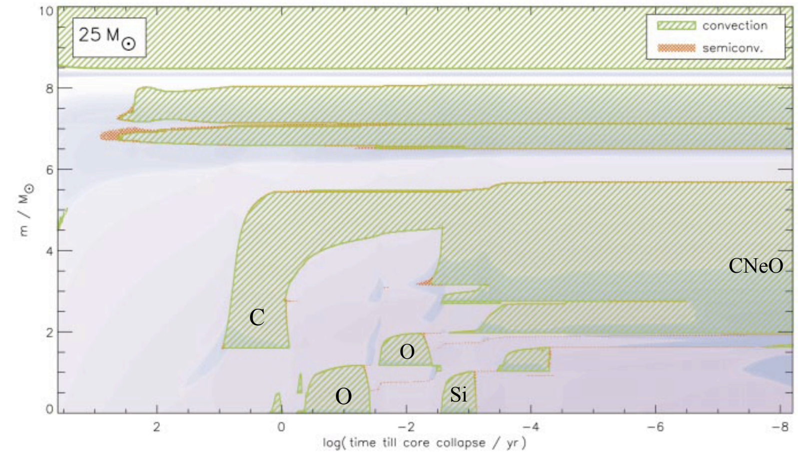
$T_{burn} \approx 0.80$

Energy is also provided by the Kelvin-Helmholz contraction of the core and this decreases the ignition temperature just a little. In more massive stars where $X(^{12}C)$ is less than about 10%, carbon burning and neon burning at the middle generate so little energy that the core never becomes convective. The carbon and neon just melt away without greatly exceeding the neutrino losses.

Further out in a shell, the burning temperature is higher (set by the gravitational potential at the bottom of the shell - similar energy generation has to come from less fuel set by the pressure scale height).

Convection





Carbon core burning not centrally convective
in more massive stars.

F. Lifetime

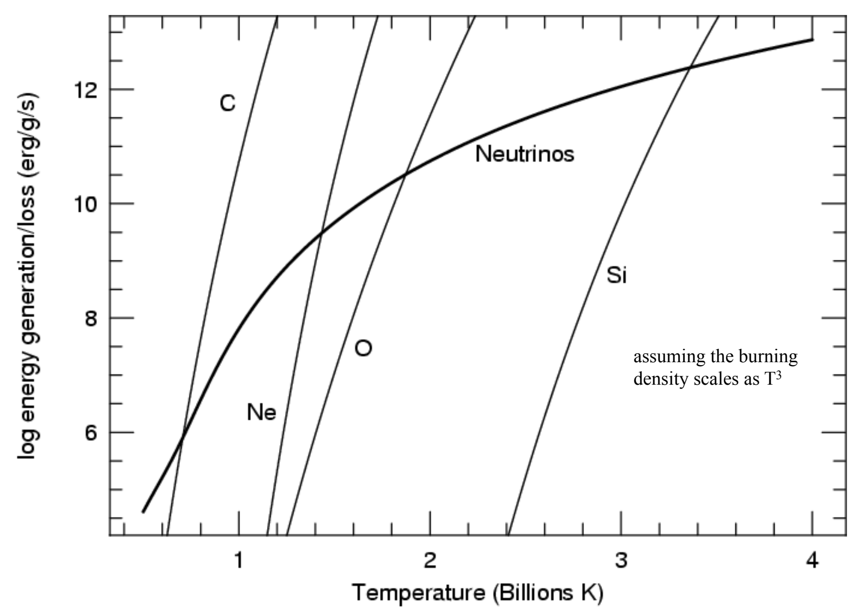
$$\tau_{12}(^{12}C) = (\rho Y_{12} \lambda_{12,12})^{-1} \qquad \text{vs} \left(\frac{1}{Y_{12}}\frac{dY_{12}}{dt}\right)^{-1}$$

@ $T_9 = 0.80$ $\quad \lambda_{12,12} = 7.9\times10^{-14}$ $\quad$ evaluate @ $Y_{12} = 0.1/12$

$$\tau_{12}(^{12}C) = \left[(2\times10^5)\left(\frac{0.1}{12}\right)(7.9\times10^{-14})\right]^{-1}$$

$$= 240 \text{ years}$$

Actual lifetime is lengthened by convection which brings fresh fuel to the center. The actual lifetime ranges from a few hundred to a few $\times10^4$ yr. More massive stars have the shorter lifetimes.



G. Freezing out of the Supergiant Envelope

One Relevant time scale is the Kelvin-Helmholtz time scale of the envelope (if L_{He} died envelope would take that long to contract; if $L\uparrow$ it would take that long to expand)

$$\tau_{KH} \sim \frac{G M_{core} M_{env}}{2 L_\gamma \; R/2}$$

eg. $15\ M_\odot$

$M_{core} = 4\ M_\odot$

$M_{env} = 11\ M_\odot$

$R = 3\times10^{13}$ cm

$L = 3\times10^{38}$ erg/s

$$\tau_{KH} \sim \frac{(6.67)(10^{-8})(10^{34})^2\,(.8)(2.2)}{(3\times10^{38})(3/2\times10^{13})}$$

≈ 100 years

actually should use a smaller radius here → a longer lifetime.

By late in carbon burning the envelope can no longer respond to changes in the core. The core evolves independently its evolution determined by ϵ_ν.

Also L_γ is provided, even in the presupernova star, by the helium burning shell. Its luminosity does not vary greatly during carbon burning.

15 solar mass star (KEPLER)

Stage	T_9	Radius	L_γ	L_ν
H-burn	0.03	4.36(11)	1.06(38)	7.0(36)
He-burn	0.18	3.21(13)	1.73(38)	7.4(36)
C-ign	0.50	4.76(13)	2.78(38)	7.1(37)
C-dep	1.2	5.64(13)	3.50(38)	3.5(41)
O-dep	2.2	5.65(13)	3.53(38)	3.8(43)
Si-dep	3.7	5.65(13)	3.53(38)	2.3(45)
PreSN	7.6	5.65(13)	3.53(38)	1.9(49)

Burning Stages in the Life of a Massive Star

hydrogen burning

$M_{\rm initial}$	T	ρ	M	L	R	τ
$M_\odot$	10^7 K	g cm^{-3}	$M_\odot$	$10^3\ L_\odot$	$R_\odot$	Myr
1	1.57	153	1.00	0.001	1.00	~1,100 0
15	3.53	5.81	14.9	28.0	6.75	11.1
20	3.69	4.53	19.7	62.6	8.03	8.13
25	3.81	3.81	24.5	110	9.17	6.70
75	4.26	1.99	67.3	916	21.3	3.16

helium burning

$M_{\rm initial}$	T	ρ	M	L	R	τ
$M_\odot$	10^8 K	10^3 g cm^{-3}	$M_\odot$	$10^3\ L_\odot$	$R_\odot$	Myr
1	1.25	20	0.71	0.044	~ 100	110
15	1.78	1.39	14.3	41.3	461	1.97
20	1.88	0.968	18.6	102	649	1.17
25	1.96	0.762	19.6	182	1,030	0.839
75	2.10	0.490	16.1	384	1.17	0.478

carbon burning

$M_{\rm initial}$	T	ρ	M	L	R	τ
$M_\odot$	10^8 K	10^5 g cm^{-3}	$M_\odot$	$10^3\ L_\odot$	$R_\odot$	kyr
15	8.34	2.39	12.6	83.3	803	2.03
20	8.70	1.70	14.7	143	1,070	0.976
25	8.41	1.29	12.5	245	1,390	0.522
75	8.68	1.39	6.37	164	0.644	1.07

neon burning

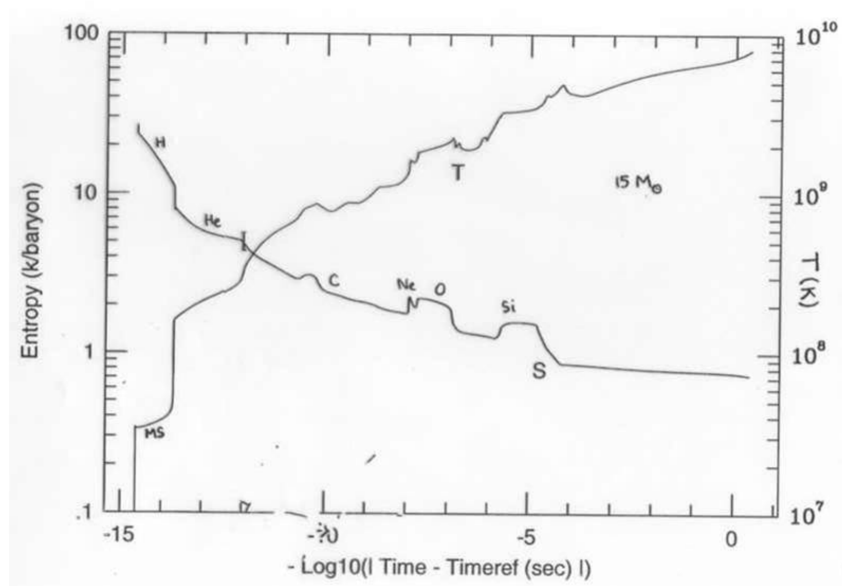
$M_{\rm initial}$	T	ρ	M	L	R	τ
$M_\odot$	10^9 K	10^6 g cm^{-3}	$M_\odot$	$10^3\ L_\odot$	$R_\odot$	yr
15	1.63	7.24	12.6	86.5	821	0.732
20	1.57	3.10	14.7	147	1,090	0.599
25	1.57	3.95	12.5	246	1,400	0.891
75	1.62	5.21	6.36	167	0.715	0.569

oxygen burning

$M_{\rm initial}$	T	ρ	M	L	R	τ
$M_\odot$	10^9 K	10^6 g cm^{-3}	$M_\odot$	$10^3\ L_\odot$	$R_\odot$	yr
15	1.94	6.66	12.6	86.6	821	2.58
20	1.98	5.55	14.7	147	1,090	1.25
25	2.09	3.60	12.5	246	1,400	0.402
75	2.04	4.70	6.36	172	0.756	0.908

silicon burning

$M_{\rm initial}$	T	ρ	M	L	R	τ
$M_\odot$	10^9 K	10^7 g cm^{-3}	$M_\odot$	$10^3\ L_\odot$	$R_\odot$	d
15	3.34	4.26	12.6	86.5	821	18.3
20	3.34	4.26	14.7	147	1,090	11.5
25	3.65	3.01	12.5	246	1,400	0.733
75	3.55	3.73	6.36	173	0.755	2.09



Neon Burning

Following carbon burning, at a temperature of about 1.5 x 10^9 K, neon is the next abundant nucleus to burn. It does so in a novel "photodisintegration rearrangement" reaction which basically leads to oxygen and magnesium (nb. not ^{20}Ne + ^{20}Ne burns to ^{40}Ca)

$$2(^{20}\text{Ne}) \rightarrow {}^{16}\text{O} + {}^{24}\text{Mg} + \text{energy}$$

The energy yield is not large, but is generally sufficient to power a brief period of convection. It was overlooked early on as a separate burning stage, but nowadays is acknowledged as such.

The nucleosynthetic products resemble those of carbon burning but lack ^{23}Na and have more of the heavier nuclei, $^{(26),27}$Al, ^{29,30}Si, and ^{31}P.

$$\alpha + {}^{16}O \underset{\text{fast}}{\rightleftarrows} {}^{20}Ne + \gamma \qquad {}^{20}Ne + \alpha \xrightarrow{\text{slow}} {}^{24}Mg + \gamma$$

The net result is that $2\left(^{20}\text{Ne}\right) \rightarrow {}^{16}\text{O} + {}^{24}\text{Mg}$ at a rate that is determined by how fast ^{20}Ne captures alpha particles from the equilibrium concentration set up by ^{16}O and ^{20}Ne.

Other secondary reactions:

^{24}Mg(α,γ)^{28}Si	^{27}Al(α,p)^{30}Si
^{25}Mg(α.n)^{29}Si	^{30}Si(p,γ)^{31}P
^{26}Mg(α,n)^{30}Si	etc.

Products:

some more ^{16}O and ^{24}Mg, ^{29,30}Si, ^{31}P, ^{26}Al and a small amount of s-process.

A. Basics: The composition following carbon burning is chiefly

^{16}O, ^{20}Ne, ^{24}Mg

but ^{16}O is not the next to burn (influence of Z = N = 8 = magic)

Species	S_α(MeV)
^{16}O	7.16
^{20}Ne	4.73
^{24}Mg	9.32

energy required to remove an α-particle.

Before the temperature becomes hot enough for oxygen to fuse (T_9 = 1.8 as we shall see), photons on the high energy tail of the Bose-Einstein distribution function begin to induce a new kind of reaction -

^{20}Ne(γ,α)^{16}O

The α-particle "photo-disintegrated" out of ^{20}Ne usually just adds back onto ^{16}O creating an "equilibrated link" between ^{16}O and ^{20}Ne. Sometimes though an α captures on ^{20}Ne to make ^{24}Mg. When this happens the equilibrium between ^{16}O and ^{20}Ne quickly restores the α that was lost.

B. Photodisintegration Reaction Rates

At high temperatures, the inverse reaction to radiative capture, [(n,γ),(p,γ),(α,γ)] becomes important as there exists an appreciable abundance of γ-rays out on the tail of the Bose-Einstein distribution that have energy in excess of several MeV. The reactions these energetic photons induce are called *photodisintegration* reactions – the major examples being (γ,n),(γ,p), and (γ,α)

Consider

$$I + j \rightarrow L + \gamma$$

and

$$L + \gamma \rightarrow I + j$$

In equilibrium, the abundances must obey the Saha equation

For the reaction $I + j \rightleftharpoons L + \gamma$

$$\frac{n_I n_j}{n_L} = \left(\frac{g_I g_j}{g_L}\right)\left(\frac{A_I A_j}{A_L}\right)^{3/2}\left(\frac{2\pi kT}{h^2 N_A}\right)^{3/2} \exp(-Q_{j\gamma}/kT)$$

(deriveable from considerations of entropy and the chemical potential and the fact that the chemical potential of the photon is zero). Thus, in *equilibrium* (a more stringent condition than "steady state")

$$\boxed{\left(\frac{n_I n_j}{n_L}\right) = 5.942\times 10^{33}\, T_9^{3/2}\left(\frac{g_I g_j}{g_L}\right)\left(\frac{A_I A_j}{A_L}\right)^{3/2} \exp(-11.60485 Q_{j\gamma}/T_9)}$$

for Q_{ij} measured in MeV

Equilibrium in the reaction $I + j \rightleftharpoons L + \gamma$ also implies

$$Y_I Y_j\, \rho \lambda_{j\gamma}(I) = Y_L\, \lambda_{\gamma j}(L)$$

and since $Y_i = \dfrac{n_i}{\rho N_A}$

$$\frac{n_I n_j}{n_L} = \frac{\rho N_A Y_I Y_j}{Y_L} = \frac{\lambda_{\gamma j}(L) N_A}{\lambda_{j\gamma}(I)} = 5.942\times 10^{33}\, T_9^{3/2}\left(\frac{g_I g_j}{g_L}\right)\left(\frac{A_I A_j}{A_L}\right)^{3/2} e^{-11.6048\, Q_{j\gamma}/T_9}$$

$$\boxed{\lambda_{\gamma j}(L) = \lambda_{j\gamma}(I) \bullet 9.868\times 10^{9}\, T_9^{3/2}\left(\frac{g_I g_j}{g_L}\right)\left(\frac{A_I A_j}{A_L}\right)^{3/2} \exp(-11.6048 Q_{j\gamma}/T_9)}$$

where $9.686\times 10^{9} = \dfrac{5.942\times 10^{33}}{N_A}$

c. <u>Energy generation during neon burning</u>

The net process $2(^{20}Ne) \rightarrow {}^{24}Mg + {}^{16}O$ releases energy. It takes 4.73 MeV to remove the α from ^{20}Ne but one gets 9.32 MeV back by adding it to another ^{20}Ne. However it must be very hot to do this. The radiation bath drives the first reaction.

Quickly $^{16}O + \alpha \rightleftharpoons {}^{20}Ne + \gamma$

<u>i.e.</u> $Y(^{16}O)\, Y_\alpha\, \rho\, \lambda_{\alpha\gamma}(^{16}O) \approx Y(^{20}Ne)\, \lambda_{\gamma\alpha}(^{20}Ne)$

$$\Rightarrow Y_\alpha = \frac{Y(^{20}Ne)\, \lambda_{\gamma\alpha}(^{20}Ne)}{Y(^{16}O)\, \rho\, \lambda_{\alpha\gamma}(^{16}O)}$$

And since

$$\frac{dY(^{20}Ne)}{dt} = -2\, Y(^{20}Ne)\, Y_\alpha\, \rho\, \lambda_{\alpha\gamma}(^{20}Ne), = -2f$$

and $\dfrac{dY(^{24}Mg)}{dt} = +f \qquad \dfrac{dY(^{16}O)}{dt} = +f$

and $\lambda_{\gamma\alpha}(^{20}Ne) = 9.87\times 10^{9}\, T_9^{3/2} \left(\frac{16\cdot 4}{20}\right)^{3/2} \lambda_{\alpha\gamma}(^{16}O)\, e^{-(11.605)(4.73)/T_9}$

$$Y_\alpha = 5.65\times 10^{10}\, T_9^{3/2}\, \frac{Y(^{20}Ne)}{\rho\, Y(^{16}O)}\, e^{-54.89/T_9}$$ (previous page)

very T sensitive

$$\epsilon_{nuc} = 9.65\times 10^{17} \left[\tfrac{1}{2}(\underbrace{198.258}_{^{24}Mg}) + \tfrac{1}{2}(\underbrace{127.62}_{^{16}O}) - \underbrace{160.646}_{^{20}Ne}\right] \cdot 2f$$

$$\epsilon_{nuc} = (9.65\times10^{17})(2.29)(2f) = 4.41\times10^{18} f \text{ erg g}^{-1}\text{s}^{-1}$$

$$\boxed{\epsilon_{nuc}^{Ne} \approx 2.49\times10^{29}\, T_9^{3/2}\, \frac{Y^2(^{20}\mathrm{Ne})}{Y(^{16}\mathrm{O})}\, \lambda_{\alpha\gamma}(^{20}\mathrm{Ne})\, e^{-54.89/T_9} \text{ erg g}^{-1}\text{s}^{-1}}$$

$$q_{nuc}^{Ne} = 9.65\times10^{17}\, \frac{2.29}{20}\, \Delta X(^{20}\mathrm{Ne})$$

$$\boxed{q_{nuc}^{Ne} = 1.1\times10^{17}\, \Delta X(^{20}\mathrm{Ne}) \text{ erg g}^{-1}}$$

relatively small

Note that the energy generation rate is independent of the density.

D. Balanced Power Condition:

$X(^{20}\mathrm{Ne}) \approx 0.2 \quad X(^{16}\mathrm{O}) \approx 0.7 \quad \rho \sim 10^6 - 10^7$

near $T_9 = 1.5 \quad \lambda_{\alpha\gamma}(^{20}\mathrm{Ne}) = 3.43\times10^{-3}\, T_9^{10.5}$

$$\therefore\ \epsilon_{nuc} \propto T_9^{1.5}\, T_9^{10.5}\, e^{-54.89/T_9}$$

$$\propto T_9^{49} \quad @ \quad T_9 = 1.5$$

very T-sensitive

Above eq^n + conditions give

$$\epsilon_{nuc}^{Ne} \approx 3.3\times10^{10} \left(\frac{T_9}{1.5}\right)^{49} \text{ erg g}^{-1}\text{s}^{-1}.$$

T_9	ϵ_{nuc}	ϵ_ν ($\rho = 10^6$)	ϵ_ν ($\rho = 10^7$)
1.4	1.1 (9)	8.9 (9)	2.9 (8)
1.5	3.3 (10)	2.2 (10)	7.8 (8)
1.6	7.8 (11)	4.9 (10)	2.0 (9)

$$\dot{\epsilon}_{nuc} = \dot{\epsilon}_\nu \left(\frac{[3u+s]^{nuc}}{[3u+s]^{\nu}}\right)^{3/2} \qquad \epsilon \propto \rho^{u-1} T^s$$

$$= \dot{\epsilon}_\nu \left(\frac{3+49}{13}\right)^{3/2} \approx 8.5\, \dot{\epsilon}_\nu$$

So neon burns at about <u>$T_9 = 1.6$</u> to <u>1.5</u> (somewhat higher in a shell where the density is less)

E. <u>Lifetime</u>

At a temperature between $T_9 = 1.5$ and 1.6

$$\tau \sim q_{nuc}/\varepsilon_{nuc} \approx \frac{1.1\times10^{17} \bullet 0.2}{10^{11}} = 2\times10^5 \text{ s}$$

This is greatly lengthened by convection – if it occurs – because of high T sensitivity. Typically ~ few years.

Nucleosynthesis from neon burning

The principal nuclei with major abundances at the end of neon burning are ^{16}O and ^{24}Mg. Most of the neutron excess resides in $^{25,26}Mg$. Most of the ^{16}O has in fact survived even since helium burning.

In terms of major production of solar material, important contributions are made to

$[^{16}O]$, $^{24,25,26}Mg$, $^{(26),27}Al$, $^{29,30}Si$, and ^{31}P

Oxygen Burning:

After neon burning the lightest nucleus remaining with appreciable abundance is ^{16}O. This not only has the lowest Coulomb barrier but because of its double magic nature, has a high α-particle separation energy. It is the next to burn.

Because of its large abundance and the fact that it is a true fusion reaction, not just a rearrangement of light nuclei, oxygen burning releases a lot of energy and is a very important part of the late stages of stellar evolution in several contexts (e.g., pair-instability supernovae).

It is also very productive nucleosynthetically. It's chief products being most of the isotopes from ^{28}Si to ^{40}Ca as well as (part of) the *p*-process.

Initial composition:

^{16}O, ^{24}Mg, ^{28}Si

Nuclear reactions:

$$
\begin{aligned}
{}^{16}O + {}^{16}O \rightarrow ({}^{32}S)^* &\rightarrow {}^{31}S + n + 1.45\,\text{MeV} && 5\% \\
&\rightarrow {}^{30}P + d - 2.41\,\text{MeV} && \leq 5\% \\
&\rightarrow {}^{31}P + p + 7.68\,\text{MeV} && 56\% \\
&\rightarrow {}^{28}Si + \alpha + 9.59\,\text{MeV} && 34\%
\end{aligned}
$$

proceeds through the ^{32}S compound nucleus with a high density of resonances. Very like carbon burning.

The deuteron, *d*, is quickly photodisintegrated into a free neutron and proton.

<u>Secondary</u> very, very many

1) $^{28}Si(\alpha,\gamma)^{32}S$ $^{29}Si(\alpha,n)^{32}S$ $^{30}Si(\alpha,\gamma)^{34}S$

$^{32}S(\alpha,\gamma)^{36}Ar$ $^{36}Ar(\alpha,\gamma)^{40}Ca$ $^{36}Ar(\alpha,p)^{39}K$ $^{32}S(\alpha,p)^{35}Cl$

$^{31}P(p,\gamma)^{32}S$ $^{35}Cl(p,\gamma)^{36}Ar$ $^{34}S(\alpha,\gamma)^{38}Ar$ $^{38}Ar(p,\gamma)^{39}K$

<u>etc.</u>, <u>etc</u>

Many isotopes begin to be made as radioactive progenitors. For example, ^{37}Cl made as ^{37}Ar, ^{41}K made as ^{41}Ca.

2) The heavy isotopes ($A \gg$ iron group) initially present in the star as well as the s-process made during helium and carbon burning begins to be stripped down by photodisintegration reactions

$$^{208}Pb(\gamma,n)^{207}Pb(\gamma,n)^{206}Pb \ldots \; ^{194}Pb(\gamma,\alpha)^{190}Hg \ldots$$

The process continues, at least during oxygen core burning until the heavy isotopes have "melted" into the Fe-group. Flows may be of some relevance to production of the p-process nuclei.

3) Onset of "quasi-equilibrium" clusters

e.g. $\quad {}^{28}Si + n \rightleftarrows {}^{29}Si + \gamma \qquad {}^{29}Si + p \rightleftarrows {}^{30}P + \gamma \qquad$ etc.

These clusters apear and grow as oxygen burning proceeds
(Woosley, Arnett, & Clayton, ApJS, 26, 271 (1973))

4) Weak interactions increase η markedly during oxygen core burning (much less so during oxygen shell burning where the density is less and the time scale shorter).

$$^{33}S(e^-,\nu_e)^{33}P \qquad ^{37}Ar(e^-,\nu_e)^{37}Cl$$

$$^{35}Cl(e^-,\nu_e)\,^{35}S$$

<u>Nuclear</u> <u>Energy</u> <u>Generation</u>

Approximation

$$^{16}O + {}^{16}O \rightarrow {}^{32}S + 16.54 \text{ MeV}$$

(actually ^{28}Si, ^{32}S, ^{36}Ar, ^{40}Ca in rough proportions 10:5:1:1)

$$\frac{dY(^{16}O)}{dt} = -2Y^2(^{16}O)\,\rho\,\lambda_{16,16}/2$$

$$\frac{dY(^{32}S)}{dt} = -\tfrac{1}{2}\,\frac{dY(^{16}O)}{dt}$$

$$\epsilon_{nuc} = 9.65\times10^{17}\; Y^2(^{16}O)\,\rho\,\lambda_{16,16}\;\underbrace{\big[\tfrac{1}{2}(\underbrace{271.775}_{^{32}S}) - \underbrace{127.617}_{^{16}O}\big]}_{8.271 \text{ MeV}}$$

$$\boxed{\dot{\epsilon}_{nuc} \approx 8.0\times10^{18}\; Y^2(^{16}O)\,\rho\,\lambda_{16,16} \;\text{erg g}^{-1}\text{s}^{-1}}$$

$$\lambda_{16,16} \propto T^n \qquad n = \frac{\tau-2}{3} = \frac{108-2}{3} = 35$$

$$\tau = 4.248\left(\frac{8^2 8^2 \frac{16\cdot16}{32}}{T_9}\right)^{1/3}$$

$$\lambda_{16,16} \approx 7.8\times10^{-12}\left(\frac{T_9}{2}\right)^{35}$$

For $X(^{16}O) \sim 0.5$

$$\boxed{\dot{\epsilon}_{nuc} \approx 6\times10^{4}\,\rho\left(\frac{T_9}{2}\right)^{35} \text{erg g}^{-1}\text{s}^{-1}}$$

$$\dot{q}_{nuc} = 9.65\times10^{17}\left(\frac{8.271}{16}\right)\Delta X_{16}$$

$$\boxed{\dot{q}_{nuc} = 5.0\times10^{17}\,\Delta X_{16} \text{ erg g}^{-1}}$$

substantial ΔX_{16} is large

$\lambda_{16,16}$	T_9
2.60E-19	1.250
3.04E-16	1.500
8.15E-14	1.750
7.81E-12	2.000
9.09E-09	2.500
1.73E-06	3.000
1.01E-04	3.500
2.61E-03	4.000
3.33E-01	5.000
9.93E+00	6.000
1.15E+02	7.000
6.[illegible]2E+02	8.000
2.[illegible]0E+03	9.000
6.30E+03	10.000

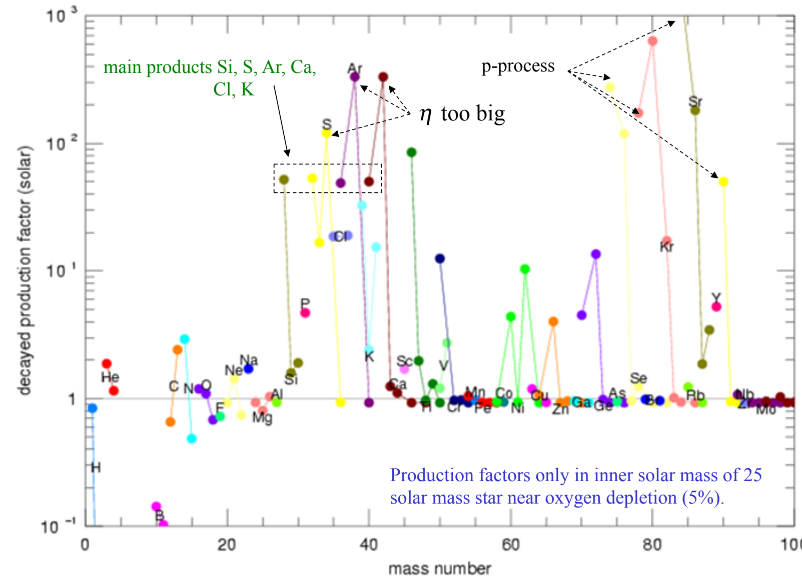
Nucleosynthesis

^{28}Si, ^{32,33,34}S, ^{35,37}Cl, ^{36,38}Ar, ^{39,41}K, ^{40,42}Ca, some p-process

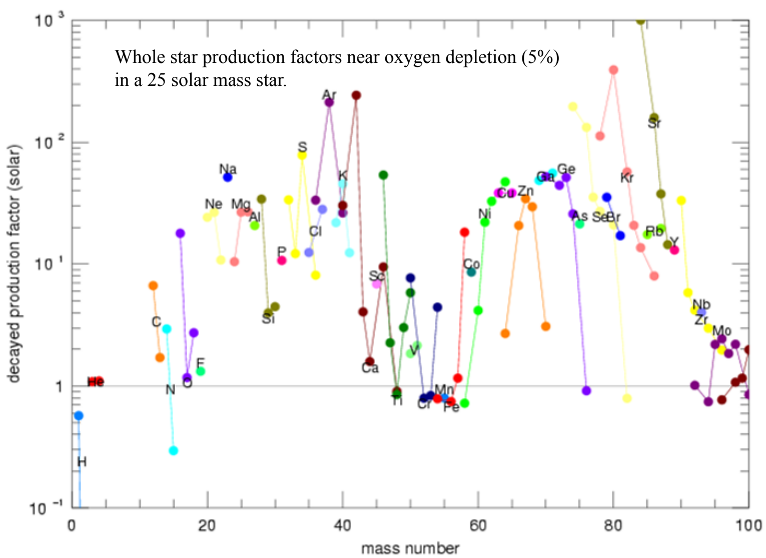
Element-wise: Si, S, Ar, Ca in roughly solar proportions.

Destruction of the s-process

Increasing neutronization, especially right after oxygen disappears from the center.



s25a28



s25a28

Simplicius (6th century AD commenting on work of Leucippus (5th century BC)

They [the atoms] move in the void and catching each other up and jostle together, and some recoil in any direction that may chance, and others become entangled with one another in various degrees according to the symmetry of their shapes and sizes and positions and order, and they remain together and thus the coming into being of composite things is effected.

But is it possible to admit that such a transmutation is occurring? It is difficult to assert, but perhaps more difficult to deny, that this is going on. Sir Ernest Rutherford has recently been breaking down the atoms of oxygen and nitrogen, driving out an isotope of helium from them; and what is possible in the Cavendish laboratory may not be too difficult in the Sun. I think that the suspicion has been generally entertained that the stars are the crucibles in which the lighter atoms which abound in the nebulæ are compounded into more complex elements. In the stars matter has its preliminary brewing to prepare the greater variety of elements which are needed for a world of life. The radio-active elements must have been formed at no very distant date; and their synthesis, unlike the generation of helium from hydrogen, is endothermic. If combinations requiring the addition of energy can occur in the stars, combinations which liberate energy ought not to be impossible.

We need not bind ourselves to the formation of helium from hydrogen as the sole reaction which supplies the energy, although it would seem that the further stages in building up the elements involve much less liberation, and sometimes even absorption, of energy. It is a question of accurate measurement of the deviations of atomic weights from integers, and up to the present

Sir Arthur Eddington *The Internal Constitution of Stars (The Observatory*, Vol. 43, p. 341-358 (1920)) p 354.

NEW GENESIS

In the beginning God created radiation and ylem. And ylem was without shape or number, and the nucleons were rushing madly over the face of the deep.

And God said: "Let there be mass two." And there was mass two. And God saw deuterium, and it was good.

And God said: "Let there be mass three." And there was mass three. And God saw tritium and tralphium, and they were good. And God continued to call number after number until He came to transuranium elements. But when He looked back on his work He found that it was not good. In the excitement of counting, He missed calling for mass five and so, naturally, no heavier elements could have been formed.

God was very much disappointed, and wanted first to contract the Universe again, and to start all over from the beginning. But it would be much too simple. Thus, being almighty, God decided to correct His mistake in a most impossible way.

And God said: "Let there be Hoyle." And there was Hoyle. And God looked at Hoyle . . . and told him to make heavy elements in any way he pleased.

And Hoyle decided to make heavy elements in stars, and to spread them around by supernovae explosions. But in doing so he had to obtain the same abundance curve which would have resulted from nucleosynthesis in ylem, if God would not have forgotten to call for mass five.

And so, with the help of God, Hoyle made heavy elements in this way, but it was so complicated that nowadays neither Hoyle, nor God, nor anybody else can figure out exactly how it was done.

Amen.

My attitude toward the steady-state theory, expressed in this piece, may account for my not receiving an invitation to the 1958 Solvay Congress on cosmology.

George Gamow in *My World Line*

Solving the wave equation in a plasma

$m_e \ddot{\mathbf{r}} = e\mathbf{E}$ $\qquad \mathbf{J} \equiv n_e\, e\, \dot{\mathbf{r}}$

$$\frac{\partial \mathbf{J}}{\partial t} = \frac{\partial}{\partial t}(n_e e \dot{\mathbf{r}}) = n_e e \ddot{\mathbf{r}} = n_e e \left(\frac{e\mathbf{E}}{m_e} \right)$$

combine a plane wave $\mathbf{E} = E_o \exp(i(kx - \omega t))\hat{\mathbf{x}}$

which satisfies the wave equation

$\nabla^2 \mathbf{E} = -k^2 E_0 e^{i(kx-\omega t)} \hat{\mathbf{x}}$ with

$$\nabla^2 \mathbf{E} - \frac{1}{c^2}\frac{\partial^2 \mathbf{E}}{\partial t^2} = \frac{4\pi}{c^2}\frac{\partial \mathbf{J}}{\partial t}$$

(e.g., http://scienceworld.wolfram.com/physics/PlasmaFrequency.html)

gives

$$v_{phase}^2 = \frac{\omega^2}{k^2} = c^2 \left(1 - \frac{4\pi n_e e^2}{m_e \omega^2} \right)^{-1}$$

Undefined if $\omega < \omega_p$